

Two years of pulsar observations with the ultra-wide-band receiver on the Parkes radio telescope

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ABSTRACT

The major programme for observing young, non-recycled pulsars with the Parkes telescope has transitioned from a narrow-band system to an ultra-wide-band system capable of observing between 704 and 4032 MHz. We report here on the initial 2 yr of observations with this receiver. Results include dispersion measure (DM) and Faraday rotation measure (RM) variability with time, determined with higher precision than hitherto, flux density measurements and the discovery of several nulling and mode changing pulsars. PSR J1703–4851 is shown to be one of a small subclass of pulsars that has a weak and a strong mode which alternate rapidly in time. PSR J1114–6100 has the fourth highest $|RM|$ of any known pulsar despite its location far from the Galactic Centre. PSR J1825–1446 shows variations in both DM and RM likely due to its motion behind a foreground supernova remnant.

Key words: pulsars: general.

1 INTRODUCTION

Results from the long-term monitoring of millisecond pulsars have the potential to do fundamental science, such as tests of theories of gravity (Kramer et al. 2006b; Archibald et al. 2018), determining the equation of state of matter (Antoniadis et al. 2013), and the detection of gravitational waves (Shannon et al. 2015; Aggarwal et al. 2019). Such experiments have limitations based on our incomplete knowledge of the pulsar emission mechanism, the pulsar spin-down and on the Galactic magneto-ionized medium through which the radio waves propagate. Therefore, observations over decades of the bulk of the population of radio pulsars are warranted as they allow science to be gleaned both from the behaviour of the pulsar itself and from the properties of the interstellar medium (ISM). As examples of the former, carried out with the Parkes telescope, Yu et al. (2013) compiled glitch statistics, Parthasarathy et al. (2019, 2020) investigated the timing noise properties of pulsars over more than a decade and Johnston & Kerr (2018) investigated the polarization properties of 600 pulsars. Meanwhile, changes in pulse profiles with time (Lyne et al. 2010) and the implications thereof were discussed in Brook et al. (2016) and Kerr et al. (2016). The properties of the ISM were explored through long-term changes in the dispersion measure (Petroff et al. 2013), extreme scattering events (Kerr et al. 2018), and

variations in flux density (Kumamoto et al. 2021) for a large sample of pulsars.

The Parkes radio telescope has been observing and timing pulsars over the past 30 yr at a variety of observing frequencies and employing signal processing techniques with increasingly better capabilities. In the 1990s and early 2000s, analogue filterbanks operating at an observing frequency of 1400 MHz were used with timing programmes described in e.g. Johnston et al. (1995) and Wang et al. (2000). Subsequently, digital filterbanks with better frequency resolution and full polarization capability were routinely employed. In 2007, a major timing programme of young, energetic pulsars was set up (Weltevrede et al. 2010a) to provide ephemerides for NASA’s *Fermi* Gamma-Ray Space Telescope mission (Smith et al. 2008). The programme used a central observing frequency of 1369 MHz and a bandwidth of 256 MHz to monitor a sample of some 150 pulsars on a monthly basis with twice-yearly observations at 3.1 and 0.7 GHz. The programme achieved its goal of increasing the sample of γ -ray pulsars more than 10-fold (Weltevrede et al. 2010b; Abdo et al. 2013; Smith et al. 2019). In late 2018, the ultra-wide-band receiver (UWL) was installed on the Parkes telescope. The UWL allows observing over the entire band between 704 and 4032 MHz. The FPGA-based digital filterbanks were replaced by a GPU-based software system known as Medusa that enables fully flexible backend configurations to be used. A comprehensive description of the capabilities of the UWL and Medusa can be found in Hobbs et al. (2020).

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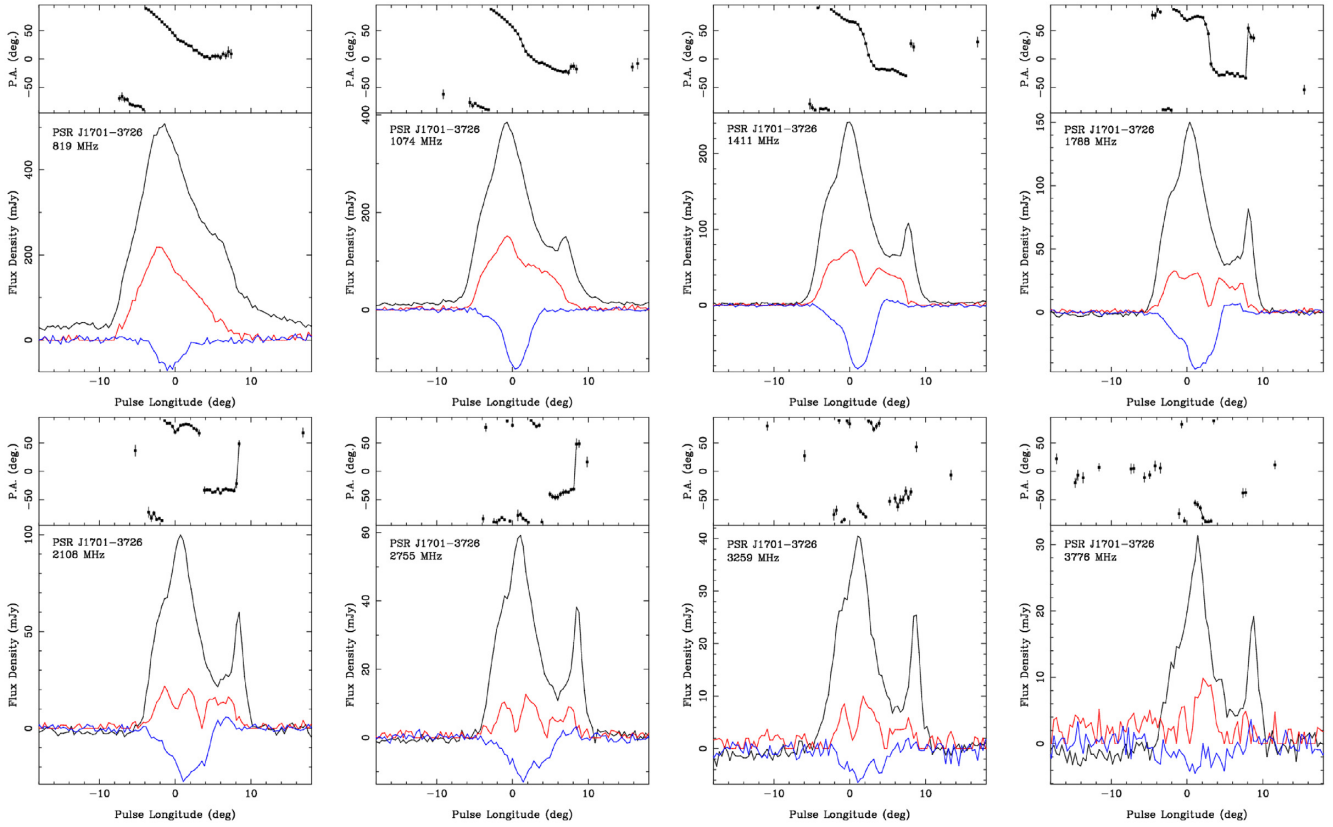


Figure 1. Polarization profiles for PSR J1701–3726 at eight different frequencies across the UWL band. Black line denotes total intensity, red and blue lines are the linear polarization and circular polarization. The position angle of the linear polarization as a function of pulse phase is also shown.

In this paper, we present results from the first 24 months of young pulsar observations with the UWL, mainly concentrating on aspects of the time variability of the sample. A polarization study of the brightest pulsars is presented in a companion paper (Sobey et al. 2021). Section 2 describes the observations. In Section 3, we look at profile variations with time, Section 4 looks at flux density variability, with Sections 5 and 6 examining dispersion measure (DM) and rotation measure (RM) changes, respectively.

2 OBSERVATIONS AND DATA REDUCTION

The initial impetus for this observation programme was the launch of the *Fermi* satellite in 2008. The programme started observing some 150 pulsars with high spin-down energy loss rates, \dot{E} , derived from the lists given in Smith et al. (2008). In 2014, the programme changed emphasis. Observations of some of the weaker, high \dot{E} pulsars were discontinued as it was evident that they were not γ -ray emitters. A substantial number of bright, lower \dot{E} pulsars were added when it was realized that these too could be γ -ray pulsars (Smith et al. 2019). Since 2014 therefore, 276 pulsars are observed in each monthly session in a single block of duration 21 h. The list of pulsars is given in Table A1.

Observations with the UWL commenced in 2018 November and have occurred roughly monthly thereafter for a total of 23 sessions up to 2020 October. The band between 704 and 4032 MHz is subdivided into 3328 frequency channels each with a channel bandwidth of 1 MHz. Data for each channel are coherently dedispersed at the DM of the pulsar and folded at the topocentric spin period to form a pulse profile with 1024 bins across the pulsar period. Data are integrated

over 30 s and written to disc. The typical observation length for each pulsar is 180 s with a handful of weaker pulsars observed for up to 480 s. Every 60 min, observations are made of a pulsed, square-wave, calibration signal to allow for polarization calibration. Flux calibration is carried out via observations of the source Hydra A. A complete end-to-end description of the system can be found in Hobbs et al. (2020).

The data are all processed using PSRCHIVE (Hotan, van Straten & Manchester 2004) and data reduction proceeds as follows. First, the calibration observations are examined and radio frequency interference (RFI) signals are removed. The pulsar observation is then calibrated. The square-wave calibration signal provides polarization calibration through correction of the gains and phases. The observation of the flux calibrator converts the digitizer units to mJy. Finally, the technique described in van Straten (2013) is used to account for the instrumental leakage terms. A two-step process is then used to remove RFI from the pulsar data. First, the data are summed in time and the RFI flagged in the frequency channels using a median-filter technique and then the individual time-steps are examined and bad time integrations are flagged before once again the data are summed in time.

As an example of the output of the data processing, the polarization profiles of PSR J1701–3726 at eight different frequencies across the band of the UWL are shown in Fig. 1.

3 PULSE PROFILE VARIABILITY

For an individual pulsar, each rotation provides a snap-shot pulse profile and these can differ significantly from pulse to pulse. Once

sufficient pulses are averaged, a characteristic pulse profile emerges that (at least to zeroth order) appears to be stable over decades. It is this stability that makes high-precision timing experiments possible. However, even in the early days of pulsar research, a phenomenon loosely known as ‘mode changing’ was observed in some pulsars (Backer 1970b). In these cases, a given pulsar appears to have two or more different, stable pulse profiles. Often, one of the modes is the off or nulling state (e.g. Backer 1970a; Biggs 1992), while in some cases a pulsar may have a ‘bright’ mode and a ‘quiet’ mode (e.g. Hermsen et al. 2013) and in yet other cases have modes of different widths (Lyne et al. 2010; Brook et al. 2016). The time-scale for switching between modes can be as short as one rotation period (Sobey et al. 2015) and as long as years (Kramer et al. 2006a), but what sets the time-scale has not yet been determined. Furthermore, without sufficient sensitivity to single pulses, it is unclear whether all pulsars exhibit mode changing or whether this is restricted to a particular class of pulsar. If we ignore the effects of mode changing, then pulsar profiles tend to stabilize to within a few per cent after several hundred rotations, with a small minority taking much longer to reach a stable profile.

In this particular set of observations, we are limited by the fact that we produce profiles every 30 s and so short-term mode changing is difficult to detect. We also only observe once per month, making it difficult to ascertain the time-scale for mode changing. In spite of these caveats, we can measure the flux density for every 30 s integration and inspect the time series for variability. We describe below those pulsars that show nulling and/or other time variable behaviour.

PSR J0729–1836: If sufficient rotations are obtained over a long integration time, the pulse profile has a strong leading component followed by a plateau before the trailing drop-off. There is little evolution with frequency. The profiles obtained in individual 180 s observations (350 rotations) can look very different due to the strong and independent modulation of the leading and trailing components.

PSR J0820–4114: This pulsar has a low \dot{E} of $5.4 \times 10^{30} \text{ erg s}^{-1}$ and a very wide profile that spans more than 100° of longitude, and it therefore appears to be an old, almost aligned rotator. It often appears to null for about 30 s and remain on for the rest of the 3 min observation so its nulling fraction is less than 15 per cent which is low compared to other aligned rotators.

PSR J1048–5832: The profile of this pulsar has at least four components (Sobey et al. 2021) and even though we average together 1450 rotations per 3-min observation, the profile can look very different from one observation to the next. We surmise that each component has very different amplitude statistics with occasional bright pulses so that different components dominate depending on the observation length. Single pulse studies are warranted.

PSR J1049–5833: This is a long-period (2.2 s) pulsar with a simple, Gaussian-like profile. Its nulls typically last for one-third of the observation.

PSR J1114–6100: This pulsar shows short-duration nulls, with a null fraction less than 20 per cent.

PSR J1428–5530: The profile of the pulsar has two components that are blended and the profile does not stabilize within the 180 s observation either in total intensity or in the fraction of linear polarization. The pulsar sometimes nulls within the 30 s subintegration time and has a low nulling fraction overall.

PSR J1646–6831: This pulsar has a spin period of 1.8 s and a low \dot{E} of $1.2 \times 10^{31} \text{ erg s}^{-1}$. The profile appears to show a classic double-conal structure although the sweep of position angle (PA) is highly distorted. The circular polarization exceeds the linear polarization

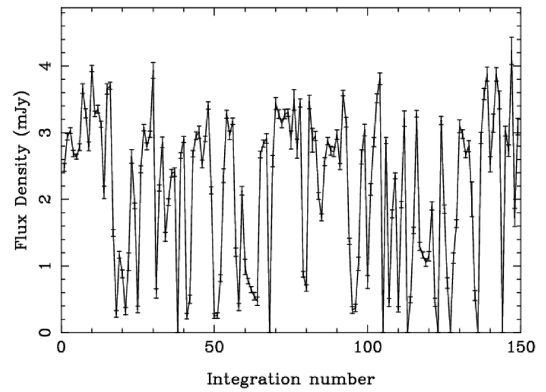


Figure 2. The flux density of PSR J1701–3726 in each 30 s integration. The integrations dominated by nulling are clearly seen.

fraction in the centre of the profile. The pulsar has short-duration nulls of which several can be seen per observation.

PSR J1701–3726: This is a long-period (2.5 s) pulsar with a complex profile that varies with frequency (see Fig. 1). It nulls within the 30 s integration and is in the null state approximately 25 per cent of the time. As a result of the nulling, the profile takes a long time to stabilize. Fig. 2 shows the flux density in each 30 s integration where the nulling is clearly seen.

PSR J1703–4851: This pulsar was flagged in Kumamoto et al. (2021) as showing unusual behaviour in its flux density variations that were clearly intrinsic to the pulsar rather than a propagation effect. In our sample, it also stands out as an exception. The pulsar shows two distinct modes. The left-hand panel of Fig. 3 shows the profile of the pulsar in the ‘bright’ mode with a strong central component and weak flanking components. The linear polarization is low and the circular polarization changes sign under the main component. In the ‘weak’ mode (right-hand panel of Fig. 3), the central component is reduced in flux density by about a factor of 10 compared to the bright mode and the flanking components stay roughly constant. Of the 14 observations, we have made of this pulsar, 5 are in the bright mode and 9 in the weak mode. In one observation, made on 2019 July 6, the pulsar changes from the bright mode to the weak mode in the middle of the observation indicating a fast switching time and a relatively short duration for each mode. This pulsar merits closer attention, a single-pulse analysis would prove beneficial.

PSR J1709–1640: The low DM of this pulsar means that it exhibits diffractive scintillation, but the time-scale of the scintillation is longer than the observation duration. The pulsar has short duration nulls, but is occasionally in the null state during the entire 3 min observation. We estimate the nulling fraction to be 10 per cent. Both Naidu et al. (2018) and Wang et al. (2020) obtain similar results, but also see occasional nulls with a duration of several hours.

PSR J1727–2739: This pulsar nulls with a typical duration of 30 s or longer, and remains on for typically twice this duration. The overall nulling fraction is therefore about 35 per cent.

PSR J1733–3716: This pulsar has a peculiar profile with two separated components each with the steep edge facing the middle of the profile. The leading component is brighter and modulates strongly, both components switch off simultaneously and the nulls are quasi-periodic and of short duration. A single-pulse study of this pulsar is given in Hu et al. (2020).

PSR J1745–3040: This pulsar also has a peculiar profile that consists of a small leading component almost disjoint from a complex trailing component. Although nulling is not evident, various parts of

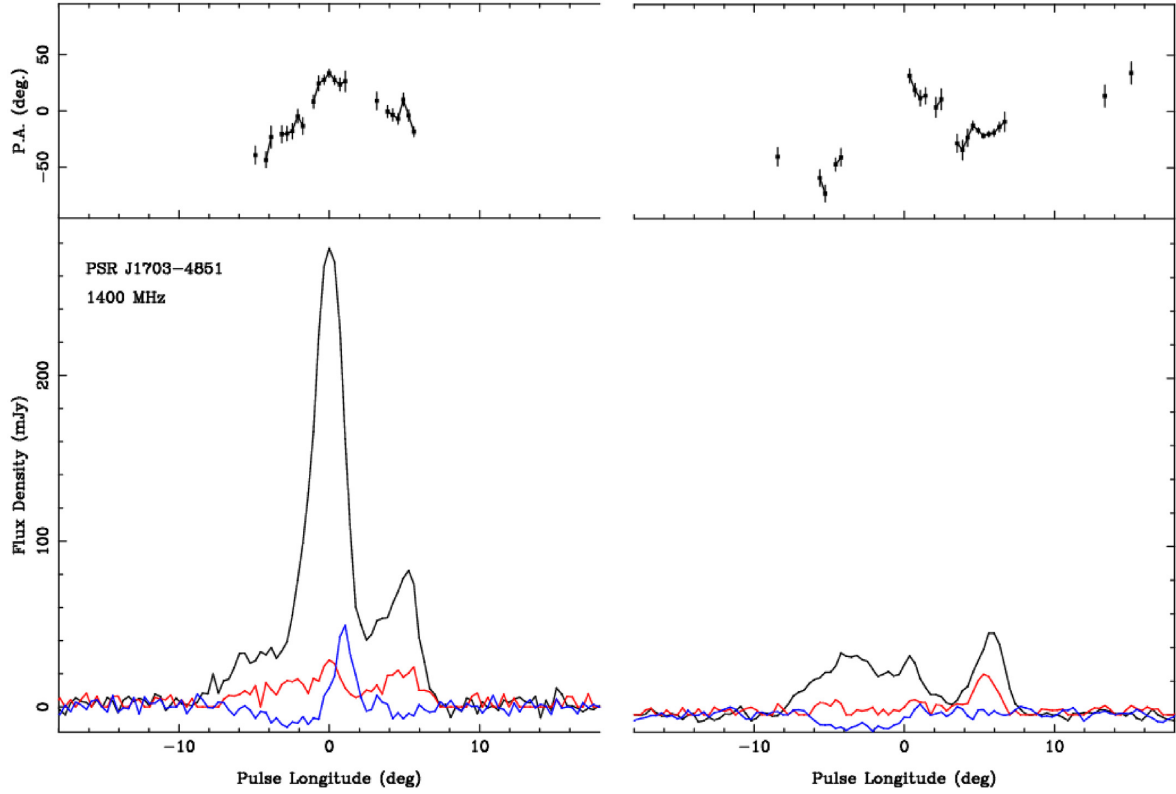


Figure 3. The two modes of PSR J1703–4851. Left-hand panel shows the bright mode with a dominant central component. The right-hand panel shows the weak mode. The central observing frequency is 1400 MHz with a total bandwidth of 256 MHz.

the profile seem to switch off at various times and the flux in the leading component is often anticorrelated with the flux of the trailing components.

PSRs J0034–0721, J1825–0935, and J1830–1059 are also part of our sample. These pulsars have a significant literature and their details will not be repeated here (see e.g. Ilie et al. 2020; Gil et al. 1994; Stairs et al. 2019, respectively).

4 FLUX DENSITIES AND SPECTRA

The UWL allows us to measure the flux density and the spectrum of a pulsar simultaneously over a wide bandwidth and to repeat this on a monthly basis. Flux densities are measured by supplying a template to the PSRCHIVE routine PSRFLUX. Fig. 4 shows the example of PSR J1157–6224, measured over 20 epochs. The spectral index is -2.6 ± 0.05 , fluctuations in the flux density are likely the result of refractive scintillation (Kumamoto et al. 2021) in this high-DM ($325 \text{ cm}^{-3} \text{ pc}$) pulsar.

A comprehensive survey of pulsar flux densities at a wide range of frequencies was carried out by Jankowski et al. (2018). In that paper, they identified five different spectral types: (i) simple power-law, (ii) broken power-law, (iii) log-parabolic, (iv) power-law with high-frequency cut-off, and (v) power-law with low-frequency turnover. Many of the pulsars in our sample were classified in that paper and the results will not be repeated here. However, our sample contains 44 pulsars not in the Jankowski et al. (2018) paper. For these pulsars, we summed together all the observations to smooth over the effects of diffractive and refractive scintillation. We then measured the flux density of each pulsar at 1400 MHz and examined their spectra. The spectrum for both a simple power-law and a log-parabolic can be

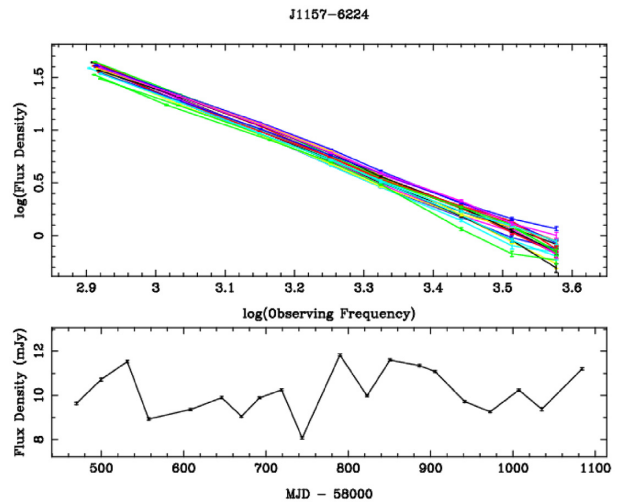


Figure 4. Top panel: Flux density versus observing frequency on a log–log scale for PSR J1157–6224 over 20 epochs. The spectral index is -2.6 ± 0.05 . Bottom panel: Flux density at 1.4 GHz as a function of epoch.

described by

$$\log_{10} S_\nu = a[\log_{10}(\nu/\nu_0)]^2 + b[\log_{10}(\nu/\nu_0)] + c, \quad (1)$$

where S_ν is the flux density at a frequency ν and ν_0 is 1400 MHz. For the power-law case, $a = 0$ and b yields the spectral index. For the log-parabolic case, a is the curvature and b gives the local value of the spectral index at ν_0 . Table 1 gives the results.

Table 1. Flux density and spectral information for 44 pulsars. Columns 2 and 3 show the flux density at 1400 MHz (S_{1400}) and the associated uncertainty, σ_S . Column 8 denotes whether the fit is a power-law (PL) or log-parabolic (LP). In the former case, a denotes the spectral index. For the latter case, a is the curvature parameter and b the spectral index, with σ_a and σ_b their respective uncertainties. PL-T denotes a low-frequency turnover, see text for details.

Jname	S_{1400} (mJy)	σ_S (mJy)	a	σ_a	b	σ_b	Type
J0525+1115	1.94	0.02	-2.27	0.08	.	.	PL
J0631+1036	1.11	0.01	-0.33	0.03	.	.	PL-T
J0738-4042	112.58	0.54	-1.02	0.10	-1.38	0.03	LP
J0842-4851	1.07	0.01	-1.78	0.06	.	.	PL
J1012-5857	1.91	0.01	-1.26	0.08	-1.24	0.02	LP
J1017-5621	2.20	0.01	-0.93	0.16	-2.13	0.05	LP
J1046-5813	1.37	0.01	-2.19	0.10	.	.	PL
J1110-5637	3.29	0.02	-0.85	0.11	-1.23	0.03	LP
J1114-6100	5.36	0.02	-0.7	0.1	-0.53	0.03	LP
J1210-5559	1.27	0.01	-2.17	0.03	.	.	PL
J1225-6408	1.26	0.01	-2.12	0.12	.	.	PL
J1306-6617	4.91	0.02	-1.6	0.2	-1.25	0.05	LP
J1418-3921	0.95	0.01	-2.52	0.04	.	.	PL
J1428-5530	7.67	0.01	-2.3	0.1	.	.	PL
J1430-6623	13.60	0.02	-1.8	0.1	.	.	PL
J1544-5308	5.82	0.01	-1.71	0.06	.	.	PL
J1555-3134	4.24	0.03	-0.60	0.15	.	.	PL
J1557-4258	3.14	0.01	-2.56	0.06	.	.	PL
J1559-4438	41.06	0.05	-2.3	0.1	.	.	PL
J1602-5100	8.23	0.03	-2.08	0.03	.	.	PL
J1603-5657	0.93	0.01	-2.06	0.04	.	.	PL
J1623-4256	2.60	0.02	-2.35	0.05	.	.	PL
J1633-4453	2.76	0.01	-2.2	0.1	.	.	PL
J1645-0317	25.76	0.03	-2.6	0.2	.	.	PL
J1649-3805	1.75	0.02	-1.75	0.08	.	.	PL
J1651-5255	2.72	0.02	-2.14	0.06	.	.	PL
J1652-2404	1.44	0.01	-2.06	0.05	.	.	PL
J1705-1906	5.66	0.03	-1.54	0.04	.	.	PL
J1716-4005	1.79	0.02	-1.52	0.06	.	.	PL
J1720-2933	1.69	0.01	-2.05	0.07	.	.	PL
J1722-3632	2.90	0.02	-0.7	0.3	-1.1	0.1	LP
J1733-2228	4.21	0.02	-2.7	0.1	.	.	PL
J1735-0724	2.76	0.01	-2.28	0.02	.	.	PL
J1739-3131	7.04	0.04	-2.20	0.03	.	.	PL
J1750-3157	1.40	0.02	-1.68	0.04	.	.	PL
J1751-3323	1.67	0.02	-0.92	0.03	.	.	PL-T
J1816-2650	2.38	0.02	-2.70	0.02	.	.	PL
J1817-3837	2.09	0.01	-1.82	0.06	.	.	PL
J1820-0427	10.07	0.02	-2.45	0.01	.	.	PL
J1825-0935	11.86	0.07	-1.5	0.1	.	.	PL
J1834-0426	19.69	0.08	-1.60	0.02	.	.	PL
J1841-0345	2.07	0.07	-1.0	0.1	.	.	PL
J1845-0434	2.92	0.02	-0.9	0.1	.	.	PL
J2155-3118	0.66	0.02	-3.0	0.1	.	.	PL

Of the 44 pulsars, only two show evidence for a low-frequency turnover in the spectrum. PSRs J0631+1036 and J1751-3323 have spectra that peak near 1 GHz, but the limited data below this frequency do not allow us to fully quantify the spectrum. We therefore simply quote the spectral index for the power-law spectrum above 1 GHz in Table 1. Of the rest, seven are best fit with a log-parabolic spectrum and 35 with a power-law. The fraction of pulsars in the different classes is in line with the results found by Jankowski et al. (2018).

Table 2. Pulsars with significant values of ΔDM . The uncertainty in the slope is given in brackets and refers to the last digit(s).

Jname	DM ($\text{cm}^{-3} \text{ pc}$)	ΔDM ($\text{cm}^{-3} \text{ pc yr}^{-1}$)
J0835-4510	67.8	-0.024(2)
J0908-4913	180.5	+0.026(2)
J1048-5832	128.7	-0.054(4)
J1105-6107	271.4	+0.063(4)
J1602-5100	170.8	-0.019(3)
J1611-5209	127.3	-0.020(3)
J1825-1446	352.7	+0.169(20)
J1826-1334	231.5	+0.084(11)
J1832-0827	301.1	+0.045(6)
J1833-0827	410.4	+0.284(23)
J1835-1106	132.6	+0.027(4)

Table 3. Pulsars without previous RM measurements. The uncertainty is given in brackets and refers to the last digit.

Jname	RM (rad m^{-2})
J0820-3826	+122(6)
J1055-6028	+343(2)
J1114-6100	-6729(2)
J1515-5720	+41(7)
J1637-4642	-43.2(5)
J1643-4505	-858(1)
J1650-4921	-208(1)
J1716-4005	-526(3)
J1843-0702	+188(3)

5 DISPERSION MEASURE AND VARIABILITY

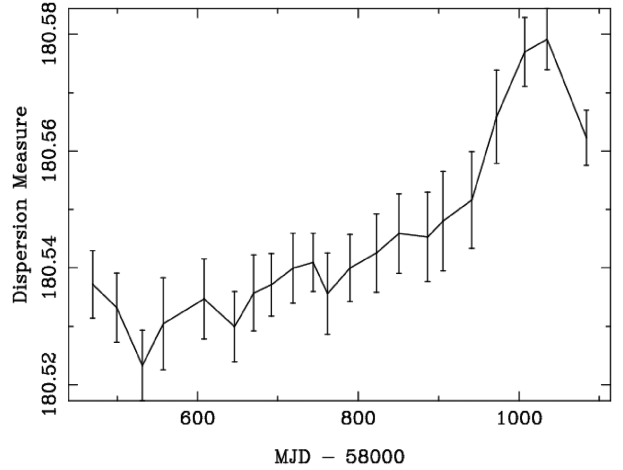
In this section, we concentrate on changes in DM over time, a quantity that is relatively simple to measure compared to the absolute DM value. Absolute DMs are difficult to measure because of profile variations with frequency and our lack of understanding of magnetospheric processes (see e.g. Oswald, Karastergiou & Johnston 2020). For millisecond pulsars, high precision to DM variations can be obtained (Pennucci 2019). You et al. (2007) and Keith et al. (2013) took advantage of this to examine the DM structure functions for a sample of 20 ms pulsars. The previous large-scale study of DM variability in slow pulsars was carried out by Petroff et al. (2013) who analysed 168 pulsars over an almost 6-yr period. They were able to detect significant variability in only four objects and placed upper limits on the rest. The work of Petroff et al. (2013) involved a mixture of narrow-band observations and non-simultaneous observations over several bands and had significant overlap with the pulsars under consideration here.

In order to measure a DM per epoch, we average over the observation duration, and reduce the number of frequency channels to eight across the band. We then use PSRCHIVE to produce arrival times (ToAs) via the routine PAT and fit the DM to these ToAs using TEMPO2. We discard observations that are strongly affected by RFI. We then fitted a straight line to the DM as a function of time. Currently 56 of the pulsars in our sample have an upper limit on DM changes of $0.01 \text{ cm}^{-3} \text{ pc yr}^{-1}$, already an improvement on Petroff et al. (2013), and which bodes well for the future of this project as the time baseline increases. We find 11 pulsars with slopes greater than 5σ deviant

Table 4. Pulsars with derived RMs that differ by more than 10 rad m^{-2} from the published values, RM_{cat} . The uncertainties, σ and σ_{cat} , are also listed.

Jname	RM (rad m^{-2})	σ (rad m^{-2})	RM_{cat} (rad m^{-2})	σ_{cat} (rad m^{-2})	Ref.
J0855–4644	+225	3	+249	22	(3)
J0857–4424	+165	4	–75	20	(5)
J0901–4624	+267	2	+289	22	(3)
J0954–5430	+86	2	+65	10	(3)
J1003–4747	+52	2	+18	4	(3)
J1012–5857	+39	4	+74	6	(3)
J1015–5719	+110	2	+96	2	(6)
J1016–5857	–520	2	–540	3	(6)
J1019–5749	–334	3	–366	10	(3)
J1034–3224	–43	2	–8	1	(5)
J1038–5831	–31	1	–15	10	(8)
J1043–6116	+189	2	+257	23	(3)
J1049–5833	+344	4	+359	11	(3)
J1146–6030	+7	2	–5	4	(8)
J1305–6203	–478	2	–436	15	(3)
J1327–6222	–327	2	–306	8	(3)
J1341–6220	–947	1	–921	3	(6)
J1356–5521	+89	2	+101	4	(8)
J1410–6132	+2270	5	+2400	30	(9)
J1412–6145	–52	2	–130	13	(6)
J1413–6141	–354	4	–35	10	(3)
J1424–5822	–643	3	–625	19	(3)
J1534–5334	+25	1	–46	17	(3)
J1536–5433	–138	3	–155	13	(3)
J1543–5459	+167	1	+28	23	(3)
J1544–5308	–42	1	–29	7	(5)
J1548–5607	+22	1	+37	10	(3)
J1600–5044	+134	1	+119	10	(11)
J1602–5100	+84	1	+71.5	1.1	(11)
J1604–4909	+7	2	+34	1	(7)
J1611–5209	–101	1	–79	5	(8)
J1638–4417	+139	3	+160	25	(12)
J1638–4608	+363	3	+335	12	(3)
J1640–4715	–422	2	–411	12	(8)
J1648–4611	–650	3	–682	26	(3)
J1715–3903	+228	1	+250	15	(3)
J1719–4006	–204	2	–218	17	(8)
J1720–2933	+10	4	+21	5	(2)
J1722–3632	–332	2	–307	8	(3)
J1731–4744	–446	2	–429.1	0.5	(11)
J1738–3211	–18	2	+7	9	(3)
J1739–2903	–301	2	–236	18	(3)
J1739–3023	–120	1	–74	18	(3)
J1740–3015	–155.3	0.4	–168.0	0.7	(8)
J1750–3157	+71	1	+111	8	(3)
J1757–2421	–27	1	+16	5	(8)
J1801–2304	–1124	2	–1156	19	(1)
J1806–2125	724	5	+796	15	(4)
J1822–4209	+54	10	–13	9	(3)
J1832–0827	+12	1	+39	7	(10)
J1833–0827	–284	3	–470	7	(10)
J1848–0123	+514	3	+580	30	(11)
J1853–0004	+627	4	+648.7	4.7	(4)

Note. References: (1) Crawford, Manchester & Kaspi (2001), (2) Hamilton & Lyne (1987), (3) Han et al. (2006), (4) Han et al. (2018), (5) Han, Manchester & Qiao (1999), (6) Johnston & Weisberg (2006), (7) Johnston et al. (2007), (8) Noutsos et al. (2008), (9) O’Brien et al. (2008), (10) Rand & Lyne (1994), (11) Taylor, Manchester & Lyne (1993), (12) Weltevrede & Johnston (2008).

**Figure 5.** Dispersion measure versus MJD for PSR J0908–4913. A straight line fit to the data yields a slope of $0.026 \pm 0.02 \text{ cm}^{-3} \text{ pc yr}^{-1}$.

from zero as listed in Table 2. Fig. 5 shows the data for PSR J0908–4913. Three of these objects (PSRs J0835–4510, J0908–4913, and J1833–0827) also had significant values of ΔDM in Petroff et al. (2013) but these were of opposite sign compared to our observations. This demonstrates that a simple slope is a crude tool to measure DM variations and that either a higher-order polynomial or a structure function analysis is required (e.g. Donner et al. 2020).

6 ROTATION MEASURE AND VARIABILITY

We note that, as with DM, values of RM for a given pulsar depend on the method used to obtain them, and can typically vary by several units. In addition, an incorrect DM can also lead to an erroneous RM (Ilie, Johnston & Weltevrede 2019). In order to determine the RM, we employ two different algorithms as implemented within the PSRCHIVE routine RMFIT. The first method uses a ‘brute-force’ approach in which trial RMs are employed and the amount of linear polarization is computed for each trial value. The algorithm returns the RM at which the linear polarization is maximized over the profile as a whole. In the second method, rotation measures are computed using the algorithm outlined in Noutsos et al. (2008). In brief, a single PA is computed per frequency channel after collapsing the pulse profile to a single-phase bin. The RM is then computed through a quadratic fit to the PA as a function of frequency and the statistical error determined through Monte Carlo trials.

There are nine pulsars in our sample that have had RMs measured for the first time. These are listed in Table 3. PSR J1643–4505, once the high RM is taken into account, joins the ranks of pulsars with high \dot{E} and very high linear polarization fraction. In contrast, the linear polarization fraction for the high \dot{E} pulsar PSR J1055–6028 remains low even after the RM correction. We note the extremely high value of RM uncovered for PSR J1114–6100, described in more detail in the subsection below. We have found a substantial number of pulsars for which our value of RM is more than 10 rad m^{-2} away from the published values as given in v1.63 of the ATNF pulsar catalogue (Manchester et al. 2005).¹ These pulsars are listed in Table 4. The published value of PSR J1413–6141 was seriously in error, applying the correct value reveals that this pulsar has only a moderate linear polarization fraction. For PSR J0857–4424, the published value from

¹<http://www.atnf.csiro.au/research/pulsar/psrcat/>

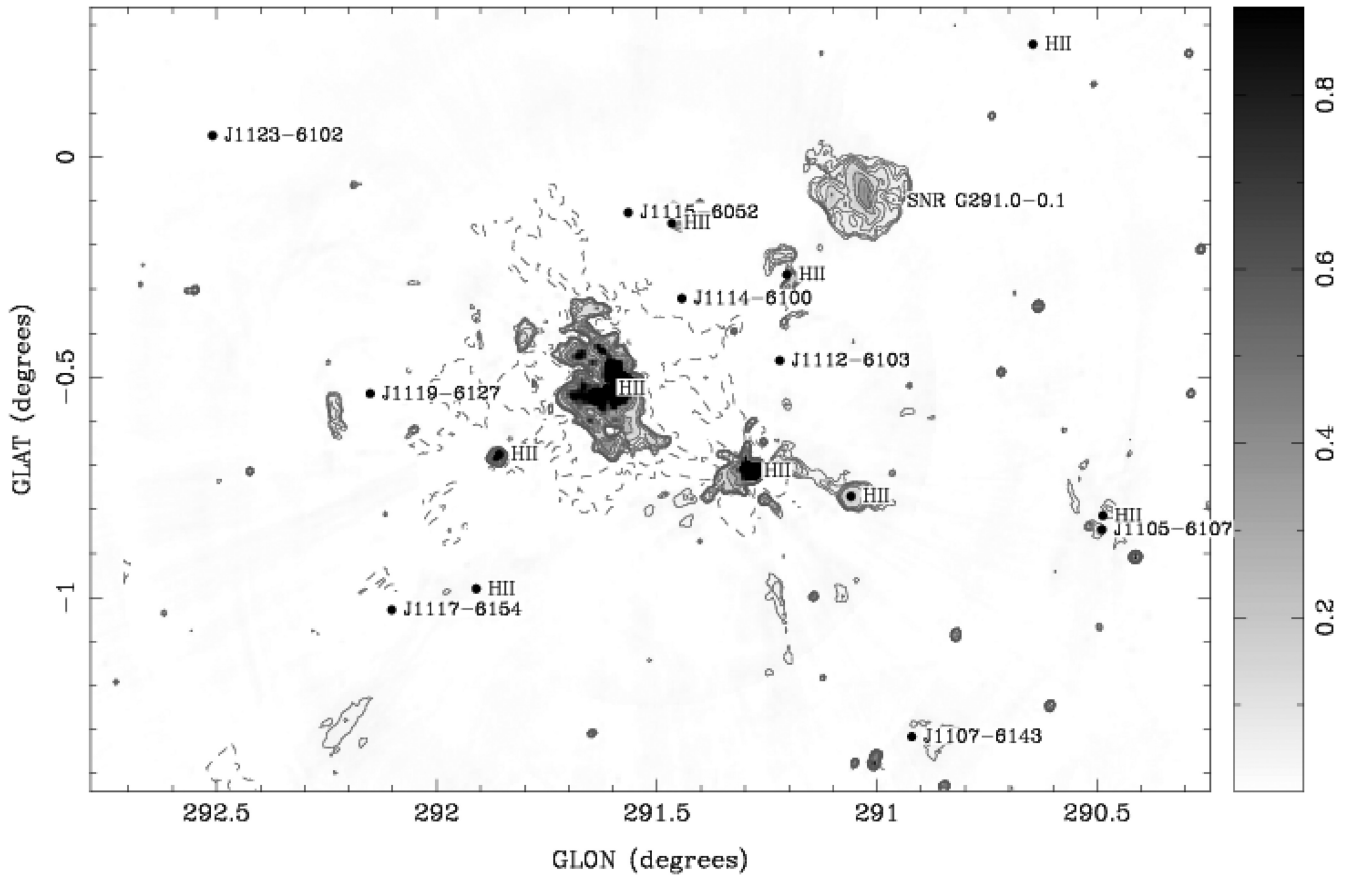


Figure 6. Radio continuum image at 843 MHz from the Molonglo Galactic Plane Survey (Green et al. 1999) of a region near PSR J1114–6100. Pulsars and H II regions are marked as is the SNR G291.0–0.1. The wedge to the right of the image shows the flux density levels in Jy.

20 yr ago of -75 rad m^{-2} is erroneous; we measure $+162 \text{ rad m}^{-2}$, a difference of more than 200 units.

We observe that there are sometimes significant differences between the RMs returned by the two methods as can be seen in Fig. 7. Karastergiou (2009) pointed out that the effects of interstellar scattering can lead to erroneous measurements of RM, and that the method of Noutsos et al. (2008) was best to use in the case of scattered profiles. In addition, Ilie et al. (2019) showed the presence of (apparent) RM variations across the pulse profile that meant that different RMs can be obtained depending on the method used. The most extreme example of this is PSR J1600–5044 for which we obtain an RM of 90 and 135 rad m^{-2} using the two methods. At low frequencies, the pulsar is heavily scattered, at high frequencies the position angle of the linear polarization has a very steep swing across the pulse profile. This is the scenario pointed out by Karastergiou (2009) as being the most detrimental to measuring correct values of RM.

6.1 PSR J1114–6100

We measured a value of $-6729 \pm 2 \text{ rad m}^{-2}$ for the RM of PSR J1114–6100. This is the fourth highest value of $|\text{RM}|$ for any known pulsar, beaten only by the Galactic centre magnetar PSR J1745–2900 (Shannon & Johnston 2013), and two other pulsars near the Galactic centre, PSRs J1746–2849 and J1746–2856 (Schnitzeler et al. 2016). The DM of the pulsar is $677 \text{ cm}^{-3} \text{ pc}$, and so the value of the magnetic field strength parallel to the line of sight, given by 1.2 RM/DM , is $-12.0 \mu\text{G}$, the third highest for any pulsar.

Fig. 6 shows the location of PSR J1114–6100 and other pulsars overlaid with a radio continuum image taken from the Molonglo Galactic Plane Survey (Green et al. 1999). Of the pulsars close on the sky to PSR J1114–6100, most have positive values of RM, including a value of $+853 \text{ rad m}^{-2}$ for PSR J1119–6127. PSRs J1112–6103 and J1115–6052 both have RMs close to $+250 \text{ rad m}^{-2}$. The prominent supernova remnant SNR G291.0–0.1 (Reynoso et al. 2006) lies close on the sky to the pulsar and there are numerous H II regions in the vicinity (Caswell & Haynes 1987). The largest and brightest H II region located at $(l, b) = (291.614, -0.525)$ has a flux density in excess of 200 Jy ; its recombination lines have a positive velocity indicating a distance of some 8 kpc. This H II region (also known as NGC 3603) is seen from radio to gamma-rays (Moffat et al. 2002; Saha et al. 2020), and is the subject of a significant body of literature; a recent study of its OB stars places its distance at 7 kpc (Drew, Monguió & Wright 2019). To the south-west lies the only H II region $(291.284, -0.713)$ with negative velocities at a distance of some 4 kpc. In the optical, the H II regions are clearly delineated and do not extend as far as PSR J1114–6100. H α images show some excess diffuse structure near the pulsar. The pulsar shows a relatively low amount of scatter broadening given the large value of DM.

The cause of the extremely high RM of PSR J1114–6100 is unclear. Its distance, derived from the DM (Yao, Manchester & Wang 2017), is roughly consistent with the distance to NGC 3603 but the H II region appears to be too distant (in angle) to contribute to the DM of pulsar. It could be that the pulsar lies behind some highly magnetized filament located perhaps in the near part of the Carina

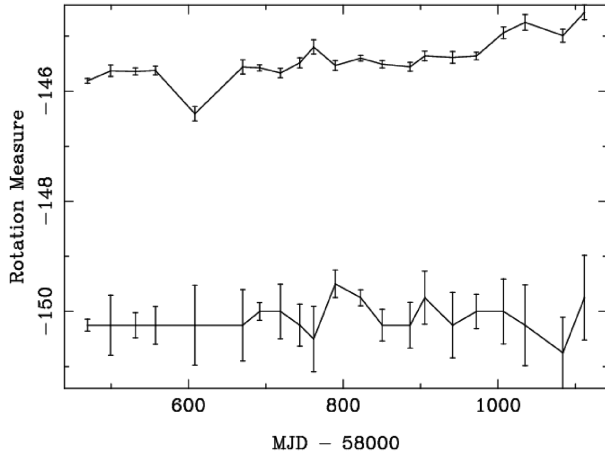


Figure 7. Rotation measure versus epoch for PSR J1048–5832. Top curve shows the RM using the ‘brute-force’ method, bottom curve shows the RM using the Noutsos et al. method.

spiral arm. In any case, it remains odd that the pulsar’s RM is an order of magnitude greater than its neighbouring pulsars, and pulsars in other, more complex regions of the Galactic plane.

6.2 Time-variable RMs

There are 10 pulsars in our sample for which the statistical error in the measured RM is smaller than 1.0 rad m^{-2} per epoch. These are the brightest and/or the most highly polarized pulsars, PSRs J0738–4042, J0742–2822, J0835–4510, J0908–4913, J1048–5832, J1359–6038, J1644–4559, J1709–4429, J1740–3015, and J1745–3040. In no case is there evidence for any change in RM. As an example, Fig. 7 shows RM versus epoch for PSR J1048–5832. The error bars per epoch reflect statistical errors and do not take into account the ionospheric RM. However, there does appear to be one pulsar that does show RM changes, PSR J1825–1446, a pulsar which we noted earlier also showed DM changes. This pulsar is discussed further below. The pulsar with the second largest $|\text{RM}|$ in our sample is PSR J1410–6132, which has a value of $+2400 \text{ rad m}^{-2}$ according to O’Brien et al. (2008) but for which we measured $+2279 \text{ rad m}^{-2}$, a substantial change. Our data show a decline in RM from 2290 to 2270 rad m^{-2} over the course of 2 yr although the significance of the change is only 2σ . However, it could be the case that the RM has decreased substantially since 2008 and further monitoring is required.

6.3 PSR J1825–1446

This pulsar is the only one of our sample that shows significant changes in both RM and DM, shown in Fig. 8. The RM increases almost linearly by about 8 rad m^{-2} over 2 yr (much larger than expected from the ionosphere) while the DM increases by about $0.3 \text{ cm}^{-3} \text{ pc}$. This implies a change in the magnetic field along the line of sight of $0.2 \mu\text{G}$ in 2 yr. An independent distance to this pulsar has not been determined but the distance inferred from the DM is $\sim 5 \text{ kpc}$. The proper motion is high, as measured by Moldón et al. (2012) and later refined by Dexter et al. (2017) and hence the pulsar has a high transverse velocity of 750 km s^{-1} . The pulsar also appears in projection to be located within the shell of the supernova remnant G16.8–1.1 although Moldón et al. (2012) argue that the pulsar’s

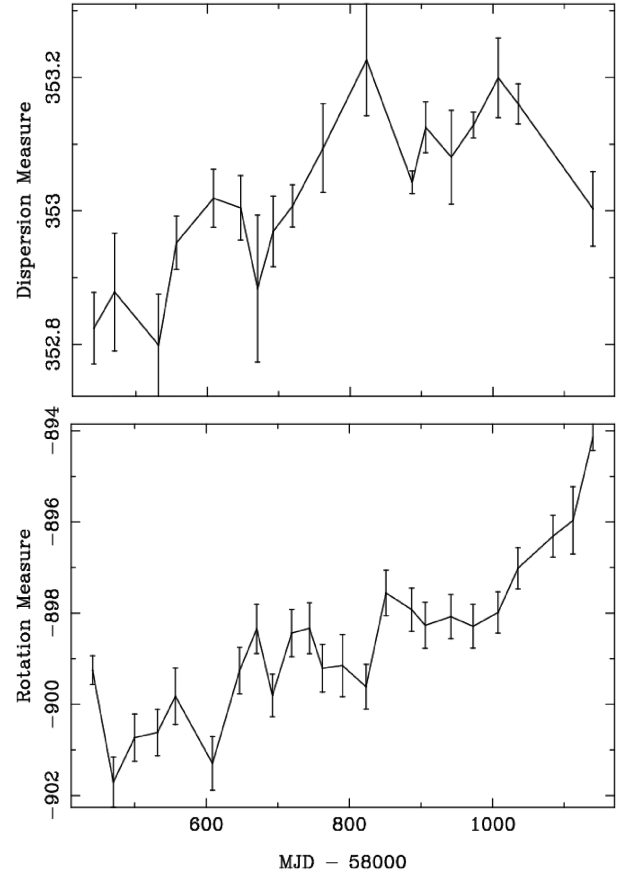


Figure 8. Dispersion measure (top panel) and rotation measure (bottom panel) as a function of epoch for PSR J1825–1446. A straight line fit to the DM versus time yields a slope of $0.15 \pm 0.02 \text{ cm}^{-3} \text{ pc yr}^{-1}$, that to the RM versus time yields $2.6 \pm 0.3 \text{ rad m}^{-2} \text{ yr}^{-1}$.

age and velocity make a true association unlikely. Nevertheless, our line of sight to the (background) pulsar intercepts the SNR and it seems likely that the RM and DM changes we see are the result of the pulsar passing behind a magnetized filament in the SNR. The scattering time-scale of the pulsar is 21 ms at 1 GHz (Oswald et al. 2021). This is an order of magnitude higher than expected from the extrapolation of the relationship of Krishnakumar et al. (2015) to 1 GHz, but is not particularly discrepant compared to other pulsars with comparable DMs (Oswald et al. 2021).

7 SUMMARY

The ultra-wide-band receiver on the Parkes telescope has been operational for 2 yr and will remain the workhorse for pulsar observations through this decade. A complementary programme of slow pulsar observations on the MeerKAT telescope is also underway (Bailes et al. 2020; Johnston et al. 2020) and will be key for single-pulse analysis. Repeated monitoring of a large sample of pulsars coupled with the wide instantaneous observing bandwidth is ideal for studying the time evolution of the pulsars themselves and the ISM through which their radio emission propagates. The early results presented here are extremely encouraging for the future of this project. We have reported on a number of new nulling pulsars, DM and RM variability, and a pulsar with an extremely high value of RM.

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DATA AVAILABILITY

Pulsar data taken for the P574 project is made available through the CSIRO's Data Access Portal (<https://data.csiro.au>) after an 18-month proprietary period.

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APPENDIX A: PULSAR LIST

Table A1. List of the 276 pulsars monitored with the UWL on a monthly basis.

PSR									
J0034–0721	J0108–1431	J0134–2937	J0151–0635	J0152–1637	J0206–4028	J0255–5304	J0304+1932	J0401–7608	J0448–2749
J0452–1759	J0525+1115	J0536–7543	J0543+2329	J0601–0527	J0614+2229	J0624–0424	J0627+0706	J0630–2834	J0631+1036
J0659+1414	J0729–1448	J0729–1836	J0738–4042	J0742–2822	J0745–5353	J0758–1528	J0809–4753	J0820–1350	J0820–3826
J0820–4114	J0834–4159	J0835–4510	J0837+0610	J0837–4135	J0842–4851	J0855–4644	J0857–4424	J0901–4624	J0904–7459
J0905–5127	J0907–5157	J0908–1739	J0908–4913	J0924–5814	J0940–5428	J0942–5552	J0954–5430	J0959–4809	J1001–5507
J1003–4747	J1012–5857	J1015–5719	J1016–5819	J1016–5857	J1017–5621	J1019–5749	J1028–5819	J1034–3224	J1038–5831
J1043–6116	J1046–5813	J1047–6709	J1048–5832	J1049–5833	J1055–6028	J1056–6258	J1057–5226	J1105–6107	J1110–5637
J1112–6103	J1114–6100	J1115–6052	J1119–6127	J1123–6259	J1136–5525	J1146–6030	J1156–5707	J1157–6224	J1210–5559
J1224–6407	J1225–6408	J1243–6423	J1253–5820	J1301–6305	J1302–6350	J1305–6203	J1306–6617	J1317–6302	J1319–6056
J1320–5359	J1326–5859	J1326–6408	J1326–6700	J1327–6222	J1327–6301	J1328–4357	J1338–6204	J1340–6456	J1341–6220
J1349–6130	J1352–6803	J1356–5521	J1357–62	J1357–6429	J1359–6038	J1401–6357	J1410–6132	J1412–6145	J1413–6141
J1418–3921	J1420–6048	J1424–5822	J1428–5530	J1430–6623	J1435–5954	J1452–6036	J1453–6413	J1456–6843	J1509–5850
J1512–5759	J1513–5908	J1515–5720	J1522–5829	J1524–5625	J1524–5706	J1530–5327	J1531–5610	J1534–5334	J1534–5405
J1535–4114	J1536–5433	J1539–5626	J1541–5535	J1543–5459	J1544–5308	J1548–5607	J1549–4848	J1555–3134	J1557–4258
J1559–4438	J1600–5044	J1600–5751	J1602–5100	J1603–5657	J1604–4909	J1605–5257	J1611–5209	J1613–4714	J1614–5048
J1623–4256	J1626–4537	J1630–4733	J1632–4621	J1633–4453	J1633–5015	J1637–4553	J1637–4642	J1638–4417	J1638–4608
J1638–4725	J1640–4715	J1643–4505	J1644–4559	J1645–0317	J1646–4346	J1646–6831	J1648–3256	J1648–4611	J1649–3805
J1649–4653	J1650–4502	J1650–4921	J1651–4246	J1651–5222	J1651–5255	J1652–2404	J1653–3838	J1653–4249	J1700–3312
J1701–3726	J1701–4533	J1702–4128	J1702–4310	J1703–3241	J1703–4851	J1705–1906	J1705–3423	J1705–3950	J1707–4053
J1707–4729	J1708–3426	J1709–1640	J1709–4429	J1715–3903	J1715–4034	J1716–4005	J1717–3425	J1718–3825	J1719–4006
J1720–2933	J1721–3532	J1722–3207	J1722–3632	J1722–3712	J1723–3659	J1727–2739	J1730–3350	J1731–4744	J1733–2228
J1733–3716	J1735–0724	J1737–3137	J1738–3211	J1739–2903	J1739–3023	J1739–3131	J1740–3015	J1741–2733	J1741–3016
J1741–3927	J1743–3150	J1745–3040	J1749–3002	J1750–3157	J1751–3323	J1751–4657	J1752–2806	J1757–2421	J1801–2304
J1801–2451	J1803–2137	J1806–2125	J1807–0847	J1809–1917	J1816–2650	J1817–3618	J1817–3837	J1820–0427	J1822–2256
J1822–4209	J1823–3106	J1824–1945	J1825–0935	J1825–1446	J1826–1334	J1828–1101	J1829–1751	J1830–1059	J1832–0827
J1833–0827	J1834–0426	J1835–0643	J1835–1106	J1837–0559	J1841–0345	J1841–0425	J1842–0905	J1843–0702	J1844–0538
J1845–0434	J1845–0743	J1847–0402	J1848–0123	J1852–0635	J1852–2610	J1853–0004	J1900–2600	J1910+0358	J1913–0440
J1932+2220	J1941–2602	J2048–1616	J2155–3118	J2330–2005	J2346–0609				

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