

# Atomic and molecular absorption in redshifted radio sources

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## ABSTRACT

We report on a survey for associated H I 21-cm and OH 18-cm absorption with the Giant Metrewave Radio Telescope at redshifts  $z \approx 0.2\text{--}0.4$ . Although the low-redshift selection ensures that our targets are below the critical ultraviolet luminosity ( $L_{\text{UV}} \sim 10^{23} \text{ W Hz}^{-1}$ ), which is hypothesized to ionize all of the neutral gas in the host galaxy, we do not obtain any detection in the six sources searched. Analysing these in context of the previous surveys, in addition to the anticorrelation with the ultraviolet luminosity (ionizing photon rate), we find a correlation between the strength of the absorption and the blue–near-infrared colour, as well as the radio-band turnover frequency. We believe that these are due to the photoionization of the neutral gas, an obscured sightline being more conducive to the presence of cold gas and the compact radio emission being better intercepted by the absorbing gas, maximizing the flux coverage, respectively. Regarding the photoionization, the compilation of the previous surveys increases the significance of the critical ionizing photon rate, above which all of the gas in the host galaxy is hypothesized to be ionized ( $Q_{\text{H I}} \approx 3 \times 10^{56} \text{ s}^{-1}$ ), to  $>5\sigma$ . This reaffirms that this is an ubiquitous effect, which has profound implications for the detection of neutral gas in these objects with the Square Kilometre Array.

**Key words:** galaxies: active – galaxies: fundamental parameters – galaxies: ISM – quasars: absorption lines – radio lines: galaxies – ultraviolet: galaxies.

## 1 INTRODUCTION

Redshifted H I 21-cm absorption can provide an excellent probe of the contents and nature of the early Universe, through surveys that are not subject to the same flux and redshift constraints suffered by optical studies. For instance, measurement of the contribution of the neutral gas content to the mass density of the Universe at redshifts ( $z \lesssim 1.7$ ), where the Lyman  $\alpha$  transition is not accessible to ground-based optical telescopes. This corresponds to the past 10 Gyr of cosmic history, which are of particular interest since this is when star formation was at its most vigorous and the stellar mass density in the Universe increased by more than a factor of four (Hopkins & Beacom 2006). There is, however, mounting evidence that the star formation history is not best traced by *intervening* galaxies, detected in the Lyman  $\alpha$  absorption of a background continuum source:

(i) While the star formation history exhibits a strong evolution (e.g. Hopkins & Beacom 2006; Lagos et al. 2014), the neutral gas

content in damped Lyman  $\alpha$  absorption systems (DLAs),<sup>1</sup> remains approximately constant with look-back time (e.g. Rao, Turnshek & Nestor 2006; Braun 2012).

(ii) Very few DLAs have been detected in Lyman  $\alpha$  or H $\alpha$  emission (Möller, Fynbo & Fall 2004; Fynbo et al. 2010, 2011; Noterdaeme et al. 2012; Péroux et al. 2012), both tracers of star formation.

(iii) The heavy element content of DLAs does not appear to be caused by star formation within the absorbers themselves, but possibly deposited via winds from nearby galaxies (Fukugita & Ménard 2015).

(iv) The relative paucity of detections of H I 21-cm absorption in high-redshift DLAs appears to be dominated by extrinsic line-of-sight effects, rather than by any intrinsic evolution (Curran 2012; Curran et al. 2016c).

Therefore, in order to fully investigate any relation between the evolution of cold neutral gas and the star formation history, we

<sup>1</sup> Intervening systems where the neutral hydrogen column density exceeds  $N_{\text{H I}} = 2 \times 10^{20} \text{ cm}^{-2}$ .

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should also quantify the population of sources in which the absorption is *associated* with the continuum source itself.

The detection of H I 21-cm is of further interest as the comparison of its redshift with that of other transitions has the potential to measure past values of the fundamental constants of nature (Tzanavaris et al. 2005, 2007), which may exhibit a spatial (Berengut et al. 2011; Webb et al. 2011), as well as a temporal (Murphy, Webb & Flambaum 2003) variation. Radio data can yield at least an order of magnitude in precision in the measurement of the constants over the optical data (see Curran, Kanekar & Darling 2004a). Furthermore, the OH radical can not only provide measurements of various combinations of constants through comparison with optical, H I or rotational (millimetre) transitions, but intercomparison of the hyperfine 18-cm transitions can remove the systematics introduced by the possible velocity offsets present between species (Darling 2003). However, the number of redshifted OH absorption systems remains a paltry five, all of which are at redshifts of  $z \leq 0.89$  (Wiklind & Combes 1994, 1995, 1996, 1998; Chengalur, de Bruyn & Narasimha 1999; Kanekar & Chengalur 2002; Kanekar et al. 2003, 2005).

In addition to these, molecular absorption has also been detected in 25 DLAs, through H<sub>2</sub> vibrational transitions redshifted into the optical band at  $z \gtrsim 1.7$  (compiled in Srianand et al. 2010 with the addition of Reimers et al. 2003; Fynbo et al. 2011; Guimarães et al. 2012; Srianand et al. 2012; Noterdaeme et al. 2015, 2017). However, extensive millimetre-wave observations have yet to detect absorption from any rotational molecular transition (e.g. Curran et al. 2004b), leading us to suspect that the choice of optically bright objects selects against dusty environments, which are more likely to harbour molecules in abundance. This is apparent when one compares the DLAs in which H<sub>2</sub> has been detected, which have molecular fractions  $\mathcal{F} \equiv \frac{2N_{\text{H}_2}}{2N_{\text{H}_2} + N_{\text{H I}}} \sim 10^{-7} - 0.3$  and optical–near-infrared colours of  $V - K \lesssim 4$  (Curran et al. 2011c), with the five known OH absorbers, where  $\mathcal{F} \approx 0.7 - 1.0$  and  $V - K \geq 4.80$  (Curran et al. 2006). This is a strong indicator that the reddening is due to dust, which protects the molecular gas from the ambient ultraviolet radiation.

Therefore by selecting very red objects, with colours of  $V - K \gtrsim 5$ , we may expect OH column densities that can be detected with current large radio telescopes. However, there exists an additional constraint when searching for absorption by gas in the hosts of radio galaxies and quasars – specifically, that neutral hydrogen has never been detected where the ultraviolet ( $\lambda = 912 \text{ \AA}$ ) luminosity of the source exceeds  $L_{\text{UV}} \sim 10^{23} \text{ W Hz}^{-1}$ . This ‘critical’ luminosity applies to all redshifts for various heterogeneous (unbiased) samples, as seen through the non-detection of H I 21-cm absorption at  $L_{\text{UV}} \gtrsim 10^{23} \text{ W Hz}^{-1}$  by Curran et al. (2008, 2011a, 2013a,b, 2016a), Allison et al. (2012), Geréb et al. (2015), Aditya, Kanekar & Kurapati (2016), Grasha et al. (2017). This is interpreted as the flux limited optical spectroscopic surveys, which yield the redshift, selecting the most ultraviolet luminous objects at high redshift (Curran et al. 2008), where the corresponding ionizing photon rates of  $Q_{\text{H I}} \gtrsim 3 \times 10^{56} \text{ s}^{-1}$  are sufficient to ionize all of the neutral gas in a large spiral galaxy (Curran & Whiting 2012).<sup>2</sup>

At  $z \gtrsim 1$  the vast majority of radio sources for which redshifts are available are believed to have luminosities above the critical

value (see fig. 4 of Morganti, Sadler & Curran 2015). Thus, in order to increase the associated absorption statistics, we can select from the large population of radio sources of known redshift at  $z \lesssim 1$  with  $B \gtrsim 17$ .<sup>3</sup> This should yield UV luminosities below the critical  $L_{\text{UV}} \sim 10^{23} \text{ W Hz}^{-1}$  in the 1400 MHz band ( $z \lesssim 0.4$ , see fig. 1 of Curran et al. 2013b). This magnitude selection also has the advantage of giving large blue–near-infrared colours, where we may expect  $B - K \approx 6 - 10$  for the five known OH absorbers, on the basis of their optical–near-infrared colours of  $V - K \approx 5 - 9$  (Curran et al. 2006).

From the first part of this survey, we (Curran et al. 2011b) obtained one detection from four targets, three of which were unaffected by radio frequency interference (RFI), following which we were awarded further observing time on the Giant Metrewave Radio Telescope (GMRT) to complete the remainder of the requested sample. Here, we present our results and discuss their implications: in Section 2, we describe the sample selection, observations and analysis, in Section 3 we present our results, in Section 4 we discuss these in context of the previous H I 21-cm searches and in Section 5 we present our conclusions.

## 2 OBSERVATIONS AND ANALYSIS

### 2.1 Sample selection

In order to obtain a  $z \lesssim 1$  sample of H I and OH absorbers, we selected sources for which both transitions would be redshifted into the 1420 MHz receiver band (which spans 1000–1450 MHz). To ensure sufficient flux against which to detect the absorption (Section 2.2), sources were selected from the Parkes Half-Jansky Flat-spectrum Sample (PHFS, Drinkwater et al. 1997; Francis, Whiting & Webster 2000), giving a total of 10 targets with redshifts of  $z = 0.219 - 0.405$ . Lastly, the sources were prioritized by faintness, for which we chose  $B \gtrsim 19$  (as quoted in the PHFS, Table 1), since this gave the 10 faintest targets for which the estimated flux density at the redshifted H I 21-cm absorption frequency,  $S_{\text{est}}$ , was confirmed to exceed 0.5 Jy.

Verifying that the magnitude selection did yield targets exceeding the critical UV luminosity/ionizing photon rate, as described in Curran et al. (2013a), for each source we obtained the photometry from NASA/IPAC Extragalactic Database (NED), the *Wide-Field Infrared Survey Explorer* (WISE, Wright et al. 2010), Two Micron All Sky Survey (2MASS, Skrutskie et al. 2006) and the *Galaxy Evolution Explorer* (GALEX, data release GR6/7)<sup>4</sup> data bases. Each datum was corrected for Galactic extinction (Schlegel, Finkbeiner & Davis 1998) and a power law was fit to the UV rest-frame data, allowing the ionizing photon rate,  $Q_{\text{H I}} \equiv \int_{\nu}^{\infty} (L_{\nu}/h\nu) d\nu$ , to be derived from

$$\int_{\nu}^{\infty} \frac{L_{\nu}}{h\nu} d\nu, \quad \text{where } \log_{10} L_{\nu} = \alpha \log_{10} \nu + C \Rightarrow L_{\nu} = 10^C \nu^{\alpha},$$

where  $\alpha$  is the spectral index and  $C$  the intercept. This gives,

$$\frac{10^C}{h} \int_{\nu}^{\infty} \nu^{\alpha-1} d\nu = \frac{10^C}{\alpha h} [\nu^{\alpha}]_{\nu}^{\infty} = \frac{-10^C}{\alpha h} \nu^{\alpha}, \quad \text{where } \alpha < 0,$$

shown by the shaded regions in Fig. 1, all of which give rates below the critical  $\log_{10} Q_{\text{H I}} = 56.5 \text{ s}^{-1}$  (Table 1).

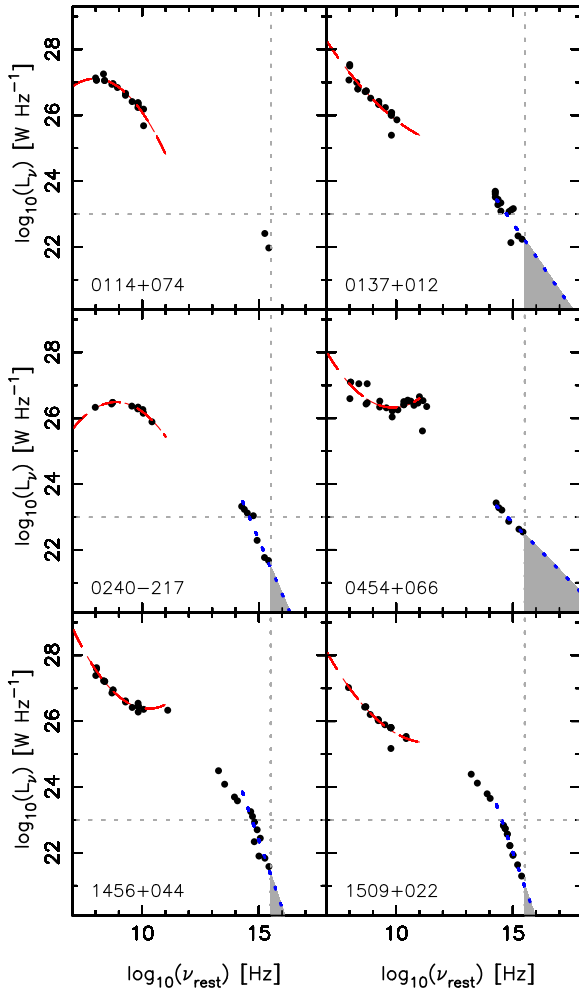
<sup>2</sup> Aditya et al. (2017) have recently detected H I 21-cm absorption at  $z = 1.223$  in TXS 1954+513, which they claim has an ultraviolet luminosity of  $L_{\text{UV}} = 4 \times 10^{23} \text{ W Hz}^{-1}$ . Although unreliable, this luminosity is still consistent with the critical value (see Section 3.1).

<sup>3</sup> The  $B$  magnitudes are generally more available than the  $V$ , as well as being given in the Parkes Half-Jansky Flat-spectrum Sample catalogue (Drinkwater et al. 1997).

<sup>4</sup> <http://galex.stsci.edu/GR6/#mission>

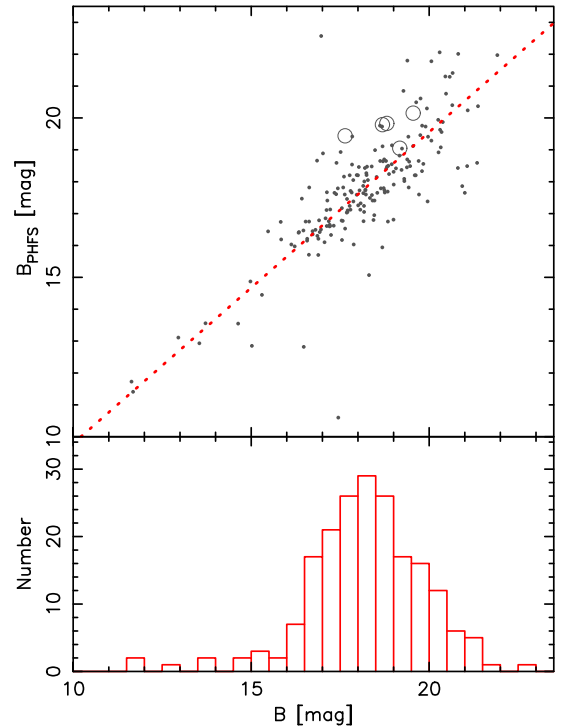
**Table 1.** The target list by IAU name (the NED names are given in Section 2.2).  $z$  is the optical redshift of the source from the PHFS (see Section 2.2),  $B$  (mag) the blue magnitude quoted in the PHFS,  $B - K$  (mag) the corresponding blue–near-infrared colour, followed by the values derived from our SED fitting. As stated in the main text, the PHFS  $B$  magnitudes may be erroneous by  $>1$  mag and the uncertainties in the  $K$  magnitudes, from 2MASS, range from 0.06 to 0.12 mag. The last five columns give the flux density at the redshifted H I 21-cm absorption frequency,  $S_{\text{fit}}$ , the rest-frame radio-band turnover frequency,  $\nu_{\text{TO}}$ , the spectral index, where  $S \propto \nu^\alpha$ , the estimated  $\lambda = 912$  Å monochromatic luminosity,  $L_{\text{UV}}$ , and the ionizing photon rate,  $Q_{\text{H I}}$ , from our fits (see Fig. 1). All quoted uncertainties are derived from the residuals to the fits.

Source	$z$	PHFS magnitudes		Derived magnitudes		$S_{\text{fit}}$	$\nu_{\text{TO}}$	$\alpha$	$L_{\text{UV}}$	$\log_{10} Q_{\text{H I}}$
		$B$ (mag)	$B - K$	$B$ (mag)	$B - K$	(Jy)	(MHz)		(W Hz $^{-1}$ )	(s $^{-1}$ )
0114+074	0.343	22.14	6.75	–	–	$1.98 \pm 0.11$	$93 \pm 1$	$-0.59 \pm 0.11$	–	–
0137+012	0.260	19.44	5.62	$17.64 \pm 1.21$	$3.95 \pm 0.27$	$1.42 \pm 0.17$	–	$-0.64 \pm 0.17$	$1.5 \times 10^{22}$	$55.4 \pm 0.3$
0240–217	0.314	19.05	5.57	$19.17 \pm 0.73$	$4.83 \pm 0.18$	$1.23 \pm 0.04$	$762 \pm 3$	$-0.12 \pm 0.04$	$2.8 \times 10^{21}$	$54.4 \pm 0.2$
0454+066	0.405	19.79	5.37	$18.69 \pm 0.20$	$3.59 \pm 0.04$	$0.65 \pm 0.18$	–	$-0.29 \pm 0.18$	$1.2 \times 10^{23}$	$55.8 \pm 0.1$
1456+044	0.391	20.15	5.57	$19.55 \pm 1.20$	$4.64 \pm 0.28$	$1.22 \pm 0.07$	–	$-0.52 \pm 0.07$	$2.0 \times 10^{21}$	$54.2 \pm 0.6$
1509+022	0.219	19.83	5.41	$18.81 \pm 0.35$	$5.16 \pm 0.10$	$1.03 \pm 0.14$	–	$-0.65 \pm 0.14$	$9.1 \times 10^{20}$	$53.8 \pm 0.1$



**Figure 1.** The rest-frame SED for each of our targets overlaid by fits to the photometry. The broken curve shows the second-order polynomial fit to the radio data and the dotted line the power-law fit to the UV data. The vertical dotted line signifies a rest-frame frequency of  $3.29 \times 10^{15}$  Hz ( $\lambda = 912$  Å) and the horizontal the critical  $\lambda = 912$  Å luminosity of  $L_{\text{UV}} \sim 10^{23}$  W Hz $^{-1}$ , with the shading showing the region over which the ionizing photon rate is derived.

Fitting the optical photometry by the same method gave the blue and near-infrared magnitudes, where the latter would be primarily obtained directly from the *WISE* and 2MASS data. From this we found brighter blue magnitudes than expected. Drinkwater et al.



**Figure 2.** The PHFS blue magnitudes (from COSMOS) versus those derived from our SED fits, where possible. The large circles designate our targets and the line the least-squares fit. The gradient of this is 0.977 (regression coefficient 99.46), showing that our  $B$  estimates are, on average, slightly higher than those in the PHFS.

(1997) obtain the blue magnitudes for the PHFS from the Cosmological Evolution Survey (COSMOS) catalogue (Yentis et al. 1992), where the uncertainty quoted is  $\pm 0.5$  mag, although Drinkwater et al. find that, due to a lack of calibration, some magnitudes may be erroneous by  $>1$  mag. In order to check this, in Fig. 2 we compare the blue magnitudes obtained from the fit to the spectral energy distribution (SED) of each PHFS source with those quoted in the catalogue. From a least-squares fit to the distribution, we see that, although our fits suggest larger  $B$  values ( $\mu = 18.19 \pm 0.12$  for the whole sample) than quoted in the PHFS ( $\mu = 17.80 \pm 0.14$ ), the magnitudes of our targets may be consistent with the PHFS values within the uncertainties (Table 1). However, the large possible uncertainties in the magnitudes were not a consideration when selecting targets and so these sightlines are generally not be as reddened as originally thought.

## 2.2 Observations and data reduction

All of the targets were observed over 2012 January 26–31, with the 1420 MHz receiver backed with the FX correlator over a bandwidth of 16 MHz spread over 512 channels in orthogonal polarizations. This gave a channel spacing of  $\approx 8 \text{ km s}^{-1}$  (cf. a full-width half-maximum of  $\text{FWHM} = 8\text{--}210 \text{ km s}^{-1}$  for the five known OH absorbers, see Curran et al. 2007), while maintaining a redshift coverage of  $\Delta z \approx \pm 0.01$ , in order to cover any uncertainties in the redshifts.<sup>5</sup> The full 30 antenna array was requested, with each transition in each source being observed for a total of 1 h, in order to reach a root mean square (rms) noise level of  $\approx 1 \text{ mJy}$  per channel. For a flux density of  $S_{\text{fit}} \gtrsim 0.5 \text{ Jy}$ , this gives a  $3\sigma$  optical depth limit of  $\tau \approx 0.005$  per channel, or a sensitivity to  $N_{\text{H I}} \sim 1 \times 10^{17} \times (T_{\text{spin}}/f) \text{ cm}^{-2}$  per  $8 \text{ km s}^{-1}$  channel, which is close to the lower limit for all of the published H I 21-cm searches. For the OH 1667 MHz transition, this corresponds to a sensitivity to  $N_{\text{OH}} \sim 2 \times 10^{13} \times (T_{\text{ex}}/f) \text{ cm}^{-2}$  per  $8 \text{ km s}^{-1}$  channel, cf.  $0.55\text{--}56 \times 10^{13} \times (T_{\text{ex}}/f) \text{ cm}^{-2}$  for the five known OH absorbers (Kanekar & Chengalur 2002; Kanekar et al. 2005).

For each source, 3C 48, 3C 147 and 3C 286 were used for band-pass calibration and a strong nearby point source for phase calibration. However, since this was performed only once every 30 min, self-calibration of the delays usually produced a superior image. The data were flagged and reduced using the MIRIAD interferometry reduction package, with flagging of the edge channels leaving the central 470 channels ( $\approx \pm 2000 \text{ km s}^{-1}$ ) from which a bad channel (144 in polarization XX) was removed from all visibilities. After averaging the two polarizations, a spectrum was extracted from the cube. Regarding each source:

**4C +07.04 (0114+074):** The H I band was observed for a total of 1.05 h. Five antennas (14, 20, 22, 23 and 30) were non-functioning, with antennas 19 and 26 also being flagged due to a severe bandpass ripple, leaving 300 baseline pairs. Self-calibration of the phases produced a superior image in which the source was unresolved by the  $5.6 \text{ arcsec} \times 4.3 \text{ arcsec}$  synthesized beam.

The OH band was observed for a total of 1.27 h and after flagging of the non-functioning antennas (14, 20, 22, 23 and 30), 300 baseline pairs remained. Again, a far superior image was obtained through self-calibration of the phases, giving a synthesized beam of  $4.1 \text{ arcsec} \times 3.5 \text{ arcsec}$ . The extracted spectrum shows a strong emission feature, which we believe to be an artefact (see Section 3.2).

**UM 355 (0137+012):** H I was searched for a total of 1.05 h. After flagging non-functioning antenna 20, antennas 23, 28, 29 and 30 were removed, due to less than ideal phase calibration. Although all of the remaining phases were well behaved, self-calibration of this source could not produce a satisfactory image. The calibration was then obtained from the phase calibrator LBQS 0056–0009, and the data from the unflagged 351 baseline pairs. The synthesized beam was  $5.3 \text{ arcsec} \times 3.4 \text{ arcsec}$  giving the partially resolved main component and a separate feature (see Section 3.3). The OH band was observed for a total of 1.28 h. After removal of the non-functioning antennas (12, 20, 22, 23 and 30), 300 good baseline pairs remained.

**PKS 0240–217:** The H I band was observed for a total of 0.90 h. Only one antenna (20) was non-functioning, leaving 406 baseline

pairs. Some minimal RFI was removed from the first half of the observation and after self-calibration of the phases a high-quality image was produced, although there are still some narrow-band RFI spikes present in the extracted spectrum. Unfortunately, this is concentrated at  $\approx 1082 \text{ MHz}$ , close to the expected absorption frequency (Fig. 3). The source was unresolved by the  $10.4 \text{ arcsec} \times 3.4 \text{ arcsec}$  synthesized beam.

For the OH band, non-functioning antennas (4 and 20) were flagged, leaving 378 baseline pairs. Again, by self-calibrating, an excellent image was produced from which the spectrum was extracted. The source was unresolved by the  $4.7 \text{ arcsec} \times 2.6 \text{ arcsec}$  synthesized beam.

**4C +06.21 (0454+066):** H I was searched for a total of 0.66 h. In addition to antenna 20, antenna 4 was found to be non-functioning with antenna 26 giving extremely noisy spectra. After flagging these, 351 baseline pairs remained from which the image was produced, using 3C 120 for phase calibration. The source was unresolved by the  $4.3 \text{ arcsec} \times 3.2 \text{ arcsec}$  synthesized beam.

For the OH band, antenna 20 was non-functioning and after the removal of the badly behaving antenna 12, 378 baseline pairs remained. However, RFI was apparent below 1182 MHz requiring the first 200 channels to be flagged, leaving 290. The source was unresolved by the  $3.3 \text{ arcsec} \times 2.7 \text{ arcsec}$  synthesized beam.

**4C +04.49 (1456+044):** H I was searched for a total of 0.47 h. After removing non-functioning antennas (4 and 26), the non-functioning baseline pair 2–8 was also removed, leaving 377 pairs. Self-calibration of the phases proved unsatisfactory and so the nearby 3C 298 (1416+067) was used. Severe RFI below 1016 MHz required removal of the first 70 remaining channels, leaving 420. The  $4.9 \text{ arcsec} \times 3.4 \text{ arcsec}$  beam reveals a double source (see Section 3.3).

For the OH band, non-functioning antennas (4 and 26) were removed, leaving 378 baseline pairs. RFI was present on all baselines at  $\gtrsim 1200 \text{ MHz}$  and using 3C 298 for the phase calibration also revealed the second feature present in the lower frequency, resolved by the  $3.2 \text{ arcsec} \times 2.9 \text{ arcsec}$  beam.

**PKS 1509+022:** Was searched in the H I band for a total of 0.73 h. After flagging the non-functioning antennas (4, 22, 23 and 26), phase calibration, using 3C 327.1, failed to produce a quality image even, after the removal of poorly performing antennas (17, 24, 28 and 30), which left 231 baseline pairs. This was probably due to the fact that this calibrator is over 1 h in distance from the target source. Self-calibration of the source required removal of the aforementioned antennas before the phases could be calibrated. Even so, no good image could be produced and so the spectrum was obtained by averaging visibilities of the remaining 231 baselines.

The OH band was observed for a total of 0.85 h and, like the H I band observations, could not be calibrated, even after the removal of non-functioning antennas. Further flagging, of the noisiest baseline pairs (where the rms exceeded 1 Jy), left 210 baseline pairs, which were averaged together to obtain a spectrum.

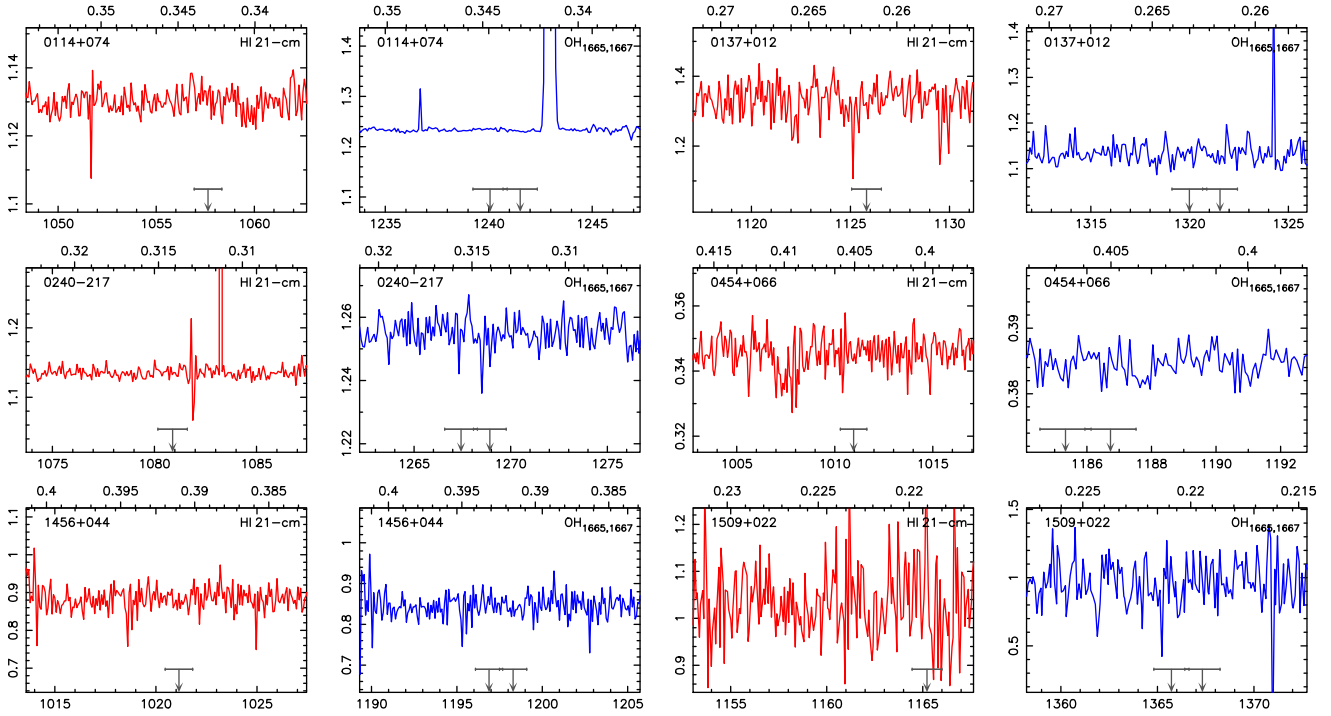
## 3 RESULTS

### 3.1 Observational results

In Fig. 3, we show the final spectra from which we have obtained no detections in the six targets searched, with the details summarized in Table 2. In addition to our targets not being sufficiently red (Section 2.1) to detect OH (Curran et al. 2006), each of the five known redshifted systems OH was detected following a clear H I detection (Carilli, Perlman & Stocke 1992; Carilli, Rupen & Yanny 1993;

<sup>5</sup> These are  $\delta z = 0.000011$  for 0137+012 (Ho & Kim 2009),  $\delta z = 0.00035$  for 1456+044 and  $\delta z = 0.00035$  for 1509+022 (Hewett & Wild 2010). Uncertainties are not available for the remaining three targets, but since the redshifts are quoted to three decimal places, the choice of  $\Delta z \approx \pm 0.01$  should ensure ample coverage.





**Figure 3.** The spectra from the H I 21-cm and OH 18-cm main line (1665 and 1667 MHz) searches shown at a spectral resolution of  $20 \text{ km s}^{-1}$ . All are extracted from the cube apart from the 1509+022 spectra, which are averages of the unflagged visibilities. The ordinate gives the flux density (Jy) and the abscissa the barycentric frequency (MHz). The scale along the top shows the redshift of either the H I 1420 MHz or OH 1667 MHz transition over the frequency range and the downwards arrow shows the expected frequency of the absorption from the optical redshift, with the horizontal bar showing a span of  $\pm 200 \text{ km s}^{-1}$  for guidance (the profile widths of the H I 21-cm detections range from 18 to  $475 \text{ km s}^{-1}$ , with a mean of  $167 \text{ km s}^{-1}$ , Curran et al. 2013a).

**Table 2.** The observational results.  $z$  is the optical redshift of the source,  $S_{\text{meas}}$  is the measured flux density,  $\Delta S$  the rms noise reached per  $\Delta \nu$  channel and  $\tau$  the derived optical depth, where  $\tau_{3\sigma} = -\ln(1 - 3\Delta S/S_{\text{cont}})$  is quoted for these non-detections. These give the quoted column densities ( $\text{cm}^{-2}$ ), where  $T_s$  is the spin temperature of the atomic gas,  $T_x$  is the excitation temperature (K) of the molecular gas and  $f$  the respective covering factor. Here, and throughout the paper, OH refers to the  $^2\Pi_{3/2}J = 3/2F = 2 - 2$  (1667 MHz) transition. Finally, we list the frequency and redshift range over which the limit applies.

Source	$z$	Line	$S_{\text{meas}}$ (Jy)	$\Delta S$ (Jy)	$\Delta \nu$ ( $\text{km s}^{-1}$ )	$\tau_{3\sigma}$	$N$ ( $\text{cm}^{-2}$ )	$\nu$ -range (MHz)	$z$ -range
0114+074	0.343	H I	1.130	0.003	9.2	$<0.009$	$<1.5 \times 10^{17} (T_s/f)$	1048–1063	0.336–0.356
–	–	OH	1.233	0.004	7.9	$<0.010$	$<1.9 \times 10^{13} (T_x/f)$	1232–1247	0.337–0.354
0137+012	0.260	H I	1.335	0.052	8.7	$<0.11$	$<1.8 \times 10^{18} (T_s/f)$	1119–1134	0.252–0.269
–	–	OH	1.133	0.022	7.4	$<0.058$	$<1.0 \times 10^{14} (T_x/f)$	1315–1327	0.256–0.268
0240–217	0.314	H I	1.135	0.007	9.0	$<0.019$	$<3.0 \times 10^{17} (T_s/f)$	1073–1088	0.306–0.324
–	–	OH	1.255	0.005	7.7	$<0.012$	$<2.2 \times 10^{13} (T_x/f)$	1262–1277	0.306–0.321
0454+066	0.405	H I	0.346	0.005	9.7	$<0.041$	$<7.2 \times 10^{17} (T_s/f)$	1002–1017	0.396–0.417
–	–	OH	0.385	0.002	8.2	$<0.016$	$<3.1 \times 10^{13} (T_x/f)$	1184–1193	0.398–0.409
1456+044	0.391	H I	0.884	0.033	9.6	$<0.11$	$<2.0 \times 10^{18} (T_s/f)$	1015–1028	0.382–0.400
–	–	OH	0.846	0.033	8.1	$<0.12$	$<2.3 \times 10^{14} (T_x/f)$	1193–1202	0.387–0.397
1509+022	0.219	H I	1.034	0.08	8.5	$<0.23$	$<3.6 \times 10^{18} (T_s/f)$	1153–1168	0.216–0.232
–	–	OH	0.972	0.16	7.1	$<0.68$	$<1.2 \times 10^{15} (T_x/f)$	1358–1373	0.214–0.228

Carilli et al. 1998; Chengalur et al. 1999; Kanekar & Briggs 2003). Furthermore, the OH absorption strength is only expected to be  $\lesssim 10^{-4}$  times that of the H I 21-cm absorption (Curran et al. 2007). Therefore, we will treat H I as a prerequisite for OH absorption and thus focus our discussion around the H I 21-cm results.

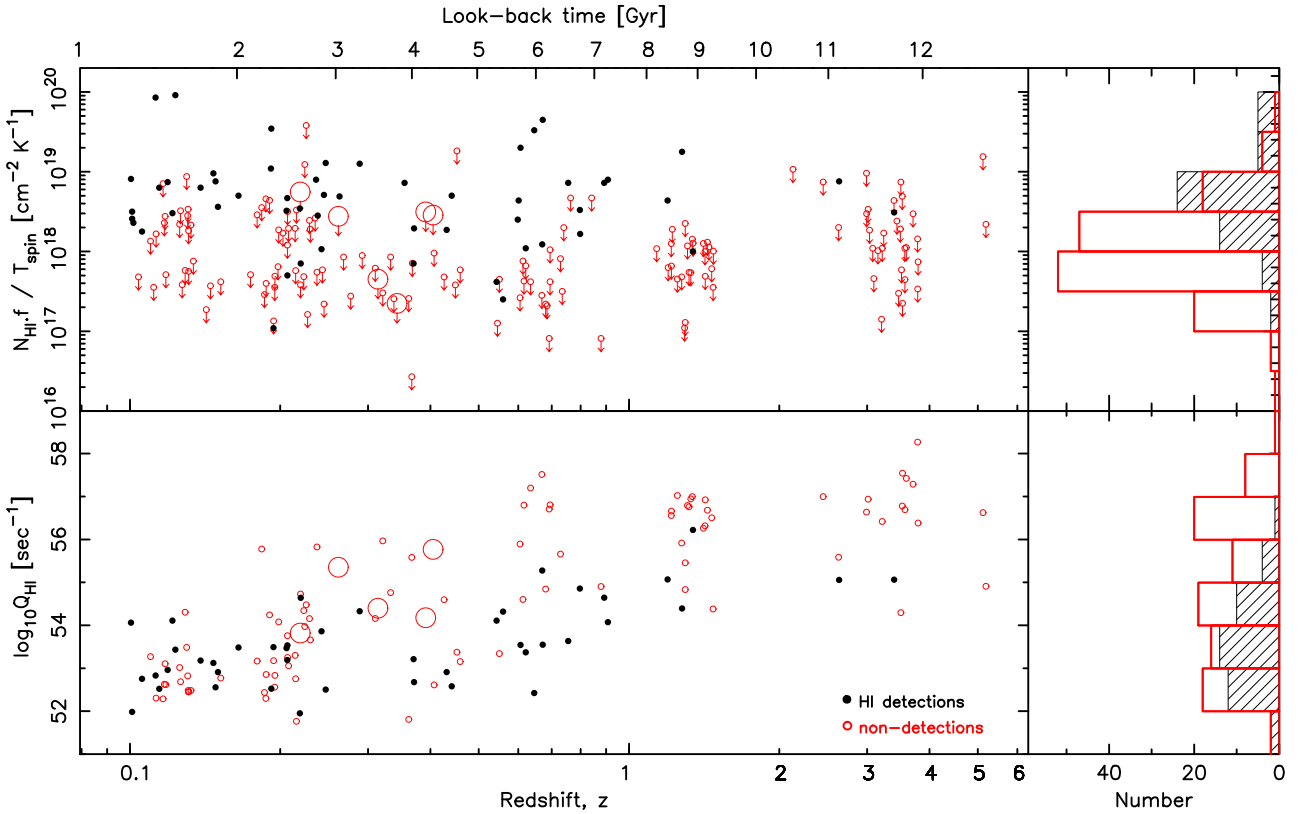
The total neutral hydrogen column density is related to the velocity integrated optical depth of the H I 21-cm absorption via,

$$N_{\text{H I}} = 1.823 \times 10^{18} T_{\text{spin}} \int \tau d\nu, \quad (1)$$

where  $T_{\text{spin}}$  is the spin temperature of the gas, which is a measure of the excitation from the lower hyperfine level (Purcell & Field 1956; Field 1959), and  $\int \tau d\nu$  is the velocity integrated optical depth of the absorption. The observed optical depth is related to this via

$$\tau \equiv -\ln \left( 1 - \frac{\tau_{\text{obs}}}{f} \right) \approx \frac{\tau_{\text{obs}}}{f}, \quad \text{for } \tau_{\text{obs}} \equiv \frac{\Delta S}{S_{\text{obs}}} \lesssim 0.3, \quad (2)$$

where the covering factor,  $f$ , is a measure of the fraction of observed background flux ( $S_{\text{obs}}$ ) intercepted by the absorber. In the optically thin regime (where  $\tau_{\text{obs}} \lesssim 0.3$ ), equation (1) can be rewritten as



**Figure 4.** The line strength ( $1.823 \times 10^{18} (T_{\text{spin}}/f) \int \tau_{\text{obs}} dv$ , top) and ionizing ( $\lambda \leq 912 \text{ \AA}$ ) photon rate (bottom) versus redshift for the  $z \geq 0.1$  H I 21-cm absorption searches. The filled circles/histogram represent the detections and the unfilled circles/histogram the  $3\sigma$  upper limits to the non-detections, with the large circles designating our targets.

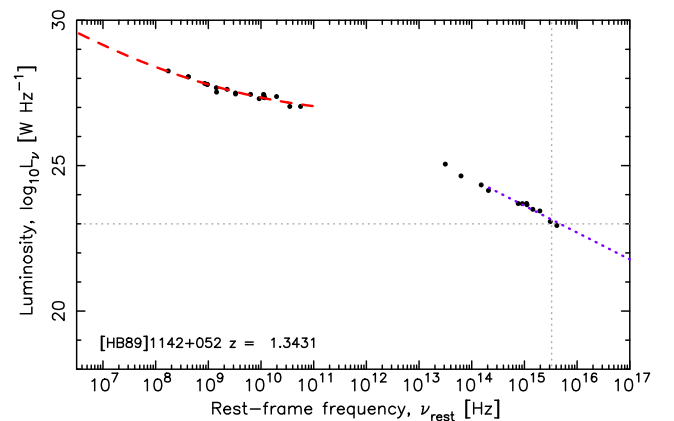
$N_{\text{H I}} \approx 1.823 \times 10^{18} (T_{\text{spin}}/f) \int \tau_{\text{obs}} dv$ . From Fig. 4 (top panel),<sup>6</sup> we see that the six targets have been searched in H I 21-cm to sensitivities comparable with previous detections.<sup>7</sup>

In the bottom panel, we show the ionizing photon rates for those sources for which there is sufficient blue/UV rest-frame photometry (Section 2.1). [HB89] 1142+052 now defines the highest value ( $Q_{\text{H I}} = 1.66^{+0.70}_{-0.49} \times 10^{56} \text{ s}^{-1}$ ) where H I has been detected (Kanekar et al. 2009).<sup>8</sup> This is close to the theoretical value of  $3 \times 10^{56} \text{ s}^{-1}$ , which is sufficient to ionize all of the neutral atomic gas in a large spiral galaxy (Curran & Whiting 2012), although from very limited

<sup>6</sup> In order to normalize the limits, each has been re-sampled to the same spectral resolution ( $20 \text{ km s}^{-1}$ , as in Fig. 3), which is used as the FWHM to obtain the integrated optical depth limit, thus giving the  $N_{\text{H I}} f / T_{\text{spin}}$  limit per channel (see Curran 2012).

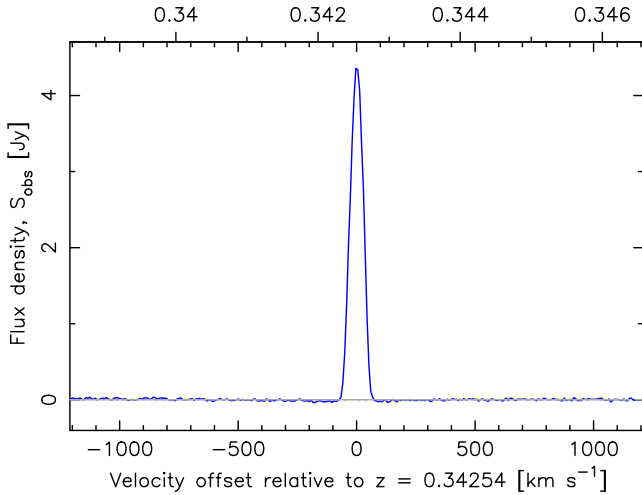
<sup>7</sup> The associated absorption searches have been compiled from de Waard, Strom & Miley (1985), Mirabel (1989), van Gorkom et al. (1989), Uson, Bagri & Cornwell (1991), Carilli et al. (1992, 1998, 2007), Moore, Carilli & Menten (1999), Peck, Taylor & Conway (1999), Peck et al. (2000), Röttgering et al. (1999), Morganti et al. (2001), Ishwara-Chandra, Dwarkanath, & Anantharamaiah (2003), Vermeulen et al. (2003), Curran et al. (2006, 2008, 2011a,b, 2013a,b, 2016a), Gupta et al. (2006), Orienti, Morganti & Dallacasa (2006), Kanekar et al. (2009), Emonts et al. (2010), Salter et al. (2010), Chandola, Sirothia & Saikia (2011), Chandola, Gupta & Saikia (2013), Allison et al. (2012, 2015), Yan et al. (2012, 2016), Geréb et al. (2015); Srianand et al. (2015), Aditya et al. (2016, 2017).

<sup>8</sup> The absorption redshift of  $z_{\text{abs}} = 1.343$  (Kanekar et al. 2009) corresponds to a velocity offset of  $\Delta v = 64 \text{ km s}^{-1}$  relative to  $z_{\text{QSO}} = 1.3425$  and so we deem this as associated rather than intervening.



**Figure 5.** The rest-frame SED of [HB89] 1142+052 overlaid by fits to the radio and optical-UV photometry (from the Sloan Digital Sky Survey data releases 2 and 6, Abazajian et al. 2004; Adelman-McCarthy et al. 2008). As per Fig. 1, the vertical dotted line signifies a rest-frame frequency of  $3.29 \times 10^{15} \text{ Hz}$  ( $\lambda = 912 \text{ \AA}$ ) and the horizontal the critical  $\lambda = 912 \text{ \AA}$  luminosity of  $L_{\text{UV}} \sim 10^{23} \text{ W Hz}^{-1}$ .

$\lambda > 912 \text{ \AA}$  photometry (Fig. 5). For  $Q_{\text{H I}} \leq 1.7 \times 10^{56} \text{ s}^{-1}$ , there are 43 detections and 66 non-detections, that is a 39.4 per cent detection rate for objects for which  $Q_{\text{H I}}$  can be estimated. Applying this to the  $Q_{\text{H I}} > 1.7 \times 10^{56} \text{ s}^{-1}$  sources, gives a binomial probability of  $4.80 \times 10^{-7}$  of obtaining 0 detections and 29 non-detections, which is significant at  $5.03\sigma$ , assuming Gaussian statistics. If we include



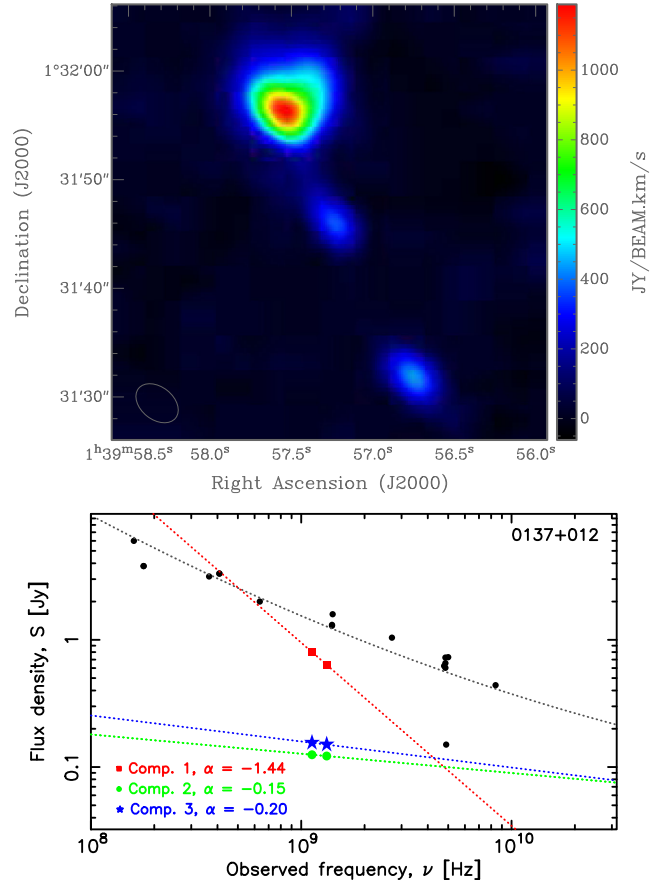
**Figure 6.** Detail of the ‘emission’ feature towards 0114+074 shown at the observed  $\Delta v = 7.85 \text{ km s}^{-1}$  resolution (cf. Fig. 3). The scale along the top shows the redshift of the OH 1667 MHz transition, which is also used to define the velocity offset.

the forthcoming results of Grasha et al. (2017),<sup>9</sup> this increases to  $6.67\sigma$ .

Aditya et al. (2017) have recently detected H I 21-cm absorption at  $z = 1.223$  in TXS 1954+513, which they claim has an ultraviolet luminosity of  $L_{\text{UV}} = 4 \pm 1 \times 10^{23} \text{ W Hz}^{-1}$ , following the method of Curran et al. (2008). The luminosity is, however, based upon only two photometry measurements that are in the rest-frame optical band and extrapolated to the UV. This is insufficient data to calculate  $Q_{\text{H I}} \equiv \int_{\nu}^{\infty} (L_{\nu}/h\nu) d\nu$ , which requires integration of the UV photometry above  $\nu_{\text{rest}} = 3.29 \times 10^{15} \text{ Hz}$  (see Section 2.1). If, for the sake of argument, we approximate  $Q_{\text{H I}} \sim L_{\text{UV}}/h \sim 6 \times 10^{56} \text{ s}^{-1}$ , our model (Curran & Whiting 2012) yields a scalelength of  $R = 3.05 \text{ kpc}$ , cf.  $3.15 \text{ kpc}$  for the Milky Way (Kalberla & Kerp 2009). That is, this ionizing photon rate is just sufficient to ionize all of the gas in a large spiral, which is consistent with the results of Curran & Whiting. Adding this value of  $Q_{\text{H I}}$  to the other data (Fig. 4, bottom), would change the above significance to  $3.30\sigma$  for the current data and  $4.50\sigma$  including the forthcoming data. Due to the lack of photometry, however, our own UV fitting method (which requires the interpolation of at least four points at  $\nu_{\text{rest}} > 4 \times 10^{14} \text{ Hz}$ , see Curran et al. 2013a) would reject a UV fit for this source and so we cannot assign an ultraviolet luminosity nor photoionization rate.

### 3.2 Emission feature in 0114+074

As seen in Fig. 3, there is a strong emission feature in the OH band of 0114+074. The feature was apparent in the averaged visibilities, before imaging, in both polarizations, close to the expected OH frequency (Fig. 6), although not apparent in the calibration sources. The emission has an integrated flux density of  $\int S_{\text{obs}} d\nu = 286 \pm 9 \text{ Jy km s}^{-1}$ , which gives  $L_{\text{OH}} = 870 L_{\odot}$ . This is within the range of known OH mega-masers ( $10^{2.2} < L_{\text{OH}} < 10^{3.8} L_{\odot}$ , e.g. Darling & Giovanelli 2002b; Lo 2005). However, given that the feature appears as absorption when off source, in conjunction with the fact that no 1665/1667 MHz doublet is apparent (e.g. Darling &



**Figure 7.** Top: GMRT 1126 MHz continuum image of 0137+012. The  $5.3 \text{ arcsec} \times 3.4 \text{ arcsec}$  synthesized beam is shown in the bottom left corner. Bottom: The spectral indices of the various components (designated 1–3 starting from the north) obtained from the redshifted H I and OH band observations. The small circles show the data from the literature overlain with the polynomial fit (see Fig. 1).

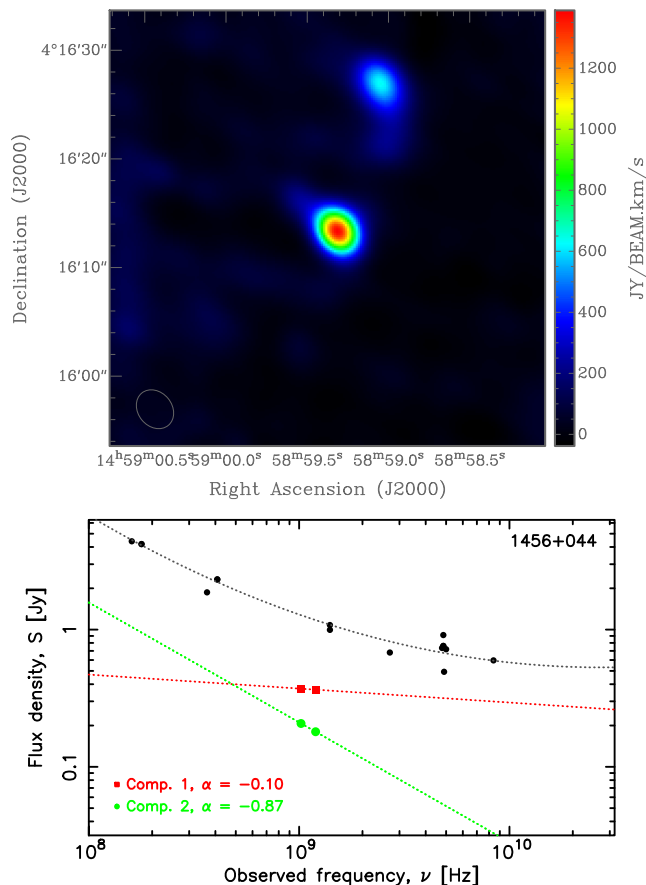
Giovanelli 2002a; Harvey-Smith et al. 2016), we attribute this to an artefact rather than being the detection of an OH mega-maser.

### 3.3 Resolved structure in 0137+012 and 1456+044

As mentioned in Section 2.2, more than one component was resolved in the observations of 0137+012 and 1456+044. In 0137+012 (Fig. 7, top) we see three separate features, previously detected at 1465 and 4885 MHz by Gower & Hutchings (1984), which they modelled as a damped precessing jet. Component 1 is by far the strongest, with flux densities of  $0.801 \pm 0.008$  and  $0.636 \pm 0.005 \text{ Jy}$  at 1125 and 1321 MHz, respectively, giving an H I optical depth limit of  $\tau_{3\sigma} < 0.029$  per  $8.7 \text{ km s}^{-1}$  channel. The other two components are significantly weaker ( $2 - 0.125 \pm 0.006 \text{ Jy}$  and  $0.122 \pm 0.008 \text{ Jy}$ , giving  $\tau_{3\sigma} < 0.14$  and  $3 - 0.155 \pm 0.005 \text{ Jy}$  and  $0.150 \pm 0.003 \text{ Jy}$ , giving  $\tau_{3\sigma} < 0.10$ ). They also have similar spectral indices which may suggest a connection, although the two frequencies are insufficiently separated to infer anything definite (Fig. 7, bottom).

For 1456+044, the image revealed a double source (Fig. 8, top). Component 1 has flux densities of  $0.371 \pm 0.010$  and  $0.365 \pm 0.09 \text{ Jy}$  at 1021 and 1198 MHz, respectively, giving an H I optical depth limit  $\tau_{3\sigma} < 0.081$  per  $8.5 \text{ km s}^{-1}$  channel. For component 2, the flux densities are  $0.207 \pm 0.06$  and  $0.180 \pm 0.06 \text{ Jy}$ , giving an H I

<sup>9</sup> Hereafter G17, which reports 0 new detections of H I 21-cm absorption out of 89 new searches over  $0.02 < z < 3.8$  (see Grasha & Darling 2011).



**Figure 8.** Top: GMRT 1021 MHz continuum image of 1456+044. The 4.9 arcsec  $\times$  3.4 arcsec synthesized beam is shown in the bottom left corner. Bottom: The spectral indices of the various components, designated 1 (main) and 2 (north), obtained from the redshifted H I and OH band observations.

optical of depth limit  $\tau_{3\sigma} < 0.093$ . From Fig. 8 (bottom), while the main component appears to have a very flat spectrum, component 2 has a similar spectral index to the cumulative emission at the measured frequencies. Furthermore, it is possible to envision the cumulative SED as being the sum of those of the two components, although there appears to be some missing flux and, again, the observed frequencies are very close to each other.

## 4 DISCUSSION

### 4.1 Factors affecting the detection rate

#### 4.1.1 Ultraviolet luminosity and survey sensitivity

Although all six sources searched have UV luminosities below the critical value, none were detected in H I 21-cm absorption. In Fig. 9, for the photoionizing rate and other important parameters, we show how the detection rate varies with the parameter in question. These are obtained from the number of detections normalized by the total number of searches within each bin, which we bin per decade for  $Q_{\text{H I}}$  (the first panel of Fig. 9). For our targets, whose values are shown by the vertical lines, we may expect a 28–52 per cent chance of detection. A decrease in the detection rate with redshift was first shown by Curran et al. (2008), which was interpreted as the high redshifts selecting UV luminosities sufficient to excite the

gas to below the detection limit.<sup>10</sup> Because of the Malmquist bias, it can be difficult to ascertain whether the decreasing detection rate is caused by an evolutionary effect at high redshift (Curran & Whiting 2010; Aditya et al. 2016), although ionization of the gas at high UV luminosities is physically motivated and is found to apply at *all* redshifts, no matter the selection criteria (Section 1). Showing the line strength versus the photoionizing rate (Fig. 10), we see a strong anticorrelation, which, for a given column density and covering factor (Section 3.1), would suggest an increase in the spin temperature due to an increased flux of  $\lambda \leq 912 \text{ \AA}$  photons (Bahcall & Ekers 1969). Again, it is also clear that there is a value of  $Q_{\text{H I}}$  above which H I 21-cm is not detected, although there is no ‘critical redshift’ apparent. For example, the two detections at  $z \gtrsim 2.5$  (Moore et al. 1999; Uson et al. 1991, which have  $Q_{\text{H I}} \approx 1 \times 10^{55} \text{ s}^{-1}$ , Fig. 4). Therefore, contrary to Aditya et al. (2016), we reaffirm that the luminosity–redshift degeneracy can be broken and that the decrease in detection rate is caused by an increase in the UV luminosity.

Referring to the second panel of Fig. 9, from the sensitivities we expect a detection rate of  $\gtrsim 20$  per cent, cf. 26 per cent for the whole sample, which rises to 29 per cent with the removal of the  $Q_{\text{H I}} \geq 3 \times 10^{56} \text{ s}^{-1}$  sources. This compares to the 25 per cent found by Gupta et al. (2006) and the 40 per cent by Vermeulen et al. (2003), although this was for compact objects, which trace less UV luminous sources (Curran & Whiting 2010; Allison et al. 2012).<sup>11</sup> From a  $\approx 30$  per cent detection rate, we may expect  $1.8 \pm 1.1$  (for binomial statistics) detections out of our six targets. This is within  $2\sigma$  of zero detections and so our results are not extraordinary. Nevertheless, it would be of value to future surveys to analyse other possible factors that could affect the detection of H I 21-cm absorption.

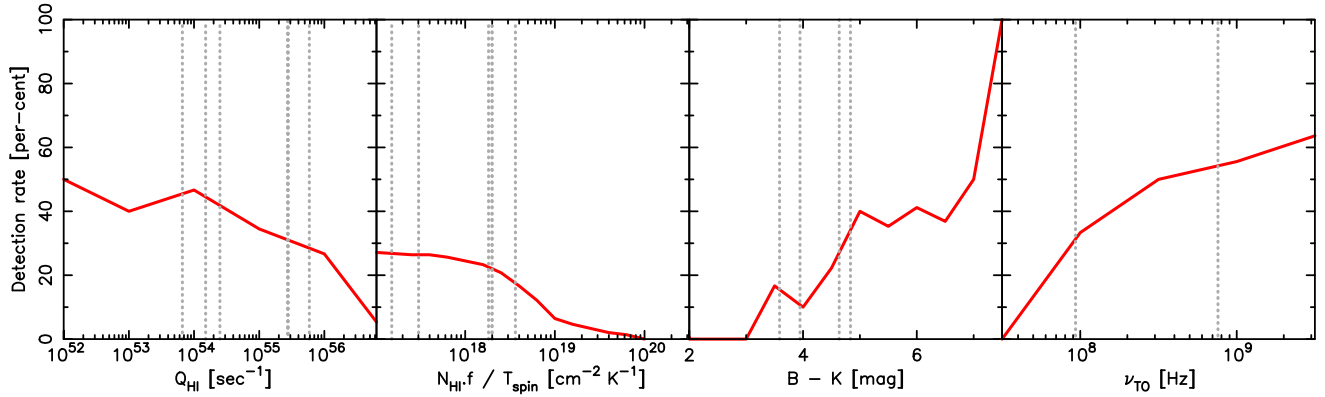
#### 4.1.2 Dust reddening

As discussed in Section 1, the correlation between molecular fraction and optical–near-infrared colour (Curran et al. 2011c) is evidence of dust reddening, where the dusty sightlines are more conducive to high molecular abundance. From a sample of five red quasars searched in associated H I 21-cm absorption, Carilli et al. (1998) obtained four detections, from which they suggested that optically selected samples, which have lower H I 21-cm detection rates, may bias against the detection of high column density absorbers. Following this, Curran et al. (2006) noted a trend between the H I 21-cm absorption strength and the optical–near-infrared colour, with Curran & Whiting (2010) reporting a  $3.63\sigma$  correlation between  $\int \tau d\nu$  and  $V - K$  for the 58 searched associated absorbers for which the colours were available. Since we have based the current sample on their faint blue magnitudes, in Fig. 11 we show the current distribution in terms of  $B - K$ . A generalized non-parametric Kendall-tau test gives a probability of  $P(\tau) = 0.0063$  [ $S(\tau) = 2.73\sigma$ ] of the correlation occurring by chance, with the inclusion of the forthcoming data (G17) increasing the significance to  $S(\tau) = 3.80\sigma$  [ $P(\tau) = 1.43 \times 10^{-4}$ ]. Furthermore, in the third panel of Fig. 9 we

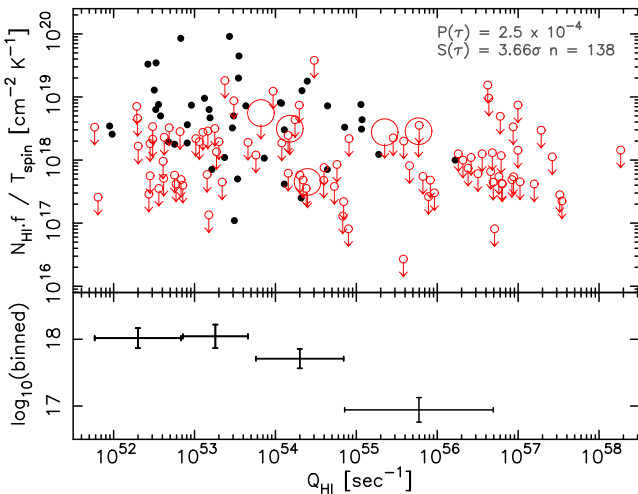
<sup>10</sup> Röttgering et al. (1999) also obtained only H I 21-cm non-detections in six  $z \gtrsim 2.4$  radio galaxies, but failed to realize the significance of the UV luminosity in their comparison with the highest redshift ( $z = 3.40$ ) H I 21-cm detection (Uson et al. 1991). This, B2 0902+34, has  $L_{\text{UV}} \approx 3 \times 10^{22} \text{ W Hz}^{-1}$ , cf.  $L_{\text{UV}} \gtrsim 10^{23} \text{ W Hz}^{-1}$  for each of their sample.

<sup>11</sup> The maximum expected detection rate is  $\sim 50$  per cent on account of the chance of the radiation being intercepted by the galactic disc along our sightline (see Curran & Whiting 2010 and references therein).





**Figure 9.** The H I 21-cm detection rate binned against the ionizing photon rate, the sensitivity, the blue–near-infrared colour and the turnover frequency (clear turnovers over the observed range only, see Fig. 1). These are binned to steps of (log) unity on the abscissa and intended only to show the overall trends, thus no uncertainties are shown. The vertical dotted lines show the values for our sample, where two overlap in the  $N_{\text{HI}}f/T_{\text{spin}}$  and  $B-K$  panels.

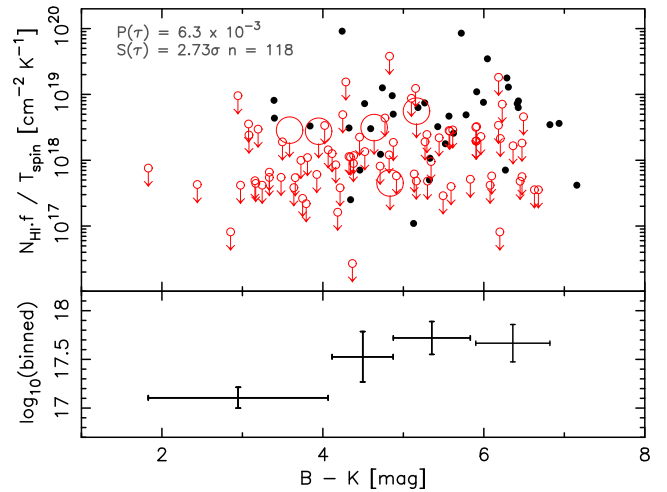


**Figure 10.** The H I 21-cm line strength versus the ionizing photon rate of the  $z \geq 0.1$  associated absorption searches. As per usual (e.g. Curran et al. 2016b), we include the H I 21-cm non-detected  $3\sigma$  limits as censored points via the *Astronomy SURVival Analysis* (ASURV) package (Isobe, Feigelson & Nelson 1986). A generalized non-parametric Kendall-tau test gives a probability of  $P(\tau) = 2.53 \times 10^{-4}$  of the correlation occurring by chance, which is significant at  $S(\tau) = 3.66\sigma$ , assuming Gaussian statistics. The inclusion of the forthcoming data (G17) increases the significance to  $S(\tau) = 5.74\sigma$  [ $P(\tau) = 9.58 \times 10^{-9}$ ]. The bottom panel shows the binned values, including the limits via the Kaplan–Meier estimator, in equally sized bins. It is not possible to derive a value for the last bin since this comprises only limits. The horizontal error bars show the range of points in the bin and the vertical error bars the  $1\sigma$  uncertainty in the mean value.

see a steep increase in the detection rate with the red colour. This and the strong positive correlation with the line strength, indicates that the reddening is caused by dust, the presence of which hinders excitation of the hydrogen above the lower hyperfine ( $F=0$ ) ground state.

#### 4.1.3 Coverage of the background continuum flux

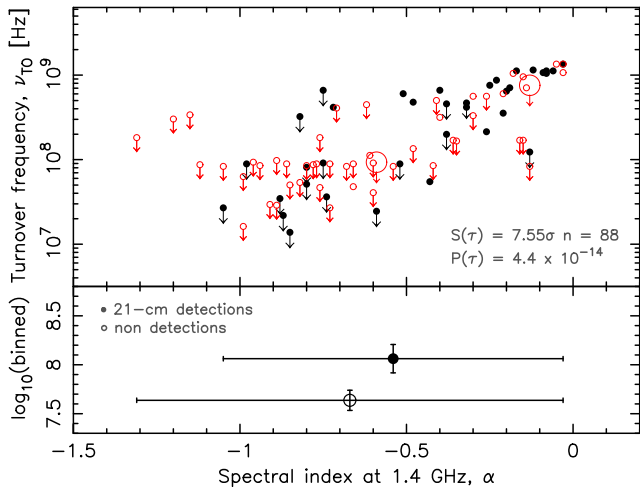
From equation (2), we see that the observed optical depth can be reduced by low coverage ( $f \ll 1$ ) of the background emission, an effect seen in both intervening and associated absorption (Curran 2012; Curran et al. 2013c, respectively). For a given (unknown) absorption cross-section, in the optically thin regime, the



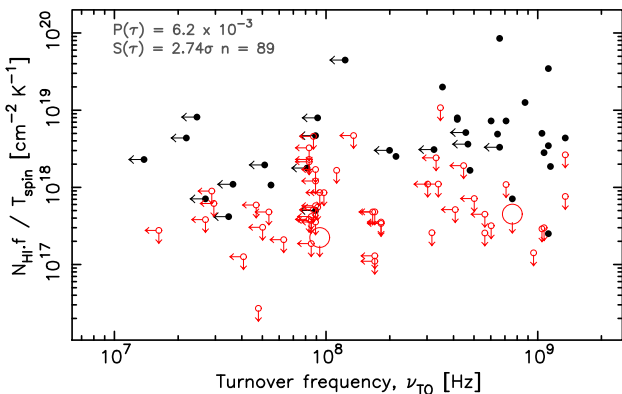
**Figure 11.** The H I 21-cm line strength versus the blue–near-infrared colours of the  $z \geq 0.1$  associated absorption searches. The significance rises to  $3.80\sigma$  ( $n = 176$ ) with the inclusion of the forthcoming data (G17).

observed optical depth will therefore be proportional to the extent of the radio emission (see Curran et al. 2005a). Extended emission may be apparent through a steep spectrum, due to the radio jets being aligned normally to our line of sight (seen edge-on), thus maximizing the apparent size of the source. Another signature of extended emission is a low/non-apparent turnover frequency, it being generally accepted that the turnover frequency of a radio source is anticorrelated with its extent (e.g. Fanti et al. 1990).

From a sample of nearby galaxies, Curran et al. (2016b) found that, while there was a large overlap in the spectral indices of the H I 21-cm detections and non-detections, detections did tend to occur towards sources with higher turnover frequencies. We obtain the turnover frequencies and the spectral indices from the fits described in Section 2.1 and plotting the H I 21-cm absorption searches in Fig. 12, we see a similar result where the detections have a mean turnover frequency of  $\langle \nu_{\text{TO}} \rangle = 115^{+45}_{-33}$  MHz, compared to  $\langle \nu_{\text{TO}} \rangle = 43^{+12}_{-9}$  MHz, for the non-detections. This is confirmed in Fig. 13 (see also the fourth panel of Fig. 9), where the positive correlation between the line strength and turnover frequency (suggesting an inverse correlation with source size), indicates that the covering factor is important. However, as seen in the top panel of the figure, not having an evident turnover frequency does not



**Figure 12.** The rest-frame turnover frequency versus the spectral index for the background sources for which these could be determined. If no turnover is apparent in the radio SED we assume that this occurs below the lowest observed frequency (typically  $\sim 10$  MHz) and use this to assign an upper limit to  $\nu_{\text{TO}}$ . Again, the filled symbols represent the detections and the unfilled the non-detections. In the bottom panel, the binned values of the detections and non-detections are shown.



**Figure 13.** The line strength versus the rest-frame turnover frequency. The inclusion of the forthcoming data (G17) does not change the significance ( $n = 100$ , cf. 89 here).

preclude a detection, with this being possible at  $\nu_{\text{TO}} < 20$  MHz. Therefore, it should be borne in mind that the spectral properties offer only an indirect measure of the emission region and the covering factor depends upon how this is related to the absorption cross-section, as well as the alignment between absorber and emitter (see Curran et al. 2013c).

#### 4.1.4 Summary

In Table 3, we summarize the probability of detection for the six targets discussed in this work, based upon the above parameters. These are obtained from the number of detections normalized by the total number of searches within the bin which hosts the target (Fig. 9). From these we see that, on an individual basis, the odds are against a detection with a maximum probability of 52 per cent for a single parameter ( $Q_{\text{H I}}$  for 1509+022 and  $\nu_{\text{TO}}$  for 0240–217). This, however, will be tempered by the probability from the other param-

**Table 3.** The probabilities of detection for each of our six targets based upon the properties of the H I 21-cm absorption searches. The rates are obtained from the intersection of the dotted vertical lines designating the values for our targets and the full line joining the binned values for all of the H I 21-cm searches (Fig. 9). The  $B - K$  rates have been interpolated between  $B - K = 3$  and 6.5 and, as per the figure, these are intended as approximate guide values only.

Source	$z$ -range	$Q_{\text{H I}}$	Sensitivity	$B - K$	$\nu_{\text{TO}}$
0114+074	0.336–0.356	–	27 per cent	–	33 per cent
0137+012	0.252–0.269	31 per cent	22 per cent	22 per cent	–
0240–217	0.306–0.324	41 per cent	26 per cent	27 per cent	52 per cent
0454+066	0.396–0.417	28 per cent	24 per cent	20 per cent	–
1456+044	0.382–0.400	44 per cent	24 per cent	26 per cent	–
1509+022	0.216–0.232	52 per cent	18 per cent	29 per cent	–

eters, which are not independent,<sup>12</sup> making the overall probability difficult to ascertain for each source.

So in addition to our hypothesis that we only expect detections where  $L_{\text{UV}} \lesssim 10^{23} \text{ W Hz}^{-1}$  ( $Q_{\text{H I}} \lesssim 3 \times 10^{56} \text{ s}^{-1}$ ), detection rates may be maximized through the selection of red, gigahertz peaked spectrum sources. The colour is indicative of a dusty environment, more conducive to the presence of cold neutral gas, and the turnover frequency indicative of an increased covering factor, due to the compact radio emission. If the projected extent of the radio emission is dictated by the orientation of the jets along the sightline, this presents an issue for the model where the absorption occurs in the obscuring torus invoked by unified schemes of active galactic nuclei (e.g. Jaffe & McNamara 1994; Conway & Blanco 1995; Morganti et al. 2001; Pihlström, Conway & Vermeulen 2003; Gupta et al. 2006; Gupta & Saikia 2006), since this is less likely to intercept the sightline in compact objects (e.g. Gupta et al. 2006). However, as well as occurring in outflowing gas (e.g. Morganti et al. 2003, 2005a, 2007; Vermeulen et al. 2003; Morganti, Tadhunter & Oosterloo 2005b; Allison et al. 2016), the bulk of the absorption is believed to occur in the large-scale galactic disc, rather than the sub-pc torus,<sup>13</sup> which is expected to have a random orientation to the radio jets (Curran & Whiting 2010 and references therein).

## 5 CONCLUSIONS

We have undertaken a survey for associated H I 21-cm and OH 18-cm absorption in six  $z \approx 0.2$ – $0.4$  radio sources with the GMRT. Despite selecting targets that have ionizing photon rates lower than the critical value, above which the neutral gas is believed to be completely ionized, and reaching the sensitivities of previous detections, we do not detect either transition. Given that the OH absorption strength is expected to be  $\lesssim 10^{-4}$  times that of the H I 21-cm absorption (Curran et al. 2007), the non-detection of this radical is not surprising, especially in the knowledge that our targets are not as reddened as originally believed. Regarding the H I, by combining our results with previous searches, in addition to the ionizing photon rate (UV luminosity), we find correlations between the detection rate and

<sup>12</sup> For example, the relationship between  $Q_{\text{H I}}$  (the UV luminosity) and the  $B - K$  colour, where  $B$  will be anticorrelated with  $L_{\text{UV}}$ , particularly at  $z \gtrsim 3$ . Since the radio and UV luminosities are correlated (Curran & Whiting 2010), there will also be a relationship between  $Q_{\text{H I}}$  and the search sensitivity.

<sup>13</sup> Although this does make a small contribution (Curran et al. 2016c).

(i) the blue–near-infrared colour, which is indicative of the reddening being due to dust, the presence of which helps in maintaining a low gas spin temperature in the case of the detections;

(ii) the turnover frequency of the radio SED of the source, which is believed to be anticorrelated with the size of the radio emission (e.g. Fanti et al. 1990). This itself is anticorrelated with the line strength, due to lower coverage of the measured flux, and so lower optical depths are expected for extended sources (Curran et al. 2013c).

The detection rate is, of course, also subject to the search sensitivity with a maximum rate of  $\approx 30$  per cent for  $N_{\text{H I}} \lesssim 10^{18} (T_{\text{spin}}/f) \text{ cm}^{-2}$ . A maximum of  $\sim 50$  per cent may be expected, due to the orientation of the absorbing gas with respect to the continuum emission (Curran & Whiting 2010), and so in addition to maximizing the sample size, future surveys should also target optically faint/obscured, radio bright objects. Since this could preclude objects with measured optical redshifts, full band ‘spectral scans’ towards sufficiently bright radio sources may be the way forward (e.g. Allison et al. 2015; Curran et al. 2005b).

We confirm that associated H I 21-cm absorption has never been detected above the theoretical value of  $Q_{\text{H I}} \approx 3 \times 10^{56} \text{ s}^{-1}$  (Curran & Whiting 2012). From the highest reliable detected value of  $Q_{\text{H I}} = 1.7 \times 10^{56} \text{ s}^{-1}$  (a monochromatic  $\lambda = 912 \text{ \AA}$  luminosity of  $L_{\text{UV}} \approx 2 \times 10^{23} \text{ W Hz}^{-1}$ ), the binomial probability of the observed 0 detections out of 29 searched above this value is just  $4.80 \times 10^{-7}$ , which is significant at  $5.03\sigma$ . Adding, the forthcoming results of Grasha et al. (2017), the probability of the 0 detections out of 63 searches is just  $2.58 \times 10^{-11}$ , which is significant at  $6.67\sigma$ . Hence, the case for a critical ultraviolet luminosity is strengthened. Given the unbiased (i.e. inhomogeneous) nature of the whole sample, this  $L_{\text{UV}} \sim 10^{23} \text{ W Hz}^{-1}$  ( $Q_{\text{H I}} \approx 3 \times 10^{56} \text{ s}^{-1}$ ) appears to be universal and the selection of sources of known optical redshift leads to the selection of higher luminosities at high redshift (see Morganti et al. 2015), where all of the gas is ionized. That is, even the Square Kilometre Array (SKA) will not detect H I 21-cm absorption in these. This leads us to reiterate that spectral scans towards optically faint objects is the best strategy and this is where the SKA will excel due to its wide instantaneous bandwidth.

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