

Long-term evolution of dim isolated neutron stars

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ABSTRACT

The X-ray dim isolated neutron stars (XDINs) have periods in the same range as the anomalous X-ray pulsars (AXPs) and the soft gamma-ray repeaters (SGRs). We apply the fallback disc model, which explains the period clustering and other properties of AXP/SGRs, to the six XDINs with measured periods and period derivatives. Present properties of XDINs are obtained in evolutionary scenarios with surface dipole magnetic fields $B_0 \sim 10^{12}$ G. The XDINs have gone through an accretion epoch with rapid spin-down earlier, and have emerged in their current state, with the X-ray luminosity provided by neutron star cooling and no longer by accretion. Our results indicate that the known XDINs are not likely to be active radio pulsars, as the low B_0 , together with their long periods place these sources clearly below the ‘death valley’.

Key words: accretion, accretion discs – stars: neutron – pulsars: individual (AXPs) – X-rays: bursts.

1 INTRODUCTION

What are the physical conditions leading to formation of different populations of young, isolated (single) neutron stars, namely anomalous X-ray pulsars (AXPs), soft gamma-ray repeaters (SGRs), dim isolated neutron stars (XDINs), central compact objects (CCOs) and rotating radio transients (RRATs)? All these systems involve a single, young or middle-aged neutron star (see Mereghetti 2011a, for a recent review of isolated neutron stars). Despite striking differences, like energetic soft gamma bursts peculiar to SGRs and AXPs, there are also striking similarities, like the periods of AXP/SGRs and XDINs being clustered in the same range (2–12 s). Considering the estimated birth rates of these objects together with the galactic supernova rate, it is likely that there are evolutionary connections between some of these populations (Keane & Kramer 2008). Recent efforts concentrate on the unification of the long-term X-ray luminosity and the rotational evolution of these neutron star systems in a single picture. With the assumption that these sources are evolving in vacuum with dipole torques, their surface dipole fields, B , inferred from the dipole torque formula range from $\sim 10^{10}$ G for CCOs to more than 10^{15} G for AXP/SGRs. In this picture, the diversity of evolutionary paths is attributed to the differences in the initial dipole and the crustal toroidal fields of the sources (Kaspi 2010). Sources with magnetar dipole fields ($B > 10^{14}$ G) are posited to go through a rapid field decay, which is required for the model to explain the recently discovered so-called ‘low-B magnetars’ (Turolla et al. 2011, Rea et al. 2012). Evolution of the X-ray luminosities of these sources, which are much higher than the rotational powers, are suggested to be governed by

the field decay history of the neutron stars depending on the initial crustal and dipole field strengths (Viganò et al. 2013). In this model, the apparently missing link between CCOs and other classes was suggested to be due to field burial to the crust by the accretion of the supernova matter in the early phase of evolution (Viganò & Pons 2012). The time-scale for the subsequent reemergence and the growth of the field to its original strength is estimated to be 10^3 – 10^5 yr depending on the initial conditions by the same authors. The dipole field’s decay is supposed to proceed after the initial burial and re-emergence. If there are fallback discs around these systems the picture is rather different. It was proposed by Alpar (2001) that the source properties, in particular the period clustering, could be explained by the presence (or absence) and the properties of fallback discs around these systems. Fallback disc models do not involve burial, re-emergence or decay of the dipole field, which is taken to be constant at its initial value.

Neutron stars with fallback discs and conventional dipole magnetic fields ($B = 10^{12}$ – 10^{13} G) could evolve into the X-ray luminosity, period and period derivative range of AXP/SGRs on time-scales of 10^3 to a few 10^5 yr (Ertan et al. 2009). The infrared (IR) and optical emission properties of these sources in the quiescent state are consistent with the emission from an irradiated active disc (Ertan & Çalışkan 2006, Ertan et al. 2007). The model fits to optical and IR data also constrain the dipole field strength to below $\sim 10^{13}$ G on the surface of the star (Ertan et al. 2007). Recently discovered ‘low-B magnetars’ SGR 0418+5729 (Rea et al. 2013) and Swift J1822.3–1606 (Scholz et al. 2012) clearly showed that the SGR bursts do not require magnetar strength dipole fields. This suggests that the SGR bursts could be powered by the quadrupole or higher multipole fields which are localized close to the surface and do not affect the rotation history of the star. The inner disc applies torque

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on the star through interaction with the large-scale dipole component of the magnetic field. The X-ray luminosity and the rotational properties of the ‘low-B magnetars’ and the so-called high-B radio pulsar PSR J1734–3333, including its anomalous braking index $n = 0.9 \pm 0.2$ (Espinoza et al. 2011), can also be reproduced by the same model and with $B = 1\text{--}2 \times 10^{12}$ G. As an independent and complementary support for this model, recent analyses show that the high-energy spectra of AXP/SGRs could be produced in their accretion columns consistently with the accretion rates of these sources inferred from the X-ray luminosities (Trümper et al. 2010, 2013; Kylafis, Trümper & Ertan 2014).

These results obtained from the application of the same model with similar basic parameters to sources with rather different observational properties encourages us to investigate the evolution and possible evolutionary links between AXP/SGRs (see Mereghetti 2008, 2011b for recent reviews) and other young neutron star populations. This will help us understand the differences in the initial conditions that lead to emergence of different classes of neutron stars. In particular, understanding whether there is a relation between the disc masses and the dipole field strengths requires detailed investigations of AXP/SGRs and other young neutron star populations. In the present work, we concentrate on the dim isolated neutron stars (XDINs).

At present, there are seven known dim isolated neutron stars (Haberl 2007; Turolla 2009). All these sources lie within a distance of ~ 400 pc. Since the solar neighbourhood is part of the Gould Belt, a ring of young stellar systems, the relatively high rate of supernova events in the Gould Belt should be taken into account in estimating the birth rate of XDINs. In the vicinity of the Sun, about two thirds of the neutron stars are born in the Gould Belt, while the remaining fraction belongs to the Galactic disc (Grenier 2000). Assuming the ages of XDINs are ~ 1 Myr, simple statistical calculations give a galactic birth rate of ~ 1 century $^{-1}$ (Popov, Turolla & Possenti 2006). The thermal X-ray luminosities of XDINs are in the $10^{31}\text{--}10^{32}$ erg s $^{-1}$ range. The ages corresponding to these luminosities on the theoretical cooling curves are a few 10^5 yr. Kinematic ages of four XDINs estimated from the space velocities and likely birth places are, with large uncertainties, in the range of $\sim 0.1\text{--}1$ Myr (Motch et al. 2009; Tetzlaff et al. 2010; Mignani et al. 2013). These kinematic ages and the cooling ages estimated from the X-ray luminosities of XDINs, imply that the characteristic ages, estimated assuming isolated pulsar spin-down, $P/2\dot{P} \sim 1\text{--}4$ Myr, could be significantly greater than the true ages of these sources.

XDINs do not show pulsed radio emission. Weak radio emission from two sources were reported (Malofeev, Malov & Teplykh 2007), but not confirmed yet (Turolla 2009; Haberl 2007). Strength of the surface dipole fields of XDINs inferred from the dipole torque formula are $\sim 10^{13}\text{--}10^{14}$ G. If these sources are indeed evolving in vacuum with dipole torques, many of them should be active as radio pulsars. Non-detection of radio pulses from these sources might be due to narrow beaming angles of long-period system. Nevertheless it is not clear why a large population of galactic XDINs with rotational properties of the known sources do not show up as radio pulsars. Alternatively, these sources may not have sufficiently strong dipole fields for radio emission. Our results support the latter possibility. In the present work, we try to explain the long-term evolution of XDINs in the frame of the fallback disc model as applied earlier to AXP/SGRs and PSR J1734–3333. We also discuss the radio properties of XDINs based on the predictions of our evolutionary model. We briefly describe the model in Section 2. The results of the simulations are given in Section 3. We summarize our conclusions in Section 4.

2 MODEL

We use the code developed to investigate the long-term evolution of AXPs and SGRs (see Ertan & Erkut 2008; Ertan et al. 2009; Alpar, Ertan & Çalıřkan 2011; Çalıřkan et al. 2013, for details and applications). We examine the period, the period derivative and the total X-ray luminosity evolution of the model sources, tracing the initial conditions, namely the initial period, P_0 , strength of the magnetic dipole field on the pole of the star, B_0 , and the initial disc mass, M_d . In addition to these initial parameters, the evolution is also affected by the irradiation efficiency C and the minimum critical temperature, T_p , for the disc to be viscously active. The magneto-rotational instability (Balbus & Hawley 1991) which generates the turbulent viscosity needed for the disc to transport mass and angular momentum will not work at temperatures below T_p , because the ionization fraction becomes too small.

The Alfvén radius of the disc could be written as

$$r_A \simeq (GM)^{-1/7} \mu^{4/7} \dot{M}_{\text{in}}^{-2/7} \quad (1)$$

(Lamb, Pethick & Pines 1973; Davidson & Ostriker 1973), where G is the gravitational constant, M and μ are the mass and magnetic dipole moment of the neutron star, and \dot{M}_{in} is the mass-flow rate arriving at the inner disc radius, r_{in} . When r_A is less than the light cylinder radius, $r_{\text{LC}} = c/\Omega_*$, we take $r_{\text{in}} = r_A$. Accretion will take place in this regime. Typically for the sources we consider, in the accretion phase $r_{\text{in}} < r_{\text{LC}}$ and the inner disc radius is greater than the co-rotation radius, $r_{\text{in}} > r_{\text{co}} = (GM/\Omega_*^2)^{1/3}$. The star is in the propeller phase, spinning down under the disc torques while accreting. Over the long-term evolution, r_{LC} increases with decreasing angular frequency of the neutron star, Ω_* , while r_A increases with decreasing \dot{M}_{in} . When r_A calculated by equation (1) is found to be greater than the current value of r_{LC} , we set $r_{\text{in}} = r_{\text{LC}}$. In this ‘tracking phase’, we assume that there is no accretion on to the neutron star.

We solve the diffusion equation for an extended disc with an initial surface-density distribution in the power-law form, $\Sigma(r) \propto r^{-3/4}$, which is the characteristic surface-density profile for steady thin accretion discs (see e.g. Frank, King & Raine 2002). At a given time, the temperature in the disc decreases with increasing radial distance r from the centre. At a given r , the temperature decreases with time because of fallback disc evolution and decreasing irradiation strength. In the calculations, for numerical reasons, we start with $r_{\text{out}} = 5 \times 10^{14}$ cm. After the first time step, the radius at which the disc temperature is as low as the minimum temperature T_p for the disc to be active is identified as the dynamical outer radius, r_{out} , of the viscously active disc. As the temperatures, starting from the outermost disc, decrease below T_p , r_{out} propagates inward. The X-ray irradiation of the disc is important in the evolution of the neutron star, since it prolongs the lifetime of the disc by delaying the inward propagation of r_{out} . X-ray irradiation flux can be written as $F_{\text{irr}} = C\dot{M}c^2/(4\pi r^2)$ (Shakura & Sunyaev 1973), where c is the speed of light. From the results of our earlier work on the IR and X-ray emission of persistent AXP/SGRs (Ertan & Çalıřkan 2006), we found the value of the irradiation efficiency C to be in the range of $\sim 1\text{--}7 \times 10^{-4}$ which is similar to the range of C estimated for low-mass X-ray binaries (Dubus et al. 1999). There is a degeneracy between the parameters C and T_p . With two extreme values of C , similar evolutionary curves can be obtained by changing T_p only by a factor of ~ 1.6 . The range of C obtained in our earlier work constrains T_p to below ~ 200 K (Ertan et al. 2009), consistent with the results of Inutsuka & Sano (2005).

Accretion on to the surface of the neutron star is the dominant source of the X-ray luminosity. The accretion rate, \dot{M} , is related to the X-ray luminosity through $L_x = GM\dot{M}/R$, where R is the radius of the neutron star. In the fallback disc model of AXP/SGRs, the sources are in the propeller regime and a fraction of the matter arriving at the inner disc radius is accreted on to the star (Chatterjee, Hernquist & Narayan 2000; Alpar 2001). This can be written as $\dot{M} = \eta\dot{M}_{\text{in}}$, where $\eta \leq 1$. For simplicity, we take $\eta = 1$ in all our calculations. Similar results can be obtained with $\eta < 1$. Depending on the initial conditions, the sources can enter the accretion phase at different epochs of the evolution. Some sources may never enter the accretion phase and probably continue their evolution as radio pulsars until their rotational power is no longer sufficient to produce pulsed radio emission. Other sources, with different initial conditions, may evolve through an initial pulsar phase into a subsequent accretion epoch. When accretion is not possible, that is, when the inner disc cannot penetrate the light cylinder, the X-ray emission is mainly due to intrinsic cooling of the star. For the cooling luminosity, L_{cool} , we use the theoretical cooling curves calculated for the neutron stars with conventional magnetic dipole fields (Page 2009). In the luminosity calculation, in addition to L_{cool} we also include the intrinsic dissipative heating of the neutron star under the dipole and disc torques acting on the star (Alpar 2007).

We use the α prescription for the kinematic viscosity $\nu = \alpha c_s h$, where c_s and h are the local sound speed and the pressure scale-height of the disc, respectively (Shakura & Sunyaev 1973). In the long-term evolution of the disc, the mass-flow rate from the outer to the inner disc is determined by the viscosities in the cold outer disc. Following the results of the detailed work on the X-ray enhancement light curves of AXP/SGRs (Çalışkan & Ertan 2012) we take $\alpha = 0.045$.

The disc could remain stable (i.e. not blown away by radiation pressure) when $r_{\text{LC}} < r_{\text{in}}$ for a certain range of r_{in} values depending on the angle θ between the magnetic dipole and the rotation axes of the neutron star (Ekşi & Alpar 2005). The maximum r_{in} for the disc to remain stable is a few r_{LC} for large θ , while the disc is stable for all r_{in} values if the axes are aligned (Ekşi & Alpar 2005). Even if the disc remains stable, it cannot apply an efficient torque on the star if the inner disc loses contact with the closed field lines. An alternative and stronger possibility in this regime is that the inner disc matter cannot be propelled efficiently from the system and piles up at the inner disc. In this case, because of the increasing surface-density gradient at the inner disc, r_{in} approaches to r_{LC} and a steady state is reached with an inner disc that remains in contact with the closed dipole field lines. The system then follows a tracking phase with $r_{\text{in}} \sim r_{\text{LC}}$ (Ertan et al. 2009; Alpar et al. 2011; Çalışkan et al. 2013). As in our earlier work, we adopt that the inner disc follows this tracking phase when the inner disc cannot penetrate the light cylinder.

In the accretion regime, the disc spin-down torque acting on the star can be written as $N = \dot{M}(GMr_{\text{in}})^{1/2} F(\omega_*)$ where the dimensionless torque $F(\omega_*)$ is a function of the fastness parameter $\omega_* = \Omega_*/\Omega_K(r_{\text{A}})$, where $\Omega_K(r_{\text{A}})$ is the angular frequency of the disc at $r = r_{\text{A}}$. From our earlier analysis, we found that the dimensionless torque in the form $F(\omega_*) = \beta(1 - \omega_*^p)$ can produce the period evolution of AXP/SGRs with $\beta = 0.5$ and $p = 2$ (Ertan & Erkut 2008; Ertan et al. 2009). Substituting $F(\omega_*)$ in the torque equation, we have

$$N = \frac{1}{2} \dot{M}_{\text{in}} (GMr_{\text{in}})^{1/2} (1 - \omega_*^2) = I \dot{\Omega}_*, \quad (2)$$

where I is the moment of inertia of the neutron star. When $r_{\text{A}} < r_{\text{LC}}$, we calculate N with $r_{\text{in}} = r_{\text{A}}$ using equation (1). When r_{A} is found to be greater than r_{LC} , we take $r_{\text{in}} = r_{\text{LC}}$.

In the spin-down with accretion phase ($r_{\text{co}} < r_{\text{in}} < r_{\text{LC}}$), substituting equation (1) in equation (2), the period derivative, \dot{P} , of the neutron star is found to be independent of both \dot{M} and P when ω_* is sufficiently beyond unity. This is probably the case for most of the persistent AXPs and SGRs currently in the accretion phase except for a few sources which are likely to be close to rotational equilibrium.

The ‘low-B magnetars’ SGR 0418+5729 and Swift J1822.3 – 1606, which seem to have completed the accretion epoch, are evolving in the tracking phase with $r_{\text{in}} \simeq r_{\text{LC}}$ (Alpar et al. 2011; Benli et al. 2013). In this phase, the disc torque is proportional to \dot{M}_{in} . The ratio of the ram pressure to magnetic pressure at r_{LC} , and the disc torque decrease with time. By the time the magnetic dipole torque dominates the disc torque, the X-ray cooling luminosity of the source is already below the detection limits in most cases. In the tracking phase following the accretion phase, the pulsed radio emission rate depends on the dipole field strength and the period of the source (see Sections 3 and 4). We assume that magnetic field decay is negligible within the observable time-scale of the AXP/SGRs and XDINs ($\tau \sim 10^6$ yr). Sources with appropriate initial conditions could start their evolution in the tracking phase. These are likely to be observed as radio pulsars until they cross the radio pulsar ‘death line’.

3 RESULTS

Out of seven XDINs, six sources have measured period and period derivatives.¹ For comparison with our model results, we have converted the reported X-ray fluxes of the six XDINs with known P and \dot{P} into the unabsorbed fluxes and luminosities using the distances, N_{H} values and blackbody temperatures given in the corresponding papers. There are large uncertainties in the distances of some XDINs which are reflected in the uncertainties of the X-ray luminosities. Corrections to luminosities could require modification of our model parameters reported here.

For RX J0420.0–5022, $P = 3.453$ s (Haberl et al. 2004) and $\dot{P} = 2.8 \pm 0.3 \times 10^{-14}$ s s⁻¹ (Kaplan & van Kerkwijk 2011). The distance of this source was estimated as ~ 345 pc by Posselt et al. (2007). The 0.1–2.4 keV observed flux was reported as $\sim 5 \times 10^{-13}$ erg s⁻¹ cm⁻² (Haberl et al. 2004). Using the properties of their spectral fit, we calculate the bolometric luminosity as $L_x \simeq 2.6 \times 10^{31}$ erg s⁻¹.

The period of RX J1308.6+2127 (RBS 1223) is $P = 10.31$ s (Haberl 2004). The most recent timing analysis gives $\dot{P} = 1.120(3) \times 10^{-13}$ s s⁻¹ (Kaplan & van Kerkwijk 2005). The distance is estimated to be 76–380 pc by Schwope et al. (2005) and 380_{-30}^{+15} pc by Hambaryan et al. (2011). Using $d = 380$ pc and the observed flux reported by Schwope, Schwarz & Greiner (1999), we obtained $L_x = 7.9 \times 10^{31}$ erg s⁻¹.

For RX J0806.4–4123, $P = 11.37$ s and $\dot{P} = 5.5(30) \times 10^{-14}$ s s⁻¹ (Kaplan & van Kerkwijk 2009a). The distance of this source was estimated as 235–250 pc by Posselt et al. (2007). Using

¹ When we were submitting this paper the detection of the period ($\simeq 3.39$ s) and a tentative (2σ) period derivative ($\sim 1.6 \times 10^{-12}$ s s⁻¹) was reported for the seventh source RX J1605.3+3249 (Pires et al. 2014). We intend to investigate the evolutionary possibilities of this source when the period derivative is confirmed.

Table 1. The ages and the disc parameters for the six XDINs. The ages are constrained by the theoretical cooling luminosity (Page 2009) and the estimated bolometric X-ray luminosities.

	B_0 (10^{12} G)	M_{disc} ($10^{-6} M_{\odot}$)	T_p (K)	C (10^{-4})	age (10^5 yr)
RX J0720.4–3125	1.1–1.3	0.8–12	106	1	1.45
RX J1856.5–3754	0.9–1.1	0.8–18	100	1	1.85
RX J2143.0+0654	1.0–1.2	1.0–12	100	1	1.9
RX J1308.6+2127	0.9–1.0	0.6–18	100	1.5	2.1
RX J0806.4–4123	0.8–0.9	0.5–18	100	2.3	3.1
RX J0420.0–5022	0.35–0.38	4.8–18	82	7	3.2

$d = 240$ pc and with the observed X-ray flux reported by Haberl et al. (2004), we find $L_x = 2.5 \times 10^{31}$ erg s $^{-1}$.

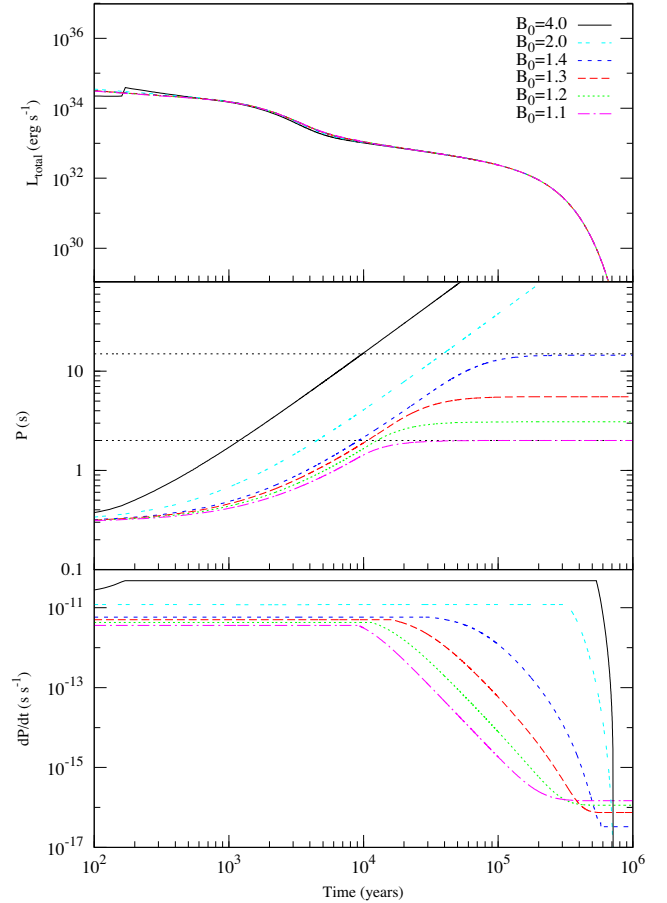
The period of RX J1856.5–3754 is 7.055 s (Tiengo & Mereghetti 2007) and $\dot{P} = 2.97(7) \times 10^{-14}$ s s $^{-1}$ (van Kerkwijk & Kaplan 2008). The distance is highly uncertain. Estimates range from 123^{+11}_{-15} pc (Walter et al. 2010) to 167^{+18}_{-15} pc (Kaplan, van Kerkwijk & Anderson 2007). Posselt et al. (2007) estimate the distance as 135 ± 25 pc. With $d = 135$ pc, L_x is found between $\sim 5 \times 10^{31}$ erg s $^{-1}$ and $\sim 1 \times 10^{32}$ erg s $^{-1}$ (Drake et al. 2002; Pons et al. 2002; Walter & Lattimer 2002; Burwitz et al. 2003). We take $L_x = 9.5 \times 10^{31}$ erg s $^{-1}$ obtained by Burwitz et al. (2003) with the lowest χ^2 .

For RX J2143.0+0654 (RBS 1774), the timing analysis by Kaplan & van Kerkwijk (2009b) gives $P = 9.428$ s and $\dot{P} = 4.1(18) \times 10^{-14}$ s s $^{-1}$. The upper limit to the distance was given as 300 pc by Posselt, Neuhäuser & Haberl (2009) and as 390–430 pc by Posselt et al. (2007). The unabsorbed X-ray fluxes reported by Kaplan & van Kerkwijk (2009b), Zane et al. (2005) and Rea et al. (2007) are 4.8, 6.1 and 5.6×10^{-12} erg s $^{-1}$ cm $^{-2}$, respectively. Using the flux reported by Rea et al. (2007), we find $L_x \simeq 1.1 \times 10^{32}$ erg s $^{-1}$ with $d = 400$ pc.

The period of RX J0720.4–3125 is $P = 8.39$ s (Haberl et al. 1997). The most recent analyses give $\dot{P} \sim 7 \times 10^{-14}$ s s $^{-1}$ (van Kerkwijk et al. 2007; Hohle et al. 2010). The distance was estimated as 235–270 pc by Posselt et al. (2007) and as 360^{+170}_{-90} pc by Kaplan et al. (2007). We use $d = 270$ pc in our calculations, since it agrees with both estimates. The X-ray flux was reported as $(9 \pm 2) \times 10^{-12}$ erg s $^{-1}$ cm $^{-2}$ by Kaplan et al. (2003), which corresponds to a bolometric luminosity of $\sim 1.6 \times 10^{32}$ erg s $^{-1}$.

In the present work, we have investigated the long-term evolution of these six XDINs (Table 1). The present day X-ray luminosities, periods and period derivatives of these sources can be produced simultaneously by the neutron star’s evolution with a fallback disc. For a neutron star with a magnetic dipole field $\sim 10^{12}$ G to evolve into the observed period range of XDINs (3–12 s) the system should pass through an accretion phase. Our results indicate that these sources slowed down to long periods during accretion epochs lasting from a few 10^4 yr to a few 10^5 yr. They completed their accretion phases in the past, and now they are in the tracking phase, still spinning down by disc torques, which are weaker than in the accretion epoch. At present, the X-ray luminosities of these sources are likely to be powered by the intrinsic cooling of the neutron star. If the inner disc had never penetrated into the light cylinder the disc torque could not have spun down a neutron star with a conventional $\sim 10^{12}$ G dipole field to XDIN periods within the cooling time-scale, $\tau_{\text{cool}} \sim 10^6$ yr (e.g. Page 2009).

We present illustrative model curves in Figs 1 and 2 to show the dependence of evolutionary paths on the disc mass, M_d , and the dipole field strength on the pole of the star, B_0 , respectively. Different phases of evolution can be followed from the \dot{P} curves given in the bottom panels of Figs 1 and 2. In the accretion phase

**Figure 1.** Total luminosity, period and period derivative evolution of the model sources with different magnetic dipole fields. For all model sources $M_d = 3 \times 10^{-6} M_{\odot}$. The magnitudes of the magnetic dipole field on the pole of the star, B_0 , are given in units of 10^{12} G in the top panel (see the text for a discussion).

($r_{\text{co}} < r_{\text{in}} < r_{\text{LC}}$), from equation (1), it is found that \dot{P} is proportional to B_0^2 , and independent of \dot{M}_{in} and P . That is, the evolutionary phases with constant \dot{P} correspond to accretion epochs. In this phase, for a given M_d , the sources with higher B_0 reach longer periods. After the accretion phase, the system enters the tracking phase in which $r_{\text{in}} \simeq r_{\text{LC}}$. The disc torque decreases with decreasing \dot{M}_{in} (the mass inflow rate arriving at $r_{\text{in}} \simeq r_{\text{LC}}$). Accretion on to the neutron star has stopped, $\dot{M} = 0$, and the X-ray luminosity is supplied by the intrinsic cooling of the neutron star. In Fig. 1, we also give illustrative model curves that could represent the evolution of AXP/SGRs (the two upper curves). These model sources with relatively high B_0 values cannot acquire the properties of XDINs. Our results

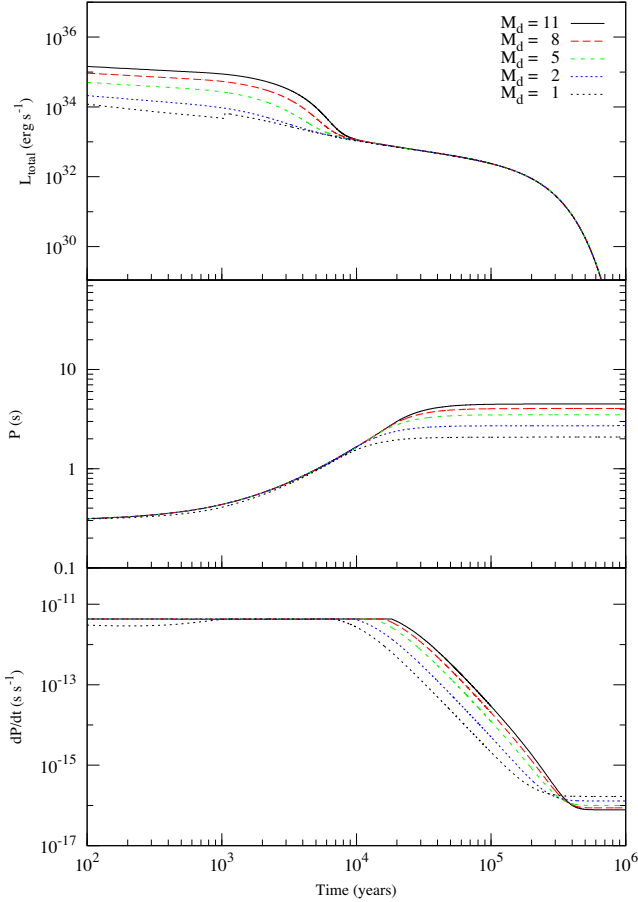


Figure 2. The luminosity, period and period derivative evolution of the model sources with different initial disc masses. The magnetic dipole field is the same ($B_0 = 1.2 \times 10^{12}$ G) for all sources. The disc masses are given in units of $10^{-6} M_{\odot}$ in the top panel.

indicate that XDINs could be distinguished from AXP/SGRs by having weaker B_0 fields.

In Fig. 2, we give the evolutionary curves for different disc masses keeping the initial dipole magnetic field constant at $B_0 \simeq 1.2 \times 10^{12}$ G. It is seen that systems with greater M_d have longer accretion phases. We obtain these illustrative model curves with $T_p \sim 200$ K and $C = 7 \times 10^{-4}$. Using $C = 1 \times 10^{-4}$, evolutionary curves similar to those in Fig. 2 can be obtained with $T_p \simeq 125$ K. Even with the highest reasonable value of C , the model curves obtained with $T_p > 300$ K do not lead to XDIN or AXP/SGR properties for any disc mass.

Comparing Figs 1 and 2, it is seen that the dipole magnetic field strength B_0 , rather than the disc mass, is the more effective initial condition for determining the long-term rotational evolution of the sources. In Fig. 1, M_d is the same for all model sources, and the B_0 values vary only by a factor of 4. We see that the model curves trace rotational properties of AXP/SGR and XDINs from the shortest to the longest observed P and \dot{P} values. For a given B_0 , different M_d values do not yield significantly different evolutionary paths. In Fig. 2, disc masses change by an order of magnitude, while the final periods at $t = 1$ Myr remain in the 2–5 s range. The \dot{P} curves of the sources with the same B_0 but different M_d are also very similar. In the accretion phase, the accretion power is the dominant source of the X-ray luminosity, and depends on the initial disc mass M_d .

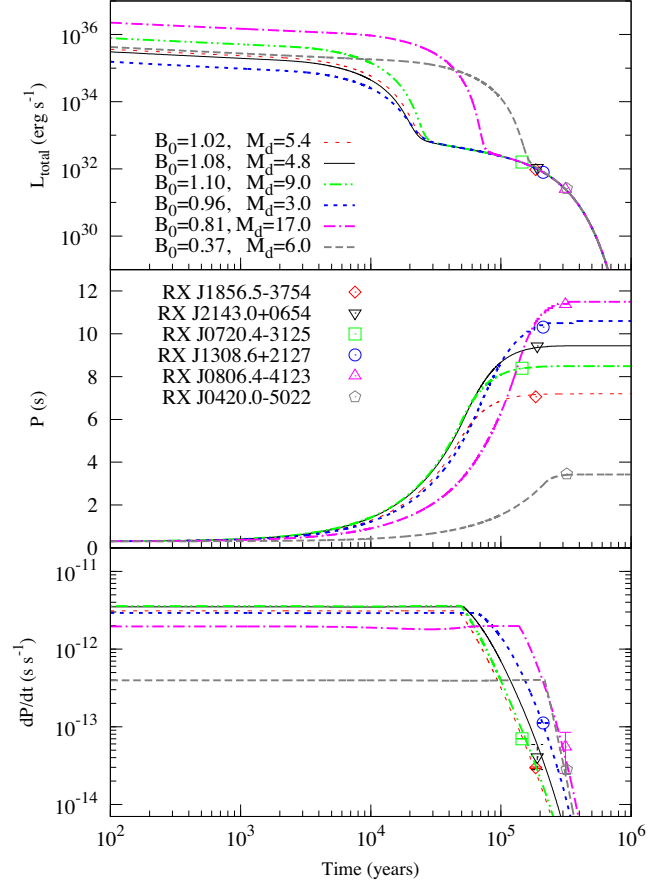


Figure 3. The model curves that can simultaneously produce the luminosity, period and period derivative of the six XDINs. For each source, we determined the ranges of M_d and B_0 that can produce the source properties (see Table 1). Here, we show illustrative model curves that can represent the long-term evolution of these XDINs. The values of B_0 and M_d used in the models are given in the top panel in units of 10^{12} G and $10^{-6} M_{\odot}$, respectively.

With more detailed analysis, we also tried to reproduce the individual properties of XDINs with known P and \dot{P} values. The evolutionary tracks of these sources are seen in Fig. 3. Since all these XDINs are currently powered by the intrinsic cooling luminosity of their neutron stars, observed luminosities and the theoretical cooling curves constrain the ages of the model sources. In all these calculations, we first take $C = 1 \times 10^{-4}$ and $T_p = 100$ K. In some cases, we obtained better fits with slightly different C and T_p values. The model curves given in Fig. 3 are obtained with T_p values that remain in a narrow range between 80 and 110 K. Assuming that AXP/SGRs and XDINs have similar disc compositions, we expect that they have similar critical temperatures. The irradiation efficiency might change with the accretion rate; it is expected to be similar for sources in the same accretion regime. We perform simulations tracing all possible values of M_d and B_0 with $P_0 = 300$ ms. For the model sources that can enter the accretion phase, the source properties at the end of this phase are not sensitive to P_0 (Ertan et al. 2009). For all XDINs, the B_0 values that can produce the reasonable evolutionary curves remain in the $\sim 0.3\text{--}1.3 \times 10^{12}$ G range. The ranges of model parameters that can produce the individual properties of the six XDINs are listed in Table 1.

Our results imply that all six XDINs completed the long-term accretion phase. They are not accreting matter from the disc at present,

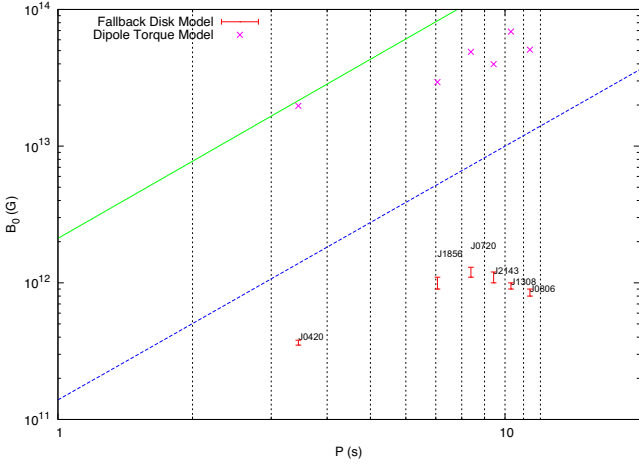


Figure 4. Magnetic dipole field strength on the pole of the neutron star versus period distribution for six XDINs. The two parallel lines (green and blue) show the upper and lower bounds of the radio pulsar death valley from Chen & Ruderman (1993). Crosses are the B_0 values inferred from the dipole torque formula. Vertical bars show the ranges of B_0 that can produce the properties of the sources in the fallback disc model (see Fig. 3). It is seen that all these sources remain below the lower boundary of the death valley (death line) indicating that these XDINs cannot be normal radio pulsars, if they are evolving with fallback discs.

while they are still being slowed down by the disc torques. Without accretion, these sources are free to emit pulsed radio emission. However, the dipole fields B_0 inferred in the fallback disc model are significantly weaker than those inferred from the dipole torque formula, so that at present the XDINs do not have sufficient voltages to sustain pulsed radio emission. Allowed ranges of B_0 values for each XDIN are given in Table 1 and plotted in Fig. 4. All these sources are indeed well below the lower border of the radio pulsar death valley in the extended galactic population, and are not expected to show radio pulses. Younger XDINs with shorter periods should be able to emit radio pulses, but they would still be evolving in the long-term accretion phase during which mass-flow on to the neutron star hinders the radio emission. Once the accretion epoch is over, XDINs have weak rotational power and are not likely to produce radio emission.

The inner disc conditions of XDINs in the efficient propeller phase could be rather different from those of persistent AXPs in the accretion phase, for which the boundary layer extends down to the co-rotation radius in a steady state with magnetic heating in the boundary layer negligible compared to viscous dissipation (Rappaport, Fregeau & Spruit 2004). By contrast, in the propeller phase, magnetic heating in the boundary layer could be a significant extra heating source in addition to viscous dissipation and irradiation. This additional heating in the propeller phase can supply the difference (~ 2 – 3 orders of magnitude) between the observed R band flux (Kaplan et al. 2011) and the model calculation, for an irradiated face on thin disc. A calculation of the magnetic heating is beyond the scope of the present work; a detailed study is in progress and will be presented in a separate paper.

4 DISCUSSION AND CONCLUSIONS

We have examined the long-term evolution of the six XDIN sources, with measured \dot{P} out of the current total seven XDINs. First, we find that the fallback disc model is successful in producing the observed individual properties of all these sources. XDIN properties

are obtained after the source has gone through a past accretion phase with spin-down, with the fallback disc protruding into the neutron star’s magnetosphere. At present, the fallback disc is tracking the light cylinder, with continuing spin-down under the disc torques, but no accretion on to the neutron star. The X-ray luminosities of XDINs are provided by cooling. Next, tracing the initial parameters, we have determined the allowed ranges of the dipole field strength on the pole of the star, B_0 , and the disc mass, M_d , for each source (Table 1). Our results constrain the B_0 values of the six XDINs into the ~ 0.3 – 1.3×10^{12} G range. Comparing this range of dipole field strength with earlier results obtained by Ertan et al. (2009) for AXP/SGRs, we surmise that XDINs tend to have dipole fields weaker than those of most known AXP/SGRs. This result should be confirmed through further detailed work on the persistent and transient AXP/SGRs.

Reasonable evolutionary model curves for XDINs can be obtained with a wide range of initial fallback disc masses, M_d , for all sources (Table 1). We could not test whether there is a correlation between M_d and B_0 , since our results do not constrain the disc masses. For the sources that do not accrete at present, like XDINs, it is not possible to constrain the initial mass of the disc in most cases. A possible M_d – B_0 correlation could be tested through further investigation of persistent AXP/SGRs which are powered by accretion on to the star. This analysis will help us understand the differences in the initial conditions of different young neutron star populations.

All known XDINs are close-by objects within ~ 400 pc of the Sun. Statistical analysis considering the properties of the Gould Belt gives a galactic birth rate of $\gtrsim 1$ century $^{-1}$ (Popov et al. 2006; see also Section 1). This raises a critical question for the models. Only the XDINs in the solar neighbourhood are observable in X-rays. The inferred large galactic population should be detectable in the radio band if they are active radio pulsars. If these sources evolve in vacuum without fallback discs, their dipole field on the pole of the star can be estimated as $B_0 \simeq 6.4 \times 10^{19} \sqrt{P \dot{P}}$ G. Observed P and \dot{P} values give $B_0 \sim 10^{13}$ – 10^{14} G. It is expected to see these sources as radio pulsars, if XDINs indeed slow down by magnetic dipole torques with these strong fields. Five out of seven XDIN sources, located within 0.4 kpc, have periods greater than 7 s. Among the currently known radio pulsars in the ATNF pulsar catalogue (Manchester et al. 2005), there is only a single radio pulsar with $P > 7$ s ($P = 7.7$ s) and $B > 10^{13}$ G which is located at a distance of 3.3 kpc. We are not sure whether this can be accounted for with the selection effects and even the most conservative beaming fraction formulas for the long-period systems, if XDINs and radio pulsars with large dipole magnetic fields are taken to be members of the same population. A detailed population synthesis or statistical analysis will not be attempted here.

This problem encountered in the dipole torque model is naturally resolved in the fallback disc model. During the initial long-term accretion phase radio emission is suppressed due to mass-flow on to the star. Is it possible to observe pulsed radio emission from these sources after the mass accretion terminates? The neutron stars that evolve like the six XDINs we studied here are not likely to produce beamed radio emission when accretion stops, because: (1) their dipole fields are much weaker than inferred from the dipole torque formula (Table 1), and (2) at the end of the accretion phase, they have already attained long periods. In Fig. 3, accretion phases correspond to the epochs with constant \dot{P} given in the lower panel. The period at the end of the accretion phase corresponds to a location in the B_0 – P plane which is already below the lower border of the death valley. Subsequent evolution moves the source even

further away from the death valley, to the present locations shown in Fig. 4.

To sum up: (1) the rotational properties and the X-ray luminosities of XDINs can be explained by the fallback disc model that was employed earlier to explain the general properties of AXP/SGRs (Fig. 1). The main disc parameters used in the present work for different XDIN sources are similar to each other and also to those used in the long-term evolution models of AXP/SGRs (Table 1). (2) The distinction between XDINs and AXP/SGRs seems to be that the former have weaker dipole magnetic fields. Finally, (3) the model can also account for non-detection of many radio pulsars evolving to the properties of known XDINs. A neutron star with a conventional dipole field must have passed through the long-term accretion phase to acquire the periods of known XDINs within the cooling time-scale. In the accretion phase, the radio emission is quenched by the mass flow on to the neutron star. After termination of the accretion phase, the sources no longer have sufficient rotational power for radio emission.

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