













The *Gaia* spectrophotometric standard stars survey – V. Preliminary flux tables for the calibration of *Gaia* DR2 and (E)DR3

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ABSTRACT

We present the flux tables of the spectrophotometric standard stars (SPSS) used to calibrate in flux the *Gaia* DR2 and (E)DR3 data releases. The latest SPSS grid version contains 112 stars, whose flux tables agree to better than 1 per cent with the CALSPEC spectra of 11 flux standards for the calibration of the *Hubble Space Telescope*. The synthetic magnitudes computed on the SPSS spectra also agree to better than 1 per cent with the Landolt magnitudes of 37 stars in common. The typical spreads in both comparisons are of the order of 1 per cent. These uncertainties already meet the initial requirements for the *Gaia* SPSS project, but further improvements are expected in the next SPSS versions, that will be used to calibrate future *Gaia* releases. We complement the SPSS flux tables with literature spectra of 60 additional stars that did not pass all the criteria to be SPSS, the Passband Validation Library (PVL). The PVL contains stars of extreme spectral types, such as bright O and B stars and late M stars and brown dwarfs, and was useful to investigate systematic effects in the previous *Gaia* DR2 release and to minimize them in the EDR3 one. The PVL literature spectra are recalibrated as accurately as possible on to the SPSS reference scale, so that the two sets together can be used in a variety of validation and comparison studies.

Key words: techniques: spectroscopic – catalogues – surveys – stars: general.

1 INTRODUCTION

The ESA astrometric mission *Gaia*¹ (Gaia Collaboration 2016a) has revolutionized our view of the Milky Way and its data are bringing progress in almost every area of astrophysical research (Perryman et al. 2001; Gilmore 2018; Pancino 2020). Three data releases were published so far,² with a constant increase in the data quality and in the quantity of data products (Gaia Collaboration 2016b, 2018, 2020). Several thousand publications in different areas of astrophysical research cite or use *Gaia* data, from Solar system objects and fundamental physics to the study of stars, stellar clusters, the Milky Way, and the Local Group. *Gaia* also tackles time-variability phenomena and provides large samples of external galaxies and distant quasars. The EDR3 (Early Data Release 3) *Gaia* catalogue contains astrometry, i.e. positions, parallaxes, and proper

motions for almost two billions of stars and non-stellar sources, down to a magnitude of $G \simeq 21$ mag ($V \simeq 20.5$; see Riello et al. 2020). The quality of the *Gaia* astrometry is unprecedented: errors on the EDR3 measurements are of the order of 10–100 μ as (20–130 μ as on parallaxes and 20–140 μ as yr⁻¹ on proper motions, see Lindegren et al. 2020). The data are complemented by three-band photometry (Riello et al. 2020) and the upcoming DR3 (Data Release 3) will present low-resolution spectra covering the optical range (330–1050 nm). Line-of-sight velocities are also presented for more than seven million stars down to $G \lesssim 14$ mag (Katz et al. 2019), thanks to higher resolution ($R = \lambda/\delta\lambda \simeq 11\,500$) spectra in the Calcium triplet region, and the *Gaia* data are complemented by stellar parameter estimates (Andrae et al. 2018). The multi-epoch observations of *Gaia* allow for time-domain studies as well (Gaia Collaboration 2019).

The complex *Gaia* data analysis system, developed within the Data Processing and Analysis Consortium (DPAC), provides essentially self-calibrated data products. However, for some deliverables it is necessary to rely on external catalogues to bring the *Gaia* measurements on to an external physical system of standardized

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¹<https://www.cosmos.esa.int/web/gaia>

²<https://www.cosmos.esa.int/web/gaia/release>

quantities. This is, for example, the case of the radial (or line-of-sight) velocities, where a specific set of thousands of carefully selected standard stars, constancy monitored to be stable within $\simeq 300 \text{ m s}^{-1}$, was assembled for the specific purpose of calibrating *Gaia* high-resolution spectra (Soubiran et al. 2018). The external calibration is even more necessary in the case of fluxes and magnitudes, because *Gaia* counts electrons like any typical modern astronomical detector, and thus an external calibration is necessary to transform the measurements into a physical scale. Unfortunately, only a few primary flux standards exist in the literature, such as Vega or Sirius (see also Section 7 for more details), and they are too bright for *Gaia*, which is not designed for bright stars and is incomplete above 3–6 mag. The most suitable and reliable sets of secondary standards to calibrate *Gaia* fluxes are the CALSPEC grid of spectrophotometric standard stars³ used to calibrate the *Hubble Space Telescope* (*HST*; Bohlin 2014) and the Landolt standards for the Johnson–Kron–Cousins broad-band photometric system (Landolt 1992a, see also Section 5.2.3). We thus decided to tie the *Gaia* flux calibration to the three pure hydrogen white dwarfs (WD) used as reference in CALSPEC, that we adopted as the pillars of our calibration (Pancino et al. 2012). However, *Gaia* is a sophisticated space observatory, capable of delivering photometry with a precision of a few mmag. To reach the best possible accuracy, the external calibration model of *Gaia* requires a large set of Spectro-Photometric Standard Stars (SPSS) subject to well-defined requirements (Pancino et al. 2012, see also Section 2.1). Unfortunately, the available standard grids did not contain the required number of stars (about 200) meeting our requirements. We thus started a new set of ground-based observing campaigns to build a specific *Gaia* SPSS grid, calibrated on the flux system defined by CALSPEC.

This is the fifth paper of the *Gaia* SPSS series, preceded by: a presentation of the SPSS survey and some preliminary results (Pancino et al. 2012); the description of the data reduction methods and the characterization of the instruments used (Altavilla et al. 2015); the results of the constancy monitoring of SPSS candidates (Marinoni et al. 2016); and the presentation of the calibrated magnitudes of candidate SPSS (Altavilla et al. 2021). We present here two preliminary versions of the SPSS flux tables. The SPSS V1 was used to calibrate the photometry in the second *Gaia* release (DR2; Evans et al. 2018; Gaia Collaboration 2018); the SPSS V2 was used to calibrate the photometry in EDR3 (Gaia Collaboration 2020; Riello et al. 2020) and will be used to calibrate the low-resolution spectra in DR3, expected in the first half of 2022. We complement the paper by presenting also the PVL, that was built with the purpose of investigating systematic effects in the *Gaia* photometry and passband reconstruction, and more in general as a validation sample. The PVL contains stars of more extreme spectral types that did not fulfil all the criteria to be selected as SPSS (Section 2.1) and is based on literature data.

The paper is organized as follows: in Section 2, we discuss the SPSS selection criteria and the observing campaigns; in Section 3, we present the procedures employed for the spectra reductions; in Section 4, we describe the flux calibration procedure and the flux tables preparation; in Section 5, we describe the SPSS results and validate them against other literature samples; in Section 6, we present the PVL; and in Section 7, we summarize our results and discuss them in the context of the existing data sets and future spectrophotometric projects.

³<https://www.stsci.edu/hst/instrumentation/reference-data-for-calibration-and-tools/astronomical-catalogs/calspec>

2 SPSS SELECTION AND OBSERVATIONS

2.1 Candidate SPSS selection

The selection criteria for the candidate SPSS stars were detailed in Pancino et al. (2012). They were set to meet the *Gaia* calibration needs and concerned the resolution ($R = \lambda/\delta\lambda \simeq 1000$, a few times the *Gaia* BP/RP resolution), S/N ratio (> 100), magnitude ($9 \lesssim V \lesssim 15$ mag) distribution, colour distribution (as wide as possible, see below), and sky distribution of the targets for the ground-based observations, having in mind that the stars should be observable with 2–4 m telescopes from both hemispheres. The *Gaia* requirements imposed the coverage of the entire wavelength range and a set of at least a hundred SPSS of different spectral types, to match the target of $\simeq 1$ per cent precision and $\simeq 1$ –3 per cent accuracy in the resulting flux tables. This led to an initial selection of about 300 candidate SPSS. An additional requirement of short-term flux variations smaller than ± 0.5 per cent (an amplitude of ± 0.005 mag) was imposed and all SPSS candidate close enough to some instability strip were monitored for constancy, as described by Marinoni et al. (2016). Not even the CALSPEC grid of flux standard met simultaneously all the above criteria and thus we started a dedicated spectrophotometric observing campaign.

The most difficult criterium to meet was the colour and spectral type coverage. To allow for the best *Gaia* flux calibration, the SPSS needed to be as different as possible from each other, with different overall spectral energy distributions. This was not easy also when considering our constancy criterium, because most of the objects with emission lines tend to be variable (including hot stars and quasars). We also realized that featureless white dwarfs (WD), that are generally considered to be the best calibrators for obvious reasons (see e.g. Moehler et al. 2014), are subject to large flux calibration uncertainties. In fact, even small wavelength calibration errors might induce large flux errors in the steep blue part of the spectrum (Fig. 1, top panel). These errors cannot be repaired using stellar features, which are absent, or sky emission and absorption lines, because we observed at low resolution and with wide slits. Similarly, cool stars with many molecular absorption bands of varying depth can have locally very uncertain flux levels (Fig. 1, bottom panel). The local flux in fact varies significantly with the unavoidable small resolution variations in wide slit observations, where the effective resolution is governed by the atmospheric seeing.

Therefore, the initial list of about 200 good SPSS candidates contained a relatively small number of O and B stars, and no spectral types later than about M2 (Pancino et al. 2012). The list of SPSS used for the calibration of *Gaia* DR2 (Gaia Collaboration 2018) and *Gaia* (E)DR3 (Gaia Collaboration 2020) is a subset of that initial list with a handful of additions, and can be found in Table 1 along with relevant information.

2.2 Observations

Observations started in the second half of 2006 and were completed in 2015 July. Overall, we were awarded more than 5000 observing hours at seven different observing facilities, of which a large fraction was carried out in visitor mode. We carried out three main campaigns: (i) a constancy monitoring campaign, described in details in Marinoni et al. (2016); (ii) an absolute photometry campaign, described in details in Altavilla et al. (2021); and (iii) a spectrophotometric campaign, for which we present the first results here. We also started a fourth campaign, devoted to the constancy monitoring on longer time-scales (about three years) (see Pancino et al. 2012, and

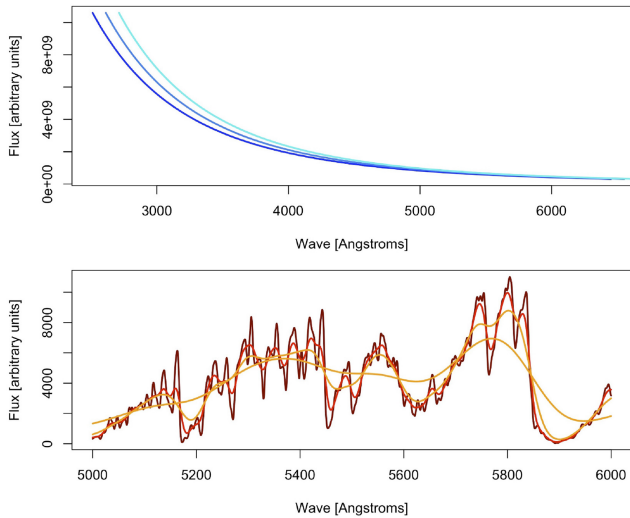


Figure 1. Cartoon illustrating the risks of using certain spectral type stars as flux calibrators. Top panel: when using a featureless WD (or a very hot star), uncertainties in the wavelength calibration can imply large flux variations in the steepest blue spectral regions. Three black-body spectra (blue shades) are shifted by 10 nm, to make the effect clearly appreciable by eye (our uncertainties are of ≈ 0.2 nm). Bottom panel: when using cool stars with deep molecular bands, variations in the effective spectral resolution, that can occur when using wide slits with variable seeing conditions, can imply large local flux variations. The mock spectra have resolution varying in a wide range from 0.5 (darkest) to 10 nm (lightest), to make the effect clearly appreciable visually. Section 2.1.

Section 5.1). However, we did not have the resources to carry it out to completion and we considered that *Gaia* itself will be able to help prune out any remaining long-term variable SPSS (Section 5.1). In this paper, we present two versions of the SPSS flux tables library: the first (SPSS V1) was used to calibrate *Gaia* DR2 and contains 94 SPSS observed in the best nights up to 2013; the second (SPSS V2) was used to calibrate *Gaia* (E)DR3 and contained more SPSS (112)⁴ and more spectra for each SPSS, including spectra observed after we completed the V1 flux tables in 2013.

The spectrophotometric observations were carried out using four different facilities, listed in Table 2, along with some relevant information. All spectra were taken with wide slits, to minimize (differential) flux losses, but narrow slit spectra were also obtained to allow for a more accurate wavelength calibration (Section 3.1). Three wide exposures and one narrow exposure were taken for each SPSS in each of two set-ups, a blue one and a red one, covering the entire *Gaia* wavelength range (about 330–1050 nm, see Table 2). Each star was observed more than once, whenever possible, in different nights and with different facilities. For the flux calibrations, we observed several times at different airmasses, a set of Pillars and Primary SPSS⁵ as the night flux calibrators (see Pancino et al. 2012, and Section 4 for more details). Daytime and twilight calibrations were taken daily (see the next section for more details) and we performed several dedicated experiments to characterize the stability and quality

⁴The actual SPSS list used to calibrate *Gaia* EDR3 contained 113 stars, but G 184-20 (SPSS 140) was later rejected because of identification problems in our ground-based observations.

⁵As described in Pancino et al. (2012), the Primary SPSS are well known stars in the literature, mostly from the CALSPEC data base, that can be used as calibrators when the Pillars are not visible.

of the calibration data for each facility (see Altavilla et al. 2015, and following sections for more details).

3 SPSS DATA REDUCTIONS

3.1 Basic spectra reductions

Every night we obtained at least 10 bias frames and 5 spectroscopic flat-field frames for each setup used during the night. Some dark frames were taken as part of our instrument familiarization plan (IFP; Altavilla et al. 2015), but we found that a dark correction was not necessary. In the few nights in which good calibrations could not be obtained, we used the results of our IFP calibration frames stability study (with a threshold of ± 1 per cent) to decide whether it was safe to use the master frames obtained in other nights or it was better to reject the data for the entire night. We used IRAF⁶ (Tody 1986, 1993) to create and apply master frames for each night, including the bias, flat-field, and illumination corrections. Whenever possible, we used the overscan strips to correct the bias level before applying the bias and flat corrections to the 2D spectroscopic frames. We also created bad pixel masks, roughly twice per year, and used them to exclude hot or dead pixels from the following reduction steps. An illumination correction, obtained with sky flats, was applied besides the usual master flat-field correction, obtained from lamp flats. More details on the pre-reduction and the quality control procedures can be found in Altavilla et al. (2015).

Once the 2D spectroscopic frames were pre-reduced as described above, we proceeded to extract the spectra and to calibrate them in wavelength using the IRAF task *apall* (Valdes 1992). To avoid flux losses, we observed the spectra with wide slits (10 or 12 arcsec, depending on the spectrograph). A spectrum was considered wide-slit if the ratio between the slit and the seeing during the observations was higher than 6, while it was considered narrow-slit if it was between 1.5 and 6, and it was only used for wavelength calibration purposes if it was below 1.5. To extract the spectra, after tracing them in the direction of dispersion and subtracting the sky level, we summed all the pixels along the slit in a window as large as six times the FWHM of the spectrum. We did not use the optimal extraction option (Valdes 1992) but we strictly summed up the pixels to minimize flux losses, profiting from the fact that the S/N ratio of our spectra was always above 100 (Fig. 2). The sigma spectrum (the square root of the counts at each wavelength) was computed to represent the actual uncertainty. The sigma spectrum was processed consistently and carried along with the spectra in every step of the reduction (see also the next sections), so that the uncertainty associated to each reduction step was taken into account.

To facilitate the wavelength calibration, that is complicated in the case of wide-slit observations, we also acquired narrow-slit spectra, generally using 2 or 2.5 arcsec slits. The narrow-slit spectra always preceded or followed immediately a wavelength calibration lamp, to minimize lamp flexure effects (Altavilla et al. 2015). In substance, we extracted the wavelength calibration lamps and used them to obtain a dispersion solution that was applied to both the narrow and wide spectra. The wide spectra were observed with long exposure times and repeated three times, thus they were generally observed far from the lamps. Moreover, the star centring within the slit was less controllable than with the narrow spectra. Therefore the wavelength calibration of wide spectra was adjusted by cross-correlation with the narrow spectra. With this procedure, and by comparing several

⁶<http://iraf.noao.edu>

Table 1. List of SPSS included in the V1 and V2 sets along with basic information: the SPSS ID and name; the *Gaia* EDR3 ID and coordinates (rounded in the pdf table version); the B and V magnitudes from Altavilla et al. (2021), except for LTT 377 (SPSS 350), where they are from Koen et al. (2010); the spectral type collected from the literature by Pancino et al. (2012); the SPSS set in which the star was included, either V1 for *Gaia* DR2 or V2 for *Gaia* (E)DR3; the quality flag (0 = Ok, 1 = warning on minor spectral defects, 2 = warning on spectral defects) and the notes. Only the first five lines are shown here; the table is available in its entirety online (see Data Availability Section).

ID	Name	<i>Gaia</i> EDR3 ID	RA (deg)	Dec (deg)	B (mag)	V (mag)	SpType	V1	V2	Flag	Notes
001	G191-B2B	266077145295627520	76.37767	52.83067	11.44	11.79	DA0	yes	yes	0	–
002	GD 71	3348071631670500736	88.11544	15.88624	12.79	13.03	DA1	yes	yes	0	–
003	GD 153	3944400490365194368	194.25949	22.03039	13.11	13.40	DA1	yes	yes	0	–
005	EG 21	4646535078125821568	47.62973	−68.60140	11.40	11.40	DA3	yes	yes	0	–
006	GD 50	3251244858154433536	57.20948	−0.97636	13.79	14.07	DA2	yes	yes	0	–

Table 2. Summary of the facilities used for the SPSS spectroscopy observations. For each facility, we report the number of used spectra (for V1 and V2) on the total spectra acquired with the facility, the number of used observing nights on the total nights awarded, and the name and the wavelength coverage (in nm) of the used grisms. Note that the weather conditions in Calar Alto and Loiano were rarely fully photometric (see Section 4), thus the spectra obtained with these facilities will mostly be used in a future SPSS release to increase the S/N ratio and to smooth out pixel-to-pixel defects for SPSS observed with NTT and TNG.

Instrument	Site	Spectra (used/all)	Nights (used/awarded)	Blue grism	Red grism	Third grism
EFOC2@NTT 3.6m	La Silla, Chile	896 / 2037	17 / 42	#11 (338–752)	#16 (602–1032)	#5 (520–935)
DOLoRes@TNG 3.6m	La Palma, Spain	650 / 2487	8 / 59	LR-B (300–843)	LR-R (447–1007)	–
CAFOS@2.2m	Calar Alto, Spain	15 / 995	1 / 90	B200 (320–900)	R200 (620–1100)	–
BFOC@Cassini 1.5m	Loiano, Italy	0 / 947	0 / 32	#3 (330–642)	#5 (480–980)	–

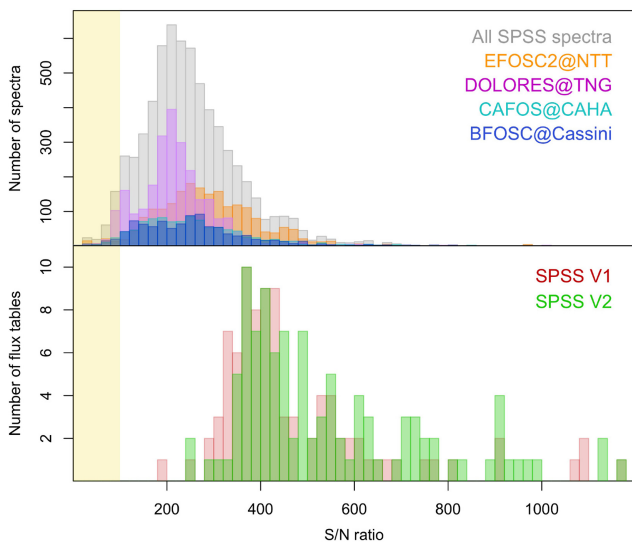


Figure 2. Histogram of the S/N ratio measured on the central third of the individual spectra and SPSS final flux tables. Top panel. The entire set of 6466 good spectra is plotted in grey in the background, while the sets obtained with each facility are plotted in different colours, as labelled. Bottom panel. The flux tables in SPSS V1 (94 stars) and SPSS V2 (112 stars) are plotted in different colours, as labelled. A few flux tables with S/N ratio between 1200 and 2300 were left out to make the plot more readable. The region with S/N ratio < 100, our S/N rejection criterion, is shaded in yellow in both panels.

spectra of the same star at the end, before creating the final flux tables (Section 4) we achieved a wavelength calibration accuracy of 0.15–0.40 nm, with a median of 0.2 nm.

Each data frame was subject to a rigorous quality control (QC) procedure (Altavilla et al. 2015) concerning the S/N ratio, the amount of saturated pixels, the presence of CCD defective pixels, the presence of artefacts or other problems. The QC system allowed for three levels: frames with no problems or with warnings were carried

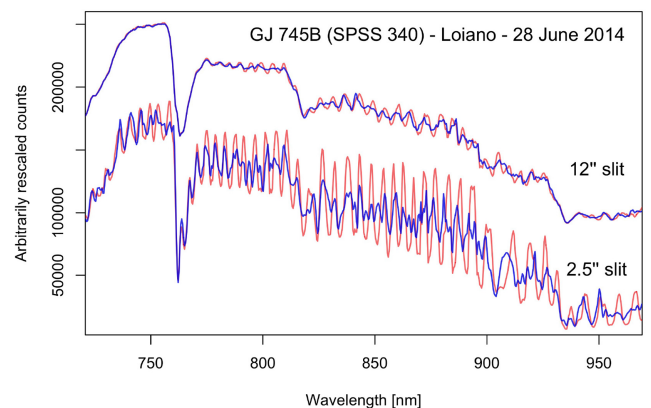


Figure 3. Example of fringing mitigation for the Secondary SPSS GJ 745B (SPSS 340), observed repeatedly with BFOC in Loiano on 2014 June 28. This is the facility with the strongest fringing pattern. Two of the seven uncorrected spectra are plotted in red, the corresponding corrected spectra are plotted in blue. The spectra taken with a 2.5 and a 12 arcsec slit are arbitrarily rescaled in flux for clarity, as labelled.

out to the following steps, while frames with serious problems were discarded.

3.2 Advanced spectra reductions

The extracted spectra were further manipulated to mitigate some typical instrumental effects: fringing, flux losses in case of narrow-slit spectra, telluric absorption, and second-order contamination. In the case of fringing, no procedure can fully remove its signature from spectra, unlike in the imaging case. Fringing only affected our red grisms and appeared much more pronounced in the case of narrow-slit spectra (see Fig. 3). We adopted the mitigation procedure described in details by Altavilla et al. (2015), that is based on the method originally adopted for STIS spectra (Malumuth et al. 2003). Briefly, the fringing pattern present in the spectroscopic flat fields is

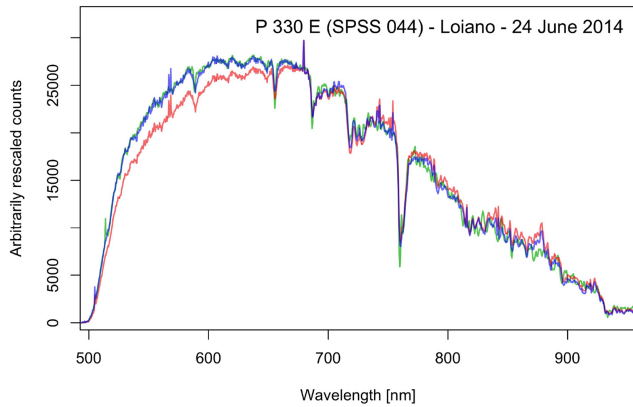


Figure 4. Example of slit-loss correction for a narrow-slit spectrum of the Primary SPSS P 330 E (SPSS 044), observed with BFOSC on the 2014 June 24. The reference wide-slit spectrum is plotted in green, the uncorrected narrow-slit spectrum in red, and the corrected one in blue.

shifted and scaled and then applied to the spectra, to minimize the fringing oscillations. The procedure is much more effective when the fringing pattern is more pronounced (see Fig. 3). The typical strength of the fringing pattern and the approximate starting wavelength for each facility can be found in table 4 by Altavilla et al. (2015). In summary, fringing starts at about 700–800 nm and varies from negligible to 15 per cent. With our procedure, we could beat down fringing to below 3 per cent in most spectra, except for a minority of cases. This uncertainty was used to update the sigma spectrum in the relevant wavelength regions. When combining different spectra for a given SPSS, fringing was further reduced to about 1 per cent or less, except for a few SPSS, because the fringing patterns are rarely aligned in wavelength for spectra obtained in different nights or with different facilities. Note that the fringing mitigation procedure was applied to the SPSS V2 spectra but not to the SPSS V1 ones.⁷

As mentioned, all the narrow-slit spectra, with a slit width smaller than six times the seeing and larger than 1.5, were later used together with the wide spectra to create the final flux tables (see Section 4.2). They were corrected for wavelength-dependent slit losses, caused by the fact that the seeing is larger at shorter wavelength and thus the flux loss is generally larger in the blue. Spectra with slits narrower than 1.5 times the seeing were only used for the wavelength calibration and then discarded. The actual slit loss at each wavelength can be complicated by other effects such as atmospheric diffraction⁸ and misalignment of the star in the slit, thus it cannot be modelled from first principles. The procedure relied on a wide-slit spectrum of the same SPSS, possibly obtained in the same night. The ratio between the narrow-slit and the wide-slit spectrum was fitted with a smooth function (typically a cubic spline of low order). The function was then applied to the narrow spectrum to partially recover the lost flux (see Fig. 4 for an example). The sigma spectra were worsened according to the correction applied.

We then took care of second-order contamination. With the goal

⁷The SPSS V1 release was used to calibrate *Gaia* DR2, which only contained integrated photometry, and thus the effect of fringing was expected to be negligible. In DR3 however, the first BP and RP spectra will be released. The fringing oscillations can create problems in the external calibration, at least locally, thus fringing mitigation was applied to V2 spectra.

⁸To minimize atmospheric diffraction effects, we always aligned the slit with the star’s parallactic angle. Exposures were rarely longer than one hour and we did not observe the SPSS too close to the horizon.

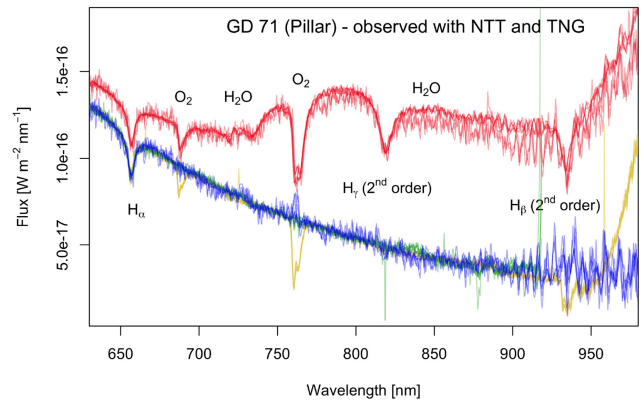


Figure 5. Example of second-order correction for the Pillar GD 71 (SPSS 002). The V2 flux table in the red wavelength range is based on 24 wide-slit spectra observed with TNG and NTT. Of these, the 12 NTT spectra observed with grism #16, affected by a significant second-order contamination, are plotted in red. The corresponding second-order corrected spectra are plotted in blue. The uncorrected TNG spectra, which are unaffected up to ≈ 950 nm, are plotted in gold. The NTT spectra obtained with grism #5, unaffected by reaching only up to ≈ 920 nm, are plotted in green.

of covering the *Gaia* wavelength range with two gratings at most, at the desired resolution, we had to use gratings affected by second-order contamination, i.e. light from the blue side of the second-order contaminating the red side of the first order spectra. This occurred for two of the four spectrographs: EFOC2 and DoLoRes. In the case of DoLoRes, only the reddest 50 nm were affected in the LR-R grism (see Table 2 and Fig. 5). In the case of EFOC2, however, almost the entire range of grism #16 was heavily contaminated, so we added some exposure with grism #5 (Table 2) that covered a smaller wavelength range, but was free from second-order contamination. For the same reason, we observed each SPSS with more than one facility whenever possible. We used the grism #5 spectra to validate our second-order correction procedure. The procedure is described in details in Altavilla et al. (2015) and is based on a method originally proposed by Sánchez-Blázquez et al. (2006). An example of the correction can be found in Fig. 5. The typical residuals of the correction are of 1–2 per cent in flux on the single spectrum and well below 1 per cent on the final flux tables. The sigma spectrum of each affected star was worsened accordingly to take into account the related uncertainty.

Finally, the removal of telluric absorption features was done using two different methods. The SPSS V1 spectra were corrected with our in-house procedure based on the IRAF telluric task. We used theoretical atmospheric absorption spectra from the HITRAN (Rothman et al. 2009) or SPECTRA (Mikhailenko, Babikov & Golovko 2005) data bases⁹ for H₂O molecules and a library of our observed SPSS spectra for O₂ ones, and explored the libraries until the residuals of the correction were minimized. New spectra added in the SPSS V2 version were instead corrected using Molecfit, which is based on similar data sources to build theoretical atmospheric absorption spectra (Kausch et al. 2015; Smette et al. 2015). Molecfit is more practical to use and we verified that the two methods produce comparable results. The typical residuals on the single spectrum after the telluric correction were smaller than 2 per cent in the weakest telluric bands, and could reach up to 5 per cent or more in the strongest ones. However, in the final merging of all the spectra for each given

⁹Accessed through: <http://spectra.iao.ru/>

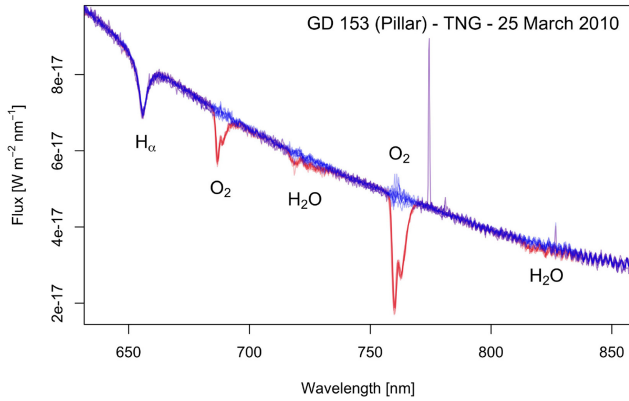


Figure 6. Example of telluric correction for the Pillar GD 153 (SPSS 003), observed repeatedly with DoLoRes at the TNG on 2010 March 25. Nine uncorrected spectra are plotted in red, the corresponding nine corrected spectra are plotted in blue. The $H\alpha$ line is marked as well as the main telluric absorption bands in the plotted wavelength range.

Table 3. List of the reference flux tables of pillars and primary SPSS used to calibrate the V1 and V2 releases.

SPSS	Name	Source	Version
001	G 191-B2B	CALSPEC	g191b2b_mod.007.fits
002	GD 71	CALSPEC	gd71_mod.007.fits
003	GD 153	CALSPEC	gd153_mod.007.fits
007	HZ 2	CALSPEC	hz2.005.fits
009	LTT 2415	Hamuy et al. (1994)	
010	GD 108	CALSPEC	gd108.005.fits
011	Feige 34	CALSPEC	feige34_stis.001.fits
013	Feige 66	CALSPEC	feige66.002.fits
015	GRW + 705824	CALSPEC	grw_70d5824_stisnic.003.fits
018*	BD + 284211	CALSPEC	bd_28d4211_stis.001.fits
020*	BD + 174708	CALSPEC	bd_17d4708_stisnic.003.fits
023	Feige 110	CALSPEC	feige110_stisnic.003.fits
036	1805292	CALSPEC	1805292_nic.002.fits
037	1812095	CALSPEC	1812095_nic.002.fits
043	P 177 D	CALSPEC	p177d_stisnic.003.fits
046	KF 06 T1	CALSPEC	kf06t1_nic.001.fits

Note. *These two primary SPSS were used as calibrators in V1. Later, they were found to be likely variable by Marinoni et al. (2016). Therefore, in V2 we did not use spectra from nights calibrated with these SPSS to build the median reference spectrum of each SPSS (see Section 4). However, once recalibrated, we used the spectra to build the final flux tables.

SPSS, the telluric residuals were reduced to less than 1–2 per cent in all but a handful of cases (see Fig. 6 for an example). The sigma spectrum was thus worsened in the telluric band regions accordingly.

4 SPSS FLUX CALIBRATION

The SPSS reference system is tied to the CALSPEC reference system (Bohlin 2014; Bohlin, Deustua & de Rosa 2019), because we used as Pillars of our flux calibration (see Pancino et al. 2012, for details) the three pure hydrogen white dwarfs (WD) G191-B2B, GD 71, and GD 153 (Bohlin, Colina & Finley 1995), which are directly calibrated on Vega¹⁰ and on Sirius (Bohlin 2014). More precisely, the SPSS V1 and SPSS V2 versions are calibrated on the CALSPEC system as it was in 2013 (see also Section 5.2.1). In practice, we observed at least one of the three CALSPEC pillars each night. When they were

¹⁰A fourth star, HZ 43, was originally considered but later discarded because of a companion at 3 arcsec (Bohlin, Dickinson & Calzetti 2001), that could disturb ground-based observations.

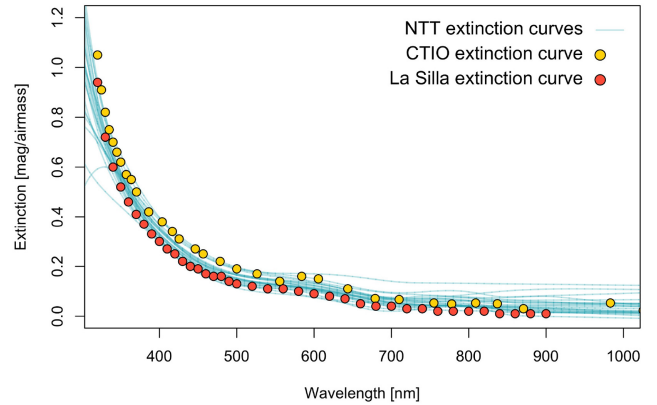


Figure 7. The extinction curves obtained during 25 NTT observing nights with varying atmospheric conditions (cyan lines), compared with the Cerro Tololo (CTIO) reference curve (Stone & Baldwin 1983; Baldwin & Stone 1984, yellow points) and with the ESO La Silla extinction curve (Schwartz & Melnik 1993, red points).

not visible, especially from the Southern hemisphere, we observed a list of well-behaved and well characterized stars from the CALSPEC data base and from the literature as our calibrators, the so-called Primary SPSS (see Pancino et al. 2012, for details). The reference flux tables used to calibrate the SPSS V1 and V2 sets are listed in Table 3.

4.1 Night calibration and quality assessment

To correct for the atmospheric extinction, we computed extinction curves, employing all the observations of stars that were repeated with a minimum time difference of 1 h and an airmass difference of at least 0.2. We monitored the nightly stability of the curves and we divided the nights in three categories: photometric when there were no significant slopes or shape variations among the different curves obtained during the night and all the curves were within 1 per cent from each other¹¹; grey when there were no slope or shape variations, but the curves differed from each other by more than 1 per cent; and non-photometric when there were slopes or shape variations, or when the curves differed by more than 3 per cent, or when the available data did not span the minimum time and airmass ranges of 4 h and 0.2, respectively. The final extinction curves were computed as the median of the curves obtained for the night and an example is shown in Fig. 7, where we compare our NTT curves, obtained in varying atmospheric conditions, with literature curves for La Silla and Cerro Tololo. The comparison is satisfactory within 400–800 nm approximately, but there is still room for improvement at the two extremes of the wavelength range.

We used the spectra obtained for the night calibrators, either Pillars or Primary SPSS (Table 3), to compute the response curves for the night calibration. We generally observed from one to three different calibrators, at different airmass during the night. The night response curve was obtained as the median of the various curves obtained from each observation of the calibrators and the uncertainty was obtained using the median absolute deviation (MAD). We note here that the final SPSS version will be calibrated in a two-step

¹¹Given the quality degradation at the spectra blue and red borders, the QC in this and in the following calibration steps was performed always in the 400–800 nm range.

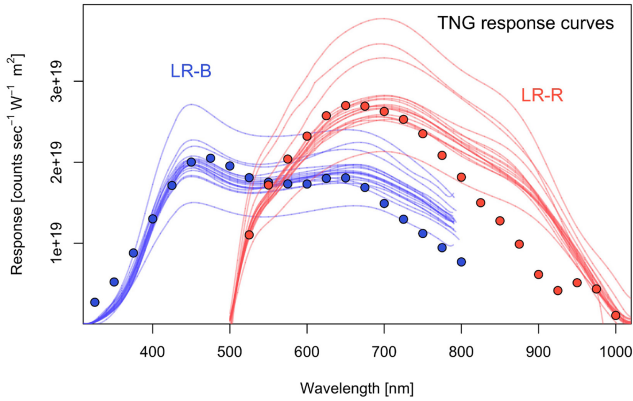


Figure 8. The red and blue response curves (red and blue lines) obtained during 19 TNG observing nights, from 2008 to 2010, with varying atmospheric conditions. The LR-R and LR-B curves obtained by the TNG staff in 2005 (with a different CCD) are shown for reference (red and blue dots, rescaled).

procedure, in which the primary SPSS will be calibrated on the pillars before being used to calibrate all the other (i.e. the secondary) SPSS (Pancino et al. 2012). In the V1 and V2 versions of the SPSS library presented here, we calibrated each night directly on the best available night calibrators, thus the internal consistency is expected to improve slightly in the next SPSS release. For the calibration presented here, we used the literature reference flux tables listed in Table 3 to compute the response curves. The curves obtained for TNG in several observing nights with different atmospheric conditions are shown in Fig. 8.

After the spectra observed in each night were calibrated with the extinction and response curves computed as described above, we further checked that repeated wide-slit exposures of the same SPSS during the night were compatible with each other within 1 per cent and did not display any significant shape variation. A warning was issued in case of zero-point variations up to 3 per cent, and the spectra were discarded in case of higher zero-point variations and/or significant shape variations.

4.2 Absolute calibration and flux tables creation

To select the spectra entering the two preliminary SPSS V1 and V2 releases, we used different criteria. For V1, we selected all the nights that were judged to be photometric according to the criteria described in the previous section. When V1 was internally released, the spectrophotometric campaign was not completed yet, thus we could count on 15 observing nights. For V2, which occurred after the end of the observing campaigns, we could count on 26 nights, using the same criterium. As an exception, we included a few nights that were judged grey or even non-photometric to include a few more red stars, that were necessary to better keep under control the calibration performance at extreme colours for the *Gaia* DR2 and (E)DR3 data.¹² The final flux tables of the few SPSS that did not contain at least one observation sequence in a photometric night were issued with a warning.

¹²The actual *Gaia* external calibration procedure is quite complex and the details can be found in Evans et al. (2018) for DR2, in Riello et al. (2020) for EDR3, and in the forthcoming DR3 release papers (De Angeli et al. and Montegriffo et al., in preparation).

Once the data were selected, we compared all the available spectra for each SPSS with each other. We first selected only the wide-slit spectra of a given SPSS that were observed in photometric nights and rejected those that were not within 1 per cent from the median spectrum, paying special attention to the region of overlap of the red and blue grisms. We then computed the median spectrum with its MAD. In doing so, we slightly adjusted the wavelength calibration of all spectra to align them to the median spectrum, using cross-correlation techniques. We then used the median of the best spectra to adjust the flux calibration of all the remaining spectra. In this way, only the best photometric data contributed to the accuracy of the final flux table, while all the spectra contributed to smooth out the pixel-to-pixel variations caused by noise and by the residuals of the telluric, fringing, and second-order corrections (when relevant). At this point, we recomputed the final median spectrum using all the spectra, the best ones plus all the re-calibrated remaining ones, to create the final SPSS flux table. In a few cases, we observed that some spectra showed a significant slope when compared to the median of the photometric spectra: we included them only when the slope could be successfully removed (within 1 per cent of maximum deviation). In future SPSS releases we will use our absolutely calibrated photometry (Altavilla et al. 2021) to validate and – if necessary and appropriate – to re-adjust the calibration of spectra taken in non-perfectly photometric conditions (see also Section 5.3).

4.3 Extension with template spectra

As mentioned, the sensitivity of the used telescope and instrument combinations drops at the borders of the spectrum. Most of the used grisms do not reach 1050 nm (Table 2) and when they do the S/N ratio is much lower than in the central part of the spectrum. Similar considerations apply to the blue side. For this reason, all the QC and validation procedures are limited to the 400–800 nm range. Some of these limitations are mitigated when combining several different spectra to obtain an SPSS flux table. Nevertheless, it was necessary to extend the spectra to cover the full nominal *Gaia* wavelength range: the passbands cover the 330–1050 nm range approximately, but the *Gaia* BP and RP spectra formally start at 300 nm and reach 1200 nm, although there is very little flux before 330 nm or after 1050 nm. Therefore, we extended the SPSS flux tables to cover the 300–1100 nm range. We note that *Gaia* does have limited sensitivity at the extremes of the spectral range, thus any imperfection in the spectra extension of a given SPSS will have a limited impact on the final accuracy of the *Gaia* flux calibration.

We followed a procedure similar to the one adopted by the CALSPEC team: we extended our observed spectra using template spectra (either observed spectra or theoretical models). In our case however, we did not use the templates to calibrate the spectra, but we adjusted the templates on to the absolutely calibrated spectra. For the stars in common with the CALSPEC library, we extended the SPSS spectra by adjusting the corresponding observed CALSPEC spectra (≤ 1 per cent adjustments). For the other SPSS, we used a series of different observed, theoretical, or semi-empirical libraries¹³ depending on the spectral type of the SPSS (Stritzinger et al. 2005; Koester 2010; Palacios et al. 2010; Sordo et al. 2011; Bailer-Jones et al. 2013; Husser et al. 2013; Coelho 2014; Levenhagen et al.

¹³Most of the spectra were retrieved from the collection at <http://svo2.cab.inta-csic.es/theory/>.

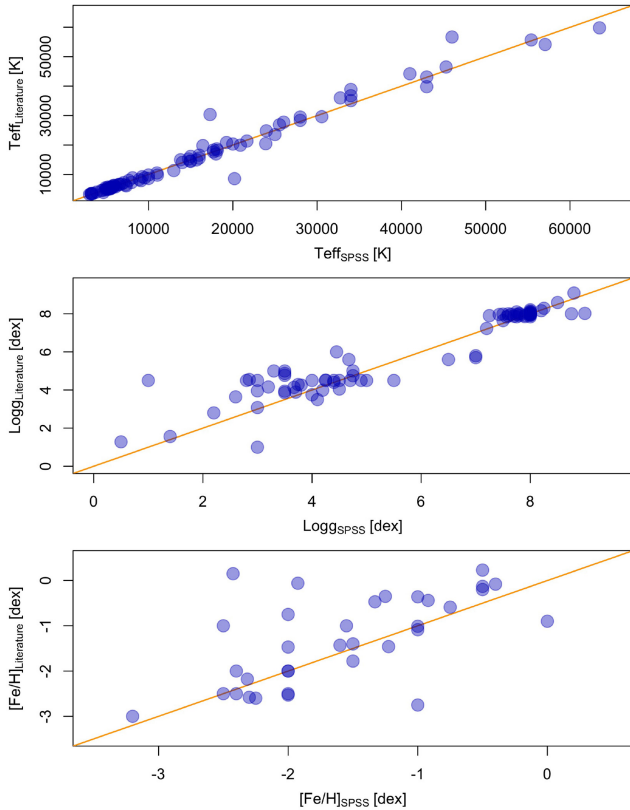


Figure 9. Comparison of the parameters of the best-fitting template spectra (abscissae in all panels) with the available literature determinations (ordinates). The top panel shows the effective temperature, the middle one the surface gravity, and the bottom one the metallicity. The 1:1 line is plotted in orange in all panels (see Section 4.3 for more details).

2017). We used a χ^2 minimization algorithm to choose the best-fitting model for each SPSS. During the comparison, the template spectra were aligned in wavelength on to the observed ones with a cross-correlation. Whenever literature parameters were available, we used them as a first guess to reduce the parameter space to explore, otherwise we used the literature spectral type (Table 1). We found a relatively good agreement between our best-fitting parameters and the literature ones, except for metallicity which is typically poorly constrained for hot stars. The typical median differences and MADs we found are $\Delta T_{\text{eff}} = -60 \pm 697$ K, $\Delta \log g = -0.14 \pm 0.33$ dex, and $\Delta [\text{Fe}/\text{H}] = -0.17 \pm 0.56$ dex (see also Fig. 9). The results of the procedure for the stars that were extended with theoretical models can be found in Table 4, where we list the best-fitting parameters along with the source of the theoretical spectra.

5 SPSS RESULTS AND VALIDATION

The final SPSS V1 and V2 libraries contain respectively 94 and 112 stars. The flux tables are available online (see Data Availability Section) as ascii files containing the wavelength resampled to 0.1 nm and the flux with its uncertainty expressed in $\text{W m}^{-2} \text{nm}^{-1}$. Seven of the SPSS in V1 were later rejected because of variability (Marinoni et al. 2016), companions, or problems with the spectra, and thus do not appear in V2. On the other hand, 25 new SPSS were added in V2. As a result, V2 contains more SPSS than V1 and on average more spectra are used to create the flux tables of each SPSS than in V1 (see also Fig. 2). In any case, the two versions of the SPSS library are on

the same spectrophotometric flux scale within ≈ 0.1 – 0.2 per cent, as detailed in Table 5.

The current list of SPSS V2 surviving candidates, along with the relevant information, can be found in Table 1. Fig. 10 shows the colour–magnitude diagram of the SPSS V1 and V2 sets, and is built using *Gaia* EDR3 parallaxes and photometry. As discussed in Section 2.1, the SPSS grid does not contain extreme spectral types: only a handful of bright and blue stars, carefully selected to be constant and single, are included, while no star cooler than about 3000 K is included. However, compared to the current CALSPEC selection, the SPSS sample extends the WD sequence significantly towards the cool end, it fills the CALSPEC gap in the M dwarfs ($1 \lesssim G_{\text{BP}} - G_{\text{RP}} \lesssim 3$ mag), and it considerably increases the samples of hot subdwarfs and FGK dwarfs and giants. In the following, we present some validation tests on the SPSS V2 sample.

5.1 SPSS list pruning with *Gaia*

As mentioned, we carried out a short-term constancy monitoring (1–2 h) on all SPSS that needed to be tested (Marinoni et al. 2016). We also started a long-term monitoring campaign (1–3 yr), but we could not conclude it for lack of resources. Up to *Gaia* DR2, we relied mostly on literature information to remove long-term variables from the SPSS list. However, *Gaia* is expected, especially in the forthcoming releases, to be extremely efficient in the detection of long-term flux variations above 1 per cent (Mowlavi et al. 2018). So far, we searched for our SPSS candidates in the *Gaia* DR2 variable stars catalogue (Holl et al. 2018) and we did not find any. With future *Gaia* releases, we will update our search to exclude all known variable stars. We also looked into the main EDR3 catalogue for companions to our SPSS V2 stars and found that 10 indeed have one, either farther than ≈ 2 arcsec (our typical seeing is ≈ 1 arcsec), or fainter than 5 mag (i.e. less than 1 per cent flux contamination), or both. The only exception (but see below) is SPSS 215, which appears to have a companion in the main *Gaia* EDR3 catalogue at about 0.7 arcsec, with a magnitude difference of about $\Delta G = 3$ mag. This means that the ground-based spectra are certainly disturbed by the companion.

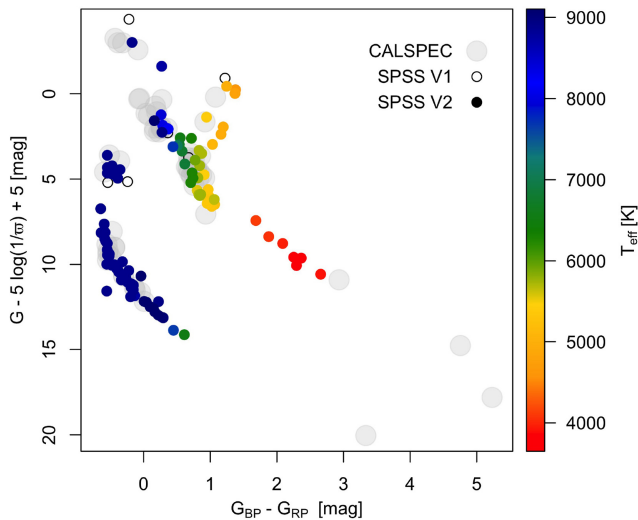
To explore the problem further, we used the Source Environment Analysis (SEAPipe), which had its first runs in the *Gaia* processing in preparation of DR3. The aim of SEAPipe is to combine the transit data for each source and to identify any additional sources in the local vicinity. The first operation in SEAPipe is image reconstruction, where a 2D image is formed from the mostly 1D transit data ($G > 13$ mag). The algorithm used to perform the image reconstructions is described in Harrison (2011). As part of the validation of SEAPipe all images reconstructed for SPSS sources were visibly inspected. Not all SPSS sources had reconstructed images, as there must be a good coverage of transits crossing the source in different orientations for a successful image reconstruction, this in fact is the main reason why SEAPipe was not run in previous releases. SEAPipe was attempted on all 119 SPSS sources (the union of V1 and V2) and successfully produced images for 116 of these sources. The remaining three sources (SPSS 20, 28, and 32) did not have enough suitable transits for the image reconstruction to proceed. The inspection of these 116 images revealed 114 SPSS sources with no evidence of a companion within 1 arcsec at least to the level of the noise in their reconstructed images. However, it was discovered that two of the sources, SPSS 192 and 215, had companions within 1 arcsec. These images are shown in Fig. 11. The angular separation (0.29 arcsec, 0.69 arcsec) and magnitude difference (1.9, 3.0 mag) of the companion from the SPSS source (192, 215) were found by performing an image parameter

Table 4. Parameters of the best-fitting models used to extend the SPSS flux tables to cover the full 300–1100 nm *Gaia* range. Stars extended with observed spectra are not listed here. Only the first five rows are shown, the table is available in its entirety online (see Data Availability Section).

ID	Name	SpType	T_{eff} (K)	$\log g$ (dex)	[Fe/H] (dex)	A_V (mag)	Library	Notes
005	EG 21	DA3	16000	8.0	0.0	0.0	Koester (2010)	Template agreement at junctions 1–3 per cent
006	GD 50	DA2	43000	8.8	0.0	0.0	Sordo et al. (2011)	Template agreement at junctions 1–3 per cent
008	LTT 3218	DA	9250	7.75	0.0	0.0	Koester (2010)	–
009	LTT 2415	G	23900	7.0	0.0	0.045	Sordo et al. (2011)	Discrepant spectral type
010	GD 108	B	26030	4.675	0.0	0.1210	Sordo et al. (2011)	Template agreement at junctions >3 per cent

Table 5. Median ratio and MAD of the SPSS V1 and V2 libraries and of the SPSS and the CALSPEC libraries (2021 version), for the 11 stars in common (see Fig. 12).

Ratio	300–400 nm	400–800 nm	800–1070 nm
V2/V1	1.0010 ± 0.0025	1.0001 ± 0.0011	1.0018 ± 0.0034
V1/CALSPEC	0.9921 ± 0.0207	1.0016 ± 0.0056	1.0012 ± 0.0152
V2/CALSPEC	0.9891 ± 0.0101	0.9995 ± 0.0089	1.0052 ± 0.0115

**Figure 10.** The *Gaia* colour–magnitude diagram of the SPSS. The CALSPEC stars are plotted in the background as grey circles for comparison. The SPSS V2 stars are coloured as a function of their T_{eff} (the colour scale saturates at $T_{\text{eff}} > 9000$ K, for readability). Seven SPSS V1 stars that were not included in SPSS V2 are plotted as empty black circles.

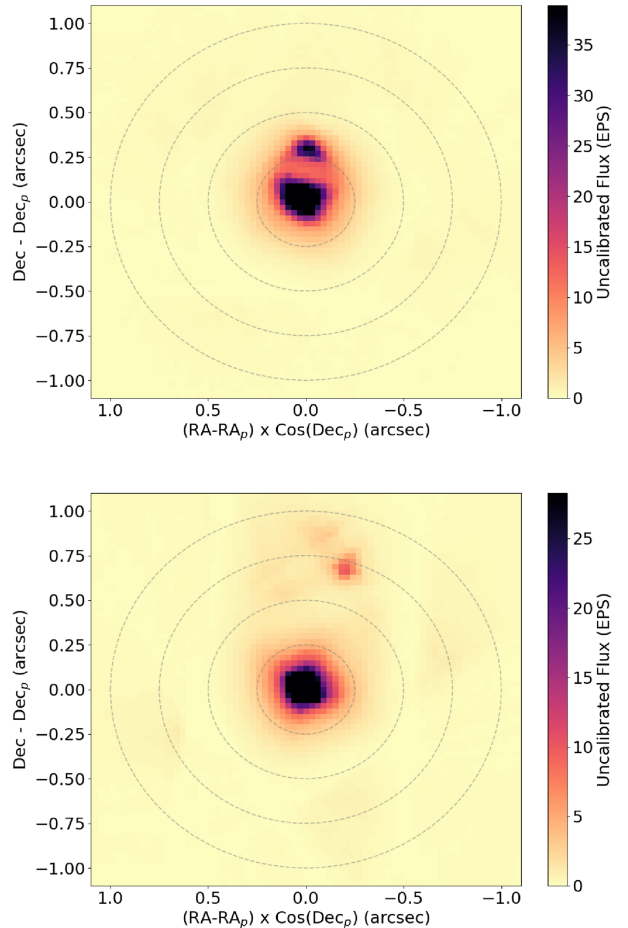
analysis on the data used to construct the images. For SPSS 215 these are in agreement with those of the companion found in EDR3.

We decided to leave SPSS 192 and 215 in the list because they were actually used in the calibration procedure for EDR3, but they will be rejected from the next SPSS release. They were annotated in Table 1 and assigned a quality flag of two.

5.2 SPSS literature comparisons

5.2.1 CALSPEC spectra

As mentioned in Section 4, the SPSS flux reference system is tied to the CALSPEC system as it was in 2013. After that, the CALSPEC system was adjusted by making it 0.6 per cent fainter on average (Bohlin 2014). More recently, the CALSPEC system was re-adjusted again by making it 0.87 per cent brighter (Bohlin, Hubeny & Rauch 2020). As a result, we expect our SPSS grid to

**Figure 11.** Images reconstructed as part of the SEAPipe analysis for two SPSS sources found to have visible companions. Images are reconstructed centred on the location of the primary source (RA_p , Dec_p), where the primary source refers to the *Gaia* source whose transit data are used in the reconstruction. The grey dashed circles indicate angular separations of 0.25, 0.50, 0.75, and 1.0 arcsec from the centre of the image (and primary source position). The colour scale of these images was set ignoring the image pixels within 0.1 arcsec of the centre, in order to better visualize the companions, hence the saturated appearance of the primary sources. Top panel: The reconstruction for SPSS 192, showing a companion at an angular separation of $\simeq 0.29$ arcsec and a magnitude difference of $\simeq 1.9$ mag. This companion is not present in EDR3. Bottom panel: The reconstruction for SPSS 215, showing the companion present in EDR3 at an angular separation of $\simeq 0.69$ arcsec and a magnitude difference of $\simeq 3.0$ mag.

agree, within uncertainties, with the latest CALSPEC release. We thus compared the SPSS V1 and V2 data sets with the latest version of the CALSPEC library, with the results reported in Table 5 and in Fig. 12. We note that in the bluest range (300–400 nm) the SPSS

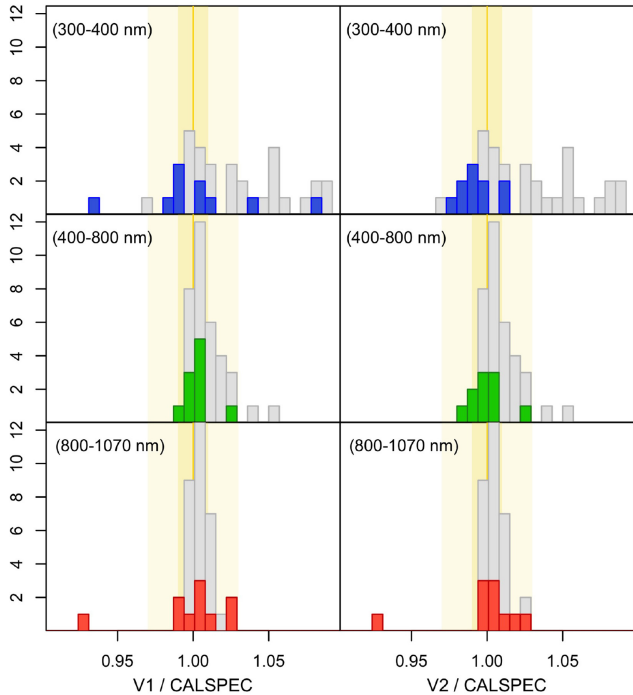


Figure 12. Ratio of the SPSS V1 (left-hand panels) and V2 (right-hand panels) flux tables with the CALSPEC ones (Bohlin et al. 2019), for the 11 stars in common. The top panels show the median ratio in the 300–400 nm range (in blue), the middle panels in the 400–800 nm range (in green), and the bottom ones in the 800–1070 nm range (in red), as labelled. The ratio of the most recent releases of the CALSPEC spectra of the same star, when available, are plotted in grey, to give an idea of the CALSPEC internal uncertainty over time. Perfect agreement is marked by an orange vertical line, and the ± 1 per cent and ± 3 per cent agreement ranges are highlighted in shades of orange in all panels (see Table 5).

spectra are about 1 per cent fainter than the CALSPEC ones. We can also note that the SPSS V2 release shows a smaller scatter in the blue than the V1 release when compared with the same set of CALSPEC spectra (Table 5 and Fig. 12), thus the precision of the blue side of the SPSS spectra has improved between the two releases, likely because of the increased number of spectra used. In the central range (400–800 nm), where the SPSS spectra have the highest S/N ratio, the SPSS and CALSPEC fluxes agree to better than 0.5 per cent, as expected. We note that V1 had a slightly better agreement with CALSPEC than V2, especially at longer wavelength, albeit with a larger scatter. In preparation for the next SPSS release, we are carefully revising all the reduction and calibration steps to understand the exact cause for the change. This slight worsening of the median ratio above 800 nm might in fact be related to the $\simeq 1$ per cent offset observed by Riello et al. (2020, fig. 25) for faint blue stars. We can finally note from Fig. 12 that the latest releases of the spectra in the CALSPEC library for the same star (plotted in grey in the background) differ from each other by amounts that are comparable to the scatter in the comparison between the SPSS and CALSPEC libraries.

5.2.2 Gaia EDR3 magnitudes

As a basic sanity check, we computed *Gaia* synthetic magnitudes on the SPSS V2 spectra and compared them with the corresponding *Gaia* EDR3 integrated magnitudes (Riello et al. 2020). We used the EDR3

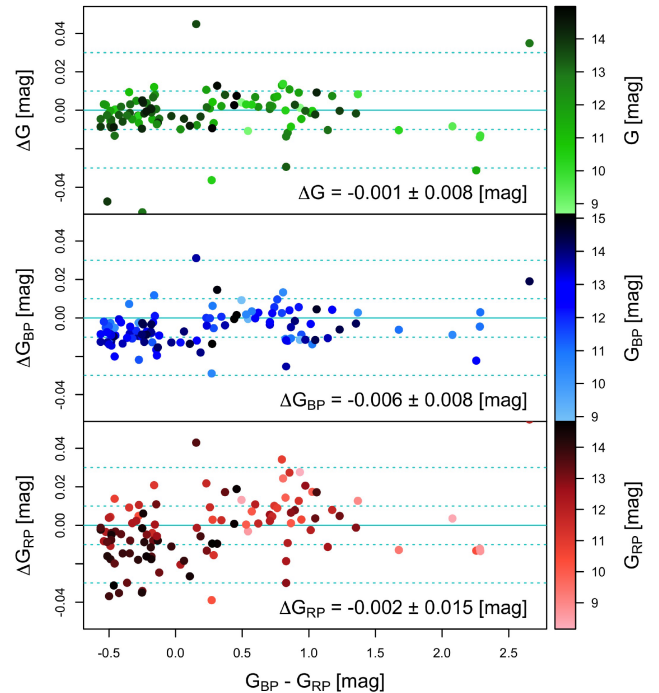


Figure 13. Difference between the synthetic magnitudes computed on the SPSS V2 flux tables (Section 5.2) and the *Gaia* EDR3 integrated magnitudes, as a function of the $G_{BP}-G_{RP}$ colour. The SPSS are coloured as a function of their *Gaia* EDR3 observed G magnitude, as indicated in the legend. The top panel shows the case of G , the middle one of G_{BP} , and the bottom one of G_{RP} . Perfect agreement is indicated by a solid cyan line in all panels, while the ± 1 and ± 3 per cent agreements are indicated by dotted lines.

passbands¹⁴ and the same Vega spectrum used by Riello et al. (2020), for consistency. The results of the comparison are shown in Fig. 13 and the synthetic magnitudes are presented in Table 6. A number of outliers is clearly present. These will be individually examined once the full set of spectra will be reduced, to identify and solve any obvious problem in the reduction chain. The median offsets and median absolute deviations, computed in the sense of SPSS synthetic magnitudes minus *Gaia* EDR3 ones, are: $\Delta G = -0.001 \pm 0.008$, $\Delta G_{BP} = -0.006 \pm 0.008$, and $\Delta G_{RP} = -0.002 \pm 0.015$ mag. An excellent agreement is found, with a slight negative bias of the order of 0.1–0.6 per cent, i.e. negligible compared to the uncertainties. A larger spread is observed in the comparison with G_{RP} , with a possible trend with magnitude among blue stars. This might in part be related to the above mentioned 1 per cent differences found between SPSS V1 and V2 redwards of 800 nm, in the sense that some extra uncertainty appears to be present in the SPSS flux tables of faint blue stars (darker symbols in the bottom panel of Fig. 13). As mentioned, this effect was also noted by Riello et al. (2020, fig. 25). If *Gaia* magnitudes were significantly affected by this minor SPSS defect, we would observe a better agreement in the bottom panel of Fig. 13. We also note some jumps at about $G_{BP}-G_{RP}$ that are not present in the comparisons with CALSPEC or the ground-based SPSS photometry.

¹⁴<https://www.cosmos.esa.int/web/gaia/edr3-passbands>

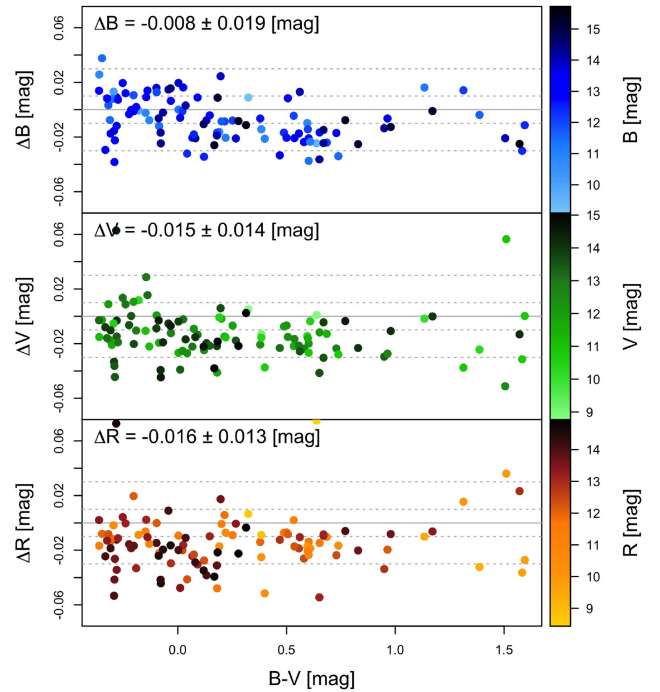
Table 6. Synthetic magnitudes in the Johnson–Kron–Cousins and the *Gaia* EDR3 systems, computed on the SPSS V2 flux tables (see Section 5.2 for details). Only the first five rows are shown here, the full table is available online (see Data Availability Section).

SPSS	<i>U</i> (mag)	<i>B</i> (mag)	<i>V</i> (mag)	<i>R</i> (mag)	<i>I</i> (mag)	<i>G</i> (mag)	<i>G</i> _{BP} (mag)	<i>G</i> _{BP} (mag)
1	10.247 ± 0.005	11.482 ± 0.006	11.790 ± 0.008	11.931 ± 0.012	12.112 ± 0.021	11.721 ± 0.009	11.541 ± 0.006	12.067 ± 0.019
2	11.667 ± 0.005	12.804 ± 0.010	13.041 ± 0.013	13.175 ± 0.016	13.338 ± 0.031	13.000 ± 0.014	12.844 ± 0.010	13.302 ± 0.026
3	11.908 ± 0.004	13.088 ± 0.008	13.355 ± 0.011	13.495 ± 0.011	13.671 ± 0.021	13.304 ± 0.011	13.136 ± 0.008	13.628 ± 0.019
5	10.689 ± 0.011	11.403 ± 0.008	11.373 ± 0.006	11.459 ± 0.012	11.527 ± 0.018	11.405 ± 0.010	11.345 ± 0.008	11.528 ± 0.018
6	12.583 ± 0.007	13.798 ± 0.010	14.067 ± 0.012	14.206 ± 0.019	14.369 ± 0.037	14.013 ± 0.015	13.845 ± 0.010	14.335 ± 0.031

5.2.3 Johnson–Kron–Cousins magnitudes

We also computed the synthetic Johnson–Kron–Cousins UBVRI magnitudes on the SPSS V2 set and compared them with our own absolutely calibrated magnitudes (Altavilla et al. 2021) and, for 36 stars, with the original measurements in a series of papers by Landolt and collaborators (Landolt 1983, 1992b; Landolt & Uomoto 2007; Landolt 2009; Bohlin & Landolt 2015). To compute the Johnson–Kron–Cousins magnitudes, we used the passbands by Bessell & Murphy (2012). The choice of the Vega spectrum is crucial, as variations of 1–2 per cent can be easily obtained in the results when using different spectra. We thus employed the same Vega spectrum used for the *Gaia* comparison above (Section 5.2.2), i.e. the Kurucz spectrum of an A0V star, normalized at 550 nm to a reference flux of $3.660 \cdot 10^{-11} \text{ W m}^{-2} \text{ nm}^{-1}$, according to Straižys (1992). If we recompute the magnitude on the Vega spectrum `alpha_lyr_stis_005` used by Bessell & Murphy (2012), we obtain $U = 0.034$, $B = 0.025$, $V = 0.027$, $R = 0.028$, and $I = 0.028$ mag, in good agreement (better than 1 per cent) with their results. This ensures that our synthetic magnitudes, presented in Table 6, are on the standard scale, within uncertainties.

The comparison of the resulting synthetic magnitudes, obtained on the SPSS V2 spectra, with the magnitudes by Altavilla et al. (2021) is shown in Fig. 14. As can be observed, the scatter is larger than in the case of Fig. 13, confirming the excellent quality and homogeneity of the *Gaia* EDR3 photometry. The *R* magnitude comparison hints at a segregation of faint blue stars (darker points in the bottom panel of Fig. 14, as noted in the previous section), which appear to have a slightly lower average difference. The SPSS V2 synthetic magnitudes appear about 1–2 per cent brighter than the Altavilla et al. (2021) ones, which in turn appeared to be about 1.5 per cent fainter than the Landolt measurements (see their fig. 8). This is confirmed when comparing our SPSS V2 synthetic magnitudes with the ones from the Landolt collection in Fig. 15: the SPSS V2 magnitudes agree with the ones of the Landolt collection to better than 1 per cent. Given the involved uncertainties in the magnitude measurements (see Clem & Landolt 2013) and in the synthetic magnitudes computation, the agreement can be considered quite satisfactory. In this case, the *R* magnitude comparison appears excellent for the faint blue stars, while some problems appear in the *I* band comparison, where the scatter is large and again the faint blue stars appear to behave slightly differently from the other SPSS. This suggests that the problems in the SPSS V2 spectra of faint blue stars, if any, should start after ≈ 700 – 800 nm. We also note the large scatter in the *U* band comparison: the *U* band is notoriously difficult to standardize (Stetson 2000; Bessell 2005; Stetson et al. 2019). This sums up with the lower S/N ratio typically reached in the SPSS observed spectra below ≈ 400 nm. As shown in the comparison between SPSS V1 and V2 (Section 5.2.1), increasing the number of spectra can significantly improve the scatter in the blue. Further improvement in this sense is expected in the next SPSS release.

**Figure 14.** As Fig. 13, but comparing with the Johnson–Kron–Cousins magnitudes of the SPSS by Altavilla et al. (2021). The top panel refers to the *B* magnitude, the middle one to *V*, and the bottom one to *R*, as labelled.

5.3 Expected improvements in the next SPSS release

As already mentioned, the current SPSS release is still preliminary, although it already meets the original requirements, showing an excellent agreement with both the CALSPEC and Landolt reference systems. The next release, SPSS V3, that will be used to calibrate *Gaia* DR4, is expected to further improve based on the following considerations:

- (i) the entire set of about 6500 spectra will be analysed, increasing the S/N ratio and smoothing out further pixel-to-pixel defects and residuals of the various advanced spectra reductions (section 3.2);
- (ii) the entire set of surviving SPSS candidates, which at the moment are about 200, will be included; the list of candidates will be further cleaned from variable and multiple stars (Section 5.1);
- (iii) the source of the ≈ 1 per cent offset observed on the red part of the spectrum ($\gtrsim 700$ – 800 nm) for faint blue SPSS (Section 5.2) will be investigated and repaired if possible;
- (iv) the calibration process used in V1 and V2 was tied to the 2013 version of the CALSPEC grid (Section 4 and Table 3), while in V3 we plan to calibrate on to the latest CALSPEC spectra of the three pillars

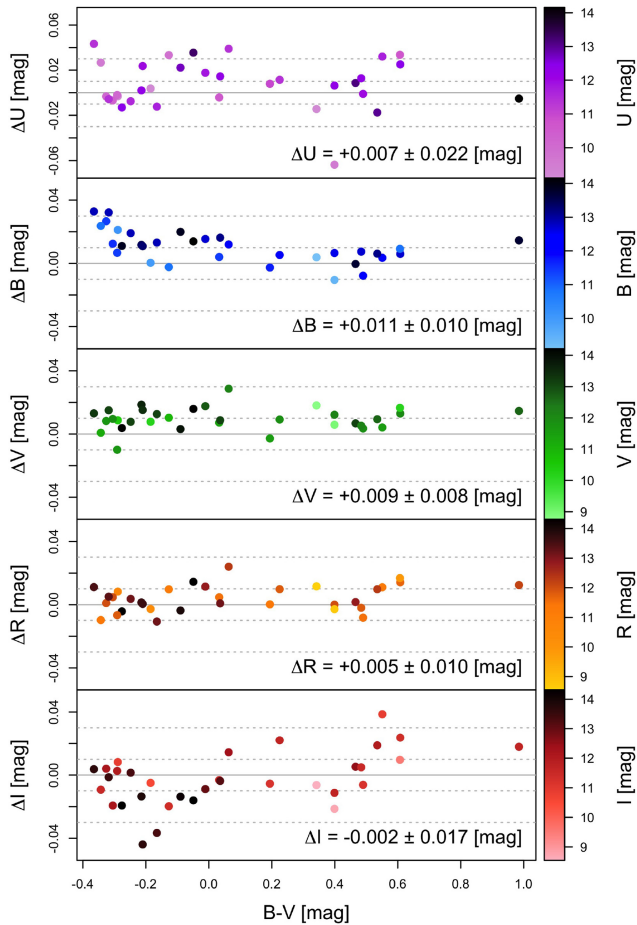


Figure 15. As Figs 13 and 14, but comparing with the Johnson–Kron–Cousins magnitudes from a series of papers by Landolt and collaborators (see Section 5.2). Each panel refers to a different passband, as labelled.

only, following the Pillar–Primary–Secondary calibration scheme (see also Section 4 Pancino et al. 2012);

(v) the extension of the observed flux tables with template spectra to replace the noisy borders will be updated with new template spectra and more refined algorithms, to minimize discrepancies at the junctions (ideally within 1 per cent or less).

All these improvements are not expected to significantly change the overall accuracy (i.e. the zeropoint) of the SPSS grid, but are expected to improve its precision (i.e. the homogeneity). In fact, the accuracy of the SPSS V2 preliminary release already meets the initial requirements set within the *Gaia* DPAC and is of the order of 1 per cent or better, which is considered the current technological limit of flux calibrations (Clem & Landolt 2013; Bohlin 2014; Altavilla et al. 2021). The precision, on the other hand, can be improved both within SPSS (flux tables precision) and among SPSS (grid homogeneity). Within SPSS, by increasing the S/N ratio and by reconciling the blue and red edges of each flux table with the high-quality central parts. Among SPSS, by calibrating more homogeneously and self-consistently each SPSS on the three CALSPEC pillars only, instead of the larger collection of calibrators listed in Table 3, and by investigating the cause of the ≈ 1 per cent offset in the red part of the flux tables ($\gtrsim 700$ – 800 nm) of the blue and faint stars.

6 THE PASSBAND VALIDATION LIBRARY (PVL)

After the publication of the *Gaia* DR2 photometry (Evans et al. 2018), a variety of passbands and tests on the quality of the integrated G, G_{BP} , and G_{RP} magnitudes appeared in the literature (Casagrande & Vandenberg 2018; Evans et al. 2018; Maíz Apellániz & Weiler 2018; Weiler 2018). In particular, given the complexity of the multiparametric model for the (internal and external) flux calibration of *Gaia* data, it was suggested that the lack of variety in stellar spectra in general – and in the SPSS set in particular – might be the cause of some small systematics observed in the *Gaia* DR2 photometry (Maíz Apellániz & Weiler 2018) and in the differences among some of the published passbands. A library of well calibrated spectra covering also extreme spectral types, that do not match our criteria as flux calibrators (especially flux constancy, see next sections), was assembled by Maíz Apellániz & Weiler (2018), using mostly *HST* spectra, including a populous sample of O and B stars and three very red stars. Unfortunately, that library of spectra is not publicly available. We thus assembled a separate library, the PVL, using the original data sources cited by Maíz Apellániz & Weiler (2018), whenever possible, as described in the following. The error estimates on the flux were computed by summing in quadrature, when known, the formal flux uncertainties, the flux variability, and additional systematic uncertainties. We made the effort, whenever appropriate, to place the PVL on the same flux reference system of the SPSS (see also Section 5.2) so that the two sets can be used together for a variety of validation tasks.

6.1 PVL stars from CALSPEC

We selected a list of 39 stars from the 2019 CALSPEC version¹⁵ (Bohlin et al. 2019) by excluding: (i) the 11 stars already included in the SPSS V2 set¹⁶; (ii) all the WDs, which are already well represented in the SPSS V2 set; (iii) stars with incomplete STIS coverage; (iv) variable stars; (v) stars with problems related to SNR, resolution, or spectral stitching; (vi) stars that are too bright to be observed accurately by *Gaia* (e.g. Sirius and α Lyr). Four CALSPEC very red stars are treated separately in Section 6.3. The CALSPEC set adds 15 A-type stars, 3 B-type stars, 1 O-type star, 6 F-type stars, 11 G-type stars, and 3 subdwarfs to the SPSS V2 coverage. For each of the 39 selected CALSPEC stars, we selected the latest and most reliable CALSPEC spectrum. We brightened the flux of each spectrum by 0.6 per cent to bring it to the pre-Bohlin (2014) CALSPEC system, on which the SPSS V2 is calibrated,¹⁷ and we combined the statistical and systematic errors in quadrature to provide one single error estimate. The spectra were not extended with models because they cover the entire range 300–1200 nm, except for star HD 60 753, for which the spectrum stops at about 1022 nm, but this is a hot star with very little flux in the red.

¹⁵ftp://ftp.stsci.edu/cdbs/current_calspec

¹⁶In reality, star 1757132 was added to the CALSPEC set in 2015, and we did not realize that it corresponds to SPSS 309 (2MASSJ17571324+6703409), that was added to the SPSS candidates in 2012 as one of the stars close to the North Ecliptic Pole for *Gaia* GDR1 testing. We decided to leave the star in PVL as a comparison object with the SPSS set.

¹⁷All these selections were made in 2019, i.e. before the latest revision of the CALSPEC grid (Bohlin et al. 2019, 2020), thus the spectra are on average calibrated on a system that is 0.6 per cent brighter than the one of the SPSS, based on the 2013 CALSPEC version.

6.2 PVL hot and bright stars

As mentioned, hot and bright stars were only included in small numbers in the SPSS list due to the high fraction of variable and binary stars (De Marco & Izzard 2017; Gaia Collaboration 2019). Therefore, they were important to include in the PVL. The Maíz Apellániz & Weiler (2018) compilation contained hot and bright stars from two separate sources, that they labelled ‘Massa’ and ‘Hot’ star sets. The former set is not publicly available, while the latter was published by Khan & Worthey (2018) as a hot star extension of the NGSL (New Generation Spectral Library; Koleva & Vazdekis 2012), containing 41 stars in the post-AGB phase, HB stars, and O-type stars, 17 of which were considered by Maíz Apellániz & Weiler (2018). We thus examined the star list in Khan & Worthey (2018) and decided to: (i) avoid all the planetary nebulae and proto-PN of the set, including one proto-PN that was in Maíz Apellániz & Weiler (2018) and is variable; (ii) discard all WDs, that are well represented in the SPSS V2 set; (iii) remove all variable stars, including three stars that were part of the Maíz Apellániz & Weiler (2018) list. The final list contains 17 stars, of which 14 are in common with the Maíz Apellániz & Weiler (2018) set. The flux calibration of these spectra is not well described in the original paper (Khan & Worthey 2018), but reference is made to NGSL (Koleva & Vazdekis 2012), which is calibrated with an accuracy of about 3 per cent. We thus added an error component of 3 per cent of the flux to compute the final error budget, and we did not apply the (comparatively small) 0.6 per cent flux correction, because it is not clear on which version of the CALSPEC system these stars were calibrated. STIS spectra reach 1025 nm approximately, but these hot stars have very little flux at the red edge and therefore the spectra were not extended with models.

6.3 PVL very red stars

Very red stars are known to be variable because of chromospheric activity, rotation, star spots, magnetic activity, and flares. Brown dwarfs are known to be variable because of clouds complexes in their atmospheres, modulated by rotation and seasonal changes (Artigau 2018). Therefore, when assembling our sample of very red stars, we took into account the known variability found in the literature, and included it in the error computation, as detailed below. We selected from CALSPEC four very red stars: GJ 555, VB 8, and two brown dwarfs (2M0036+18 and 2M0559–14). The two brown dwarfs have no observations bluewards of 500 nm, because in that range they emit very little flux. Nevertheless, they were detected by *Gaia*. We extended their spectra with the appropriate models from the Allard, Homeier & Freytag (2012) grid. All the spectra were treated like the CALSPEC spectra in Section 6.1: we selected the latest spectrum and brightened to the level of the pre-Bohlin (2014) CALSPEC reference system, on which the SPSS V2 is calibrated. We used the typical flux variations of the respective spectral types (Artigau 2018) as the flux error component caused by possible brown dwarf variability, amounting to 2 per cent for 2M0036+18 and to 1 per cent for 2M0559–14. For VB 8 we used the estimated variability by Martin, Zapatero Osorio & Rebolo (1996) to add a 2 per cent component to the flux uncertainty. For GJ 555, we used the variability estimated by Hosey et al. (2015) of 1.4 per cent.

We further added Proxima Centauri, that was part of the Maíz Apellániz & Weiler (2018) grid. We used the spectrum by Ribas et al. (2017), although the uncertainties in this case are high: (i) the various spectra used by Ribas et al. (2017) agreed with each other within 2–3 per cent; (ii) the sources of flux variability examined by Ribas et al.

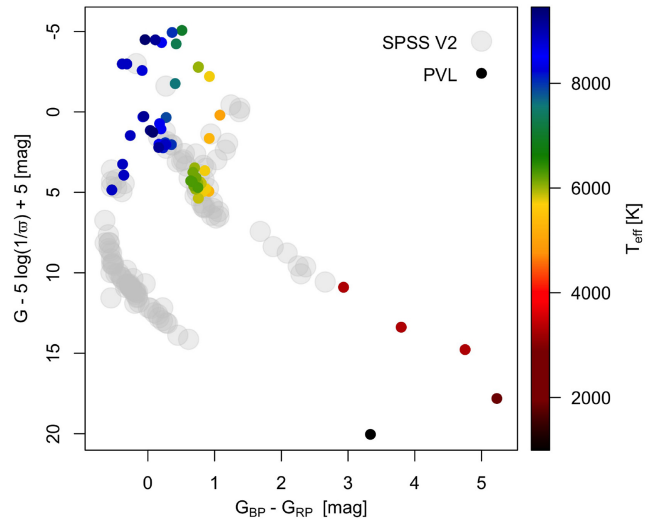


Figure 16. The *Gaia* colour–magnitude diagram of the PVL. The SPSS V2 stars are plotted in the background as grey circles for comparison. The PVL stars are coloured as a function of their T_{eff} (the colour scale saturates at $T_{\text{eff}} > 9000$ K, for readability).

(2017) amount to about 5 per cent; and (iii) the authors used STIS spectra taken with a narrow slit (0.2 arcsec). We thus conservatively estimated the uncertainty of this spectrum to be about 5 per cent. This star should be used with caution also because of a suspect disc of warm polarized dust, but mostly because of flares (MacGregor et al. 2018), and in fact it displays a very large scatter in figs 6 to 9 by Maíz Apellániz & Weiler (2018). We also note that Proxima Centauri was not included in the CALSPEC grid, presumably for the above reasons.

6.4 PVL content and results

The PVL library contains 61 stars. Its content is illustrated in Fig. 16 and the list of stars is presented in Table 7. The flux tables are available online (see Data Availability Section) and the files are in the same format as the SPSS ones. As can be noted, the PVL contains many more hot and bright O and B stars, three more hot subdwarfs, a few FGK stars, and five extremely red stars, covering a range that is not included in the SPSS set (see also Fig. 10). As mentioned, the two sets were calibrated as much as possible on the same flux reference system, so that they could safely be used together in a variety of validation procedures during the preparation of *Gaia* EDR3 (see also Riello et al. 2020).

As an example, the PVL was instrumental in investigating the apparent discrepant behaviour of blue stars found in DR2 by Weiler (2018) and Maíz Apellániz & Weiler (2018). These authors noted that the blue stars fainter than $G \simeq 10.87$ mag behaved differently from those brighter than that limit, and postulated the need for two different sensitivity curves to reconcile the magnitudes of bright and faint blue stars. We used the combined PVL and SPSS V2 global sample to explore the problem from a different point of view. We computed synthetic magnitudes using the same Vega spectrum used in Section 5.2.2 and the REV passbands by Evans et al. (2018). As shown in Fig. 17, while faint blue stars do have consistent synthetic and observed G_{BP} magnitudes, bright blue stars do belong to both the branches of an apparent bifurcation for colours bluer than $G_{\text{BP}} - G_{\text{RP}}$ of about 1 mag. In practice, there are bright blue stars that are well-behaved and other bright blue stars that do not

Table 7. List of stars included in the PVL along with basic information: the PVL ID and name, the *Gaia* EDR3 ID and coordinates (rounded), the literature *B* and *V* magnitudes and the spectral types from the original literature sources of the spectra. Only the first five lines are shown here, the table is available in its entirety online (see Data Availability Section).

PVL ID	Name	<i>Gaia</i> EDR3 ID	RA (deg)	Dec (deg)	<i>B</i> (mag)	<i>V</i> (mag)	Sp. Type	Source	Notes
401	BD−11 3759	6324325225803432320	218.56848	−12.51701	12.9	11.3	M3.5V	CALSPEC	Flare star
402	Prox Cen	5853498713190525696	217.39347	−62.67618	12.9	11.1	M5.5V	Ribas et al. (2017)	Flare star; dust rings
403	VB8	4339417394313320192	253.89336	−8.39840	18.7	16.92	M7V	CALSPEC	Flare star; variable
404	2M0036 + 18	2794735086363871360	9.06713	+18.35286	−	−	L3.5	CALSPEC	Brown dwarf
405	2M0559-14	2997171394834174976	89.82995	−14.08034	−	−	T4.5	CALSPEC	Brown dwarf

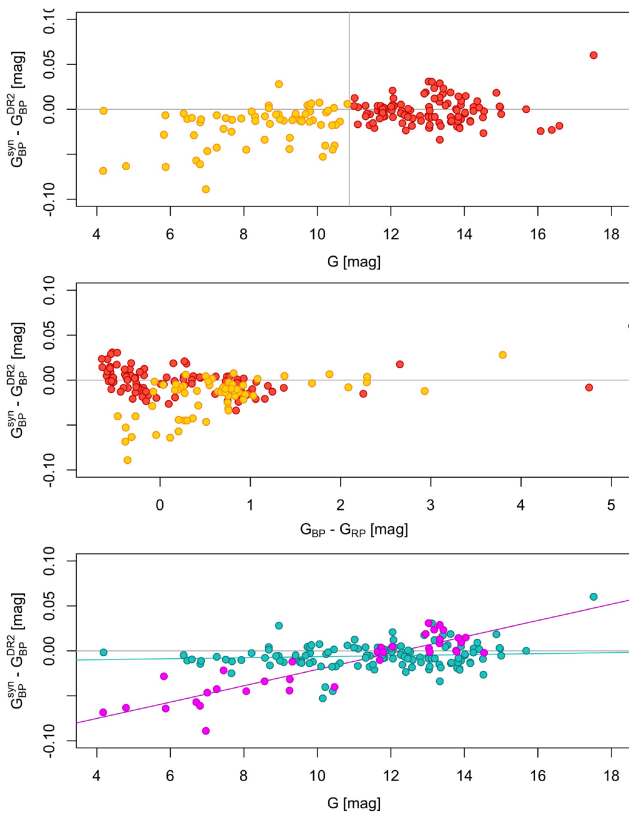


Figure 17. Example of how the SPSS + PVL library was used to investigate the problems of bright blue stars in *Gaia* DR2, that are now much less prominent in EDR3. Top panel. The difference between the G_{BP} synthetic magnitude and the DR2 one for the global SPSS + PVL sample. The cut at $G = 10.87$ proposed by Maíz Apellániz & Weiler (2018) is marked by a vertical line: fainter stars are coloured in red and brighter ones in gold. Middle panel. Same as the top panel, but as a function of $G_{BP} - G_{RP}$ instead of G . Fainter stars (in red) have on average consistent G_{BP} , while bright blue stars (in gold) show a fan-shaped discrepancy at blue colours. Bottom panel. Same as the top panel, but now stars are coloured according to T_{eff} rather than G magnitude. Linear fits to the samples with T_{eff} lower (in cyan) and higher (in magenta) than 30 000 K are also plotted. In all panels, perfect agreement is marked by a horizontal grey line.

behave well (middle panel of Fig. 17).¹⁸ This is true regardless of the exact cut in magnitude: we used values ranging from $G = 10.5$ to 12.5 mag obtaining qualitatively the same result (top panel of

¹⁸This is slightly different from the Maíz Apellániz & Weiler (2018) conclusion that all the bright blue stars behave differently from faint blue ones.

Fig. 17). However, we realized that hot stars with $T_{\text{eff}} > 30\,000$ K lie on the discrepant branch, while hot stars cooler than 30 000 K lie on the well behaved branch. Moreover, faint WDs hotter than 30 000 K appear to lie on the opposite side of the zero line (bottom panel of Fig. 17), suggesting that these extreme spectral types, that are not so numerous, were not well represented in the internal calibration model used for DR2. As can be noted in fig. 5.3 in section 5.1.3 of the *Gaia* online DR2 documentation,¹⁹ the shape of BP spectra changes above $\approx 30\,000$ K in the blue side of the spectra. As a result of this and many other investigations, the systematic effects in *Gaia* EDR3 are now greatly reduced compared to the ones in DR2. Further improvement is expected in future releases, as the amount of data increases as well as our knowledge and understanding of the data properties.

7 DISCUSSION AND CONCLUSIONS

We presented the flux tables for the first two preliminary releases of the *Gaia* SPSS: the V1, used to calibrate *Gaia* DR2 and the V2, used to calibrate *Gaia* (E)DR3. We complemented the SPSS grid with the PVL, a library of stars that do not pass all of our criteria to be flux standards, but which possess well calibrated spectra in the literature. Most importantly, the PVL contains very red and bright blue stars that are normally not used in great quantities as flux calibrators. The SPSS + PVL combined set was used as a validation set during the processing of *Gaia* (E)DR3 photometry.

The SPSS preliminary releases already meet the required accuracy for the *Gaia* project in the flux reference system: they agree to better than 1 percent with the CALSPEC grid for 11 stars in common (Section 5.2.1) and with a collection of 37 stars from a series of papers by Landolt and collaborators (Section 5.2.3). The current SPSS releases do not contain all the observed SPSS spectra and all the surviving SPSS candidates, thus improvements are expected in the upcoming SPSS release, that will be used to calibrate *Gaia* DR4 (Section 5.3). Concerning the internal homogeneity of the SPSS grid, we observe a scatter of the order of 1 percent in various literature comparisons, but the current SPSS V2 version shows a slightly discrepant (≈ 1 percent) behaviour for faint blue stars in the red part of the SPSS flux tables ($\gtrsim 700$ nm) that will be investigated in future releases.

The current SPSS release is tied to the CALSPEC grid, as it was in 2013. The CALSPEC reference system oscillates by 0.5–2.0 percent from release to release (Bohlin 2014; Bohlin et al. 2020, see also Section 4), but the latest three CALSPEC releases appear to converge on sub-1 percent readjustments of the overall flux zero-point in the optical and near-IR range. One important limitation of the CALSPEC grid, and indeed of all the available

¹⁹<https://gea.esac.esa.int/archive/documentation/GDR2/>

sets of flux standards, is that very few direct flux measurements exist (Deustua, Kent & Smith 2013). Thus the CALSPEC grid is tied to the best flux models available in the literature for the three WD pillars (Bohlin et al. 1995). WDs are relatively simple, their models are well understood and reliable, and the stars are rarely variable if one avoids the ZZ Ceti, V777 Her, and GW Vir instability strips (Rauch et al. 2014; Hermes et al. 2017). However, an accurate and updated direct measurement of stellar fluxes is still lacking and very difficult to attain in practice. In other words, the vast majority of the spectrophotometric flux standards available in the literature (including the Landolt and CALSPEC systems) are in practice made of secondary standards, only indirectly based on actual standardized measurements. Few primary standards exist for the UV, optical, and near-IR ranges, such as Vega, (Blackwell et al. 1983; Selby et al. 1983; Megessier 1995) Sirius (in the IR, Price et al. 2004), and the Sun (Thuillier et al. 2003).

Several astrophysical investigations, including the dark energy surveys based on observations of Type Ia supernovae, do require an accurate (i.e. absolute) flux calibration to 0.5 per cent or better (Albrecht et al. 2006). Therefore, updated and numerous direct measurements of flux standards are a pressing need, to go as much as possible beyond the 1 per cent limit currently experienced by the best flux standard sets available. Vega, the best measured primary standard so far, displays $\simeq 4$ per cent flux variations and an IR excess caused by a debris disc (Megessier 1995). Therefore, Sirius has been proposed as a more reliable primary standard, but it can count on fewer standard flux measurements so far (Price et al. 2004). The ‘NIST stars’²⁰ is an ongoing project to measure the irradiance of a few reliable stars, known to vary less than 1 per cent in flux, in a dedicated and carefully designed experiment (Smith et al. 2009). To complement NIST stars, ACCESS (Kaiser et al. 2017) is a project consisting in series of sub-orbital rocket and ground-based experiments, that will target a few stars, including Sirius and Vega, to improve the knowledge of their flux beyond the 1 per cent limit. Another US-funded experiment, ALTAIR (Albert et al. 2015) will launch balloons with artificial sources to calibrate a few primary standards. The project is presently in a testing phase, evaluating a series of commercial photodiodes.

In any case, for the moment, the main reference systems for the optical range remain CALSPEC for spectroscopy and Landolt for photometry. Our selection of SPSS – optimized for the flux calibration needs of *Gaia* and homogeneously calibrated – compares well to both systems, to the 1 per cent level or better. When new primary standards will be established and their calibration will be transferred to stars in the magnitude range of the SPSS, we will be able to refine our calibration, because the SPSS grid contains some of the most widely used spectrophotometric stars, such as the three CALSPEC pillars.

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²⁰<https://www.nist.gov/programs-projects/nist-stars>

DATA AVAILABILITY

The SPSS and PVL flux tables are available through the dedicated SPSS@SSDC database and archive interface (<http://gaiaextra.ssdsc.a si.it:8900>), together with all the relevant metadata and information. All the paper tables are available online as supplementary material to this paper on MNRAS. The data can also be obtained through the Vizier catalogue server (<http://vizier.u-strasbg.fr/viz-bin/VizieR>) at the CDS (<https://cds.u-strasbg.fr/>).

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SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](https://www.mnras.org/) online.

Table 1. List of SPSS included in the V1 and V2 sets along with basic information: the SPSS ID and name; the *Gaia* EDR3 ID and coordinates (rounded in the pdf table version); the *B* and *V* magnitudes from Altavilla et al. (2021), except for LTT 377 (SPSS 350), where they are from Koen et al. (2010); the spectral

type collected from the literature by Pancino et al. (2012); the SPSS set in which the star was included, either V1 for *Gaia* DR2 or V2 for *Gaia* (E)DR3; the quality flag (0 = Ok, 1 = warning on minor spectral defects, 2 = warning on spectral defects) and the notes.

Table 4. Parameters of the best-fitting models used to extend the SPSS flux tables to cover the full 300–1100 nm *Gaia* range. Stars extended with observed spectra are not listed here.

Table 7. List of stars included in the PVL along with basic information: the PVL ID and name, the *Gaia* EDR3 ID and coordinates (rounded), the literature *B* and *V* magnitudes and the spectral types from the original literature sources of the spectra.

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